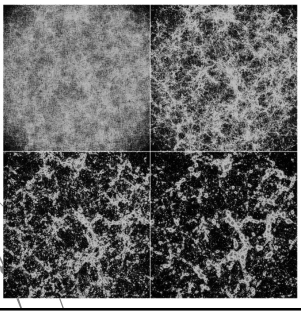
Cosmology AS7009, 2008
Lecture 10

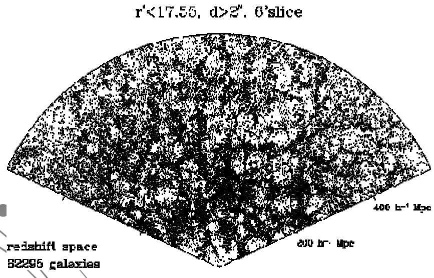


Outline

- Structure formation
  - Jeans length, Jeans mass
  - Structure formation with and without dark matter
  - Cold versus hot dark matter
  - Dissipation
  - The matter power spectrum
  - Baryon acoustic oscillations
- Reionization and high-z objects
  - When did reionization take place?
  - Caused by what?

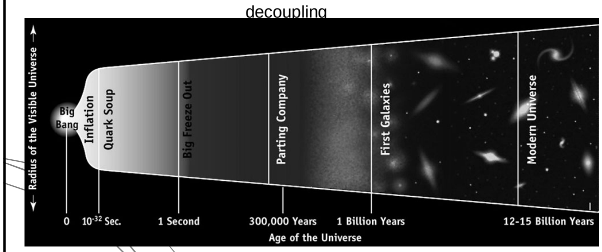
Covers chapter 12 in Ryden + extra stuff

Walls, Filaments, Voids



- Voids ~ 70 Mpc

Cosmic epochs



Matter-radiation decoupling

Likely seeds of galaxy formation: Quantum fluctuations expanded to macroscopic scales by inflation

Jeans length I

Which baryonic objects will collapse under the force of gravity?

- Two time scales:
  - Dynamical collapse time, t_{dyn}
  - Characteristic time scale for pressure build-up, t_{pre}
- $t_{pre} > t_{dyn} \rightarrow$ Object collapses
- $t_{pre} < t_{dyn} \rightarrow$ Hydrostatic equilibrium attained; collapse prevented

Jeans length II

Jeans length λ_J: Size of overdense regions for which $t_{pre} = t_{dyn} \rightarrow$

Regions of size > λ_J will collapse

Regions of size < λ_J will not

$$\lambda_J = \sqrt{\frac{\pi c^2}{G\bar{\varepsilon}}} c_s$$

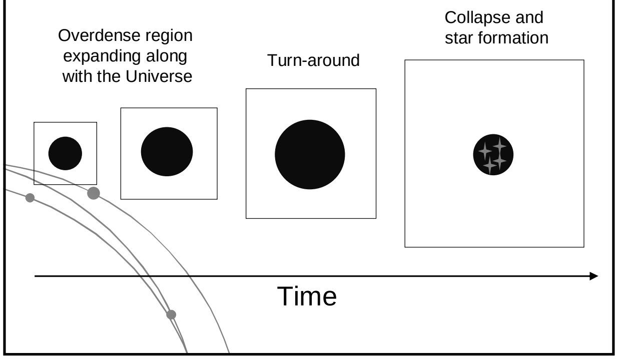
Sound speed in overdense region

Mean density of overdense region

Jeans mass

- Jeans mass M_J: Mass of baryons inside sphere of radius λ_J
  - $M > M_J \rightarrow$ Collapse
- Before decoupling: photon-baryon fluid with very high M_J (~ 10^{19} M_{solar})
- After decoupling: M_J drops to (~10^4-10^5 M_{solar}) in baryon fluid → Baryons lose pressure support

Collapse in an expanding Universe


Overdense region expanding along with the Universe

Turn-around

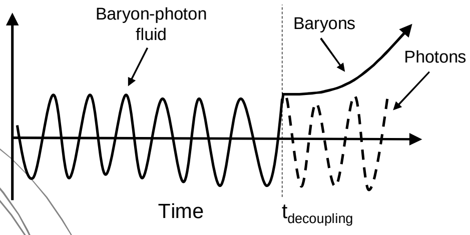
Collapse and star formation

Time

Structure Formation *Without* Non-Baryonic Dark Matter

- Density perturbations that will eventually form galaxies and galaxy clusters cannot start to grow until after decoupling (t ≈ 0.35 Myr)

$$\delta = \frac{\varepsilon - \bar{\varepsilon}}{\varepsilon}$$


Baryon-photon fluid

Baryons

Photons

Time

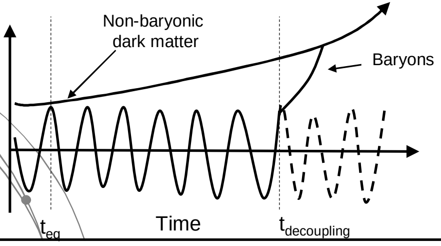
$t_{decoupling}$

Problem: Too slow structure formation – fails to explain the observed structures at high and low redshift

Structure Formation *With* Non-Baryonic Dark Matter

- Density perturbations will start to grow at the epoch of matter-radiation equality (t ≈ 0.047 Myr)
- Baryons will fall into the potential wells already produced by the dark matter

$$\delta = \frac{\varepsilon - \bar{\varepsilon}}{\varepsilon}$$


Non-baryonic dark matter

Baryons

t_{eq}

Time

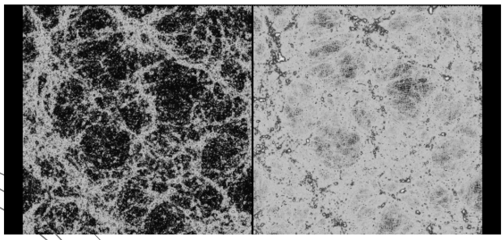
$t_{decoupling}$

Hot & cold dark matter I

- Hot dark matter (HDM): Relativistic velocities at decoupling
- Cold dark matter (CDM): Non-relativistic velocities at decoupling
- Warm dark matter (WDM): Intermediate velocities at decoupling

Velocities of the dark matter particles regulate how massive the first collapsing objects are

Hot & cold dark matter II


Cold dark matter

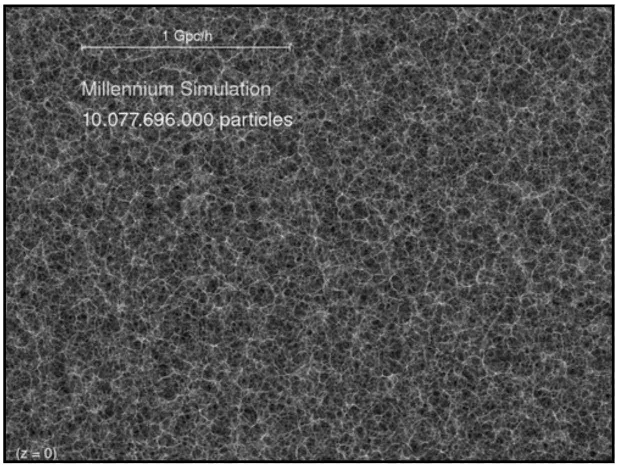
Cold + hot dark matter

HDM → Top-down structure formation

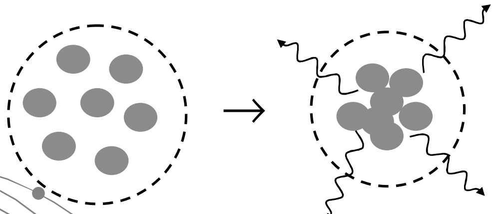
- Free-streaming wipes prevents growth of density perturbations on small scales
- Top-down: Big structures form first, small ones later
- Overdensitities of galaxy cluster mass collpase before the galaxies inside are formed
- Massive galaxies form before dwarf galaxies

CDM → Bottom-up structure formation

- Bottom-up = Small structures form first, big ones later
- Potential wells in non-baryonic CDM form before decoupling, into which baryons may fall after decoupling
- Small objects form first, galaxy clusters last (some are still collapsing)



Dissipation inside dark matter halos



Isolated dark matter halo with baryons (gas) inside

The baryons cool (dissipate) by radiating photons and sink towards the centre. The dark matter halo contracts slightly due to the changes in the gravitational field.

The Matter Power Spectrum

Most inflation models predict an adiabatic, power-law spectrum of Gaussian perturbations

Amplitude

Wave number

$$P \propto k^n$$

Corresponding potential fluctuations:

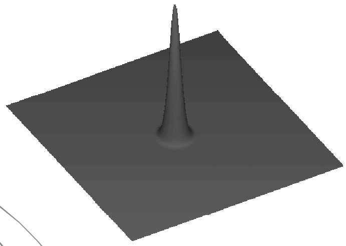
$$\delta\Phi \propto M^{(1-n)/6}$$

$n = 1 \Rightarrow$ Harrison - Zeldovich spectrum (scale - free)

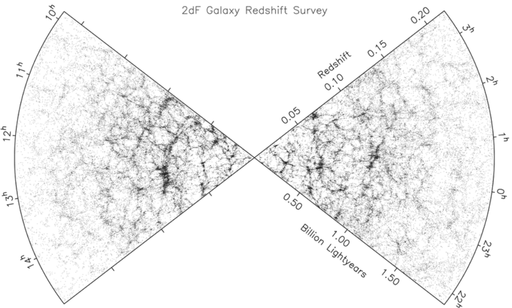
Baryon Acoustic Oscillations

- Overdensities (in baryons and dark matter), eject spherical sound waves
- Sound speed ~0.5 c
- Photons decouple → Sound speed drops
- Wave stalls at R~150 Mpc
- This overdensity of gas acts as seed for galaxy formation and can be detected in large galaxy surveys
- The 150 Mpc radius serves as a standard ruler

Baryon Acoustic Oscillations II

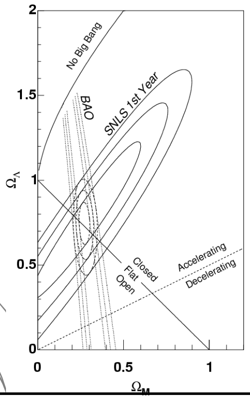


Baryon Acoustic Oscillations III

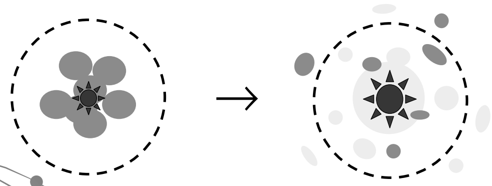


The 150 Mpc galaxy overdensity cannot be seen with the naked eye, but can be extracted using statistical analysis methods → Cosmological parameters can be derived

Baryon Acoustic Oscillations IV

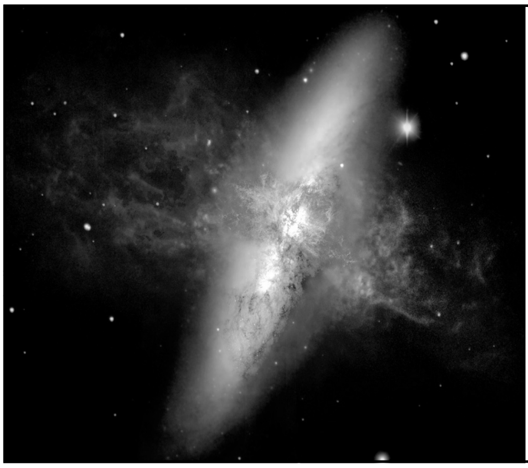


Star Formation and Feedback



Stars form
in the cooled gas

Feedback from star formation kicks in. Young, hot stars ionize part of the gas. Stellar winds and supernovae blow the gas outwards – some of it may leave the halo.



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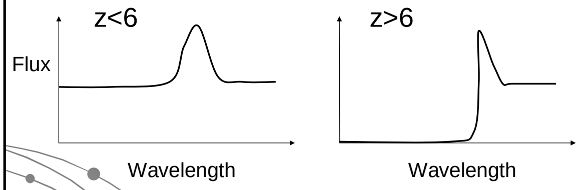
Reionization

- The Universe cooleds and becomes neutral at the epoch of recombination
- But most of the gas in the local Universe is ionized → Somewhere along the way the Universe must have experienced reionization
- Conjecture: Reionization is caused by the formation of astronomical objects (sources of Lyman continuum photons)
- The first astronomical light sources are expected to light up at around z = 30—15 (100—300 Myr after the Big Bang)

When did reionization take place?

Quasar Lyα spectra (The Gunn—Peterson test)

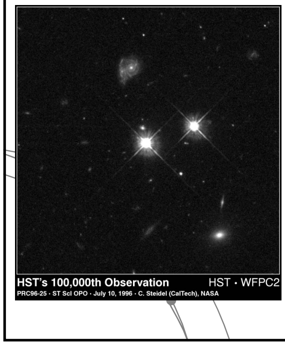


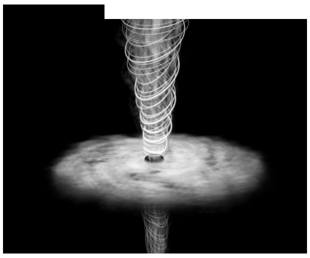
- Gunn-Peterson test → Reionization at $z \approx 6$
- But WMAP → Reionization at $z \approx 20$
- What is going on?!?
  - Prolonged reionization ($z=20$—6)?
  - Reionization happened twice?

What caused reionization?

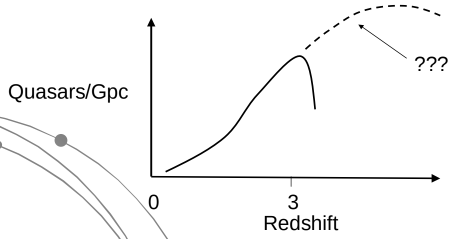
- Quasars? ⇦
- Starburst galaxies? ⇦
- Population III stars? ⇦
- Evaporating primordial black holes?
- Decaying dark matter?

Quasars



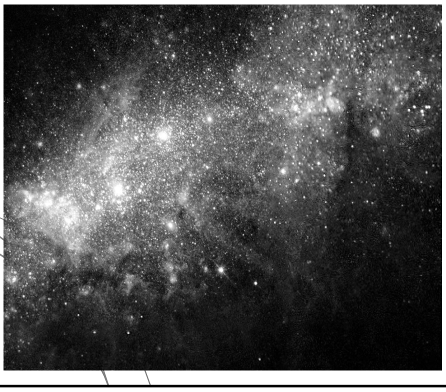


Quasars as Sources of Lyman Continuum Radiation at High Redshift



Probably too few…

Starburst Galaxies



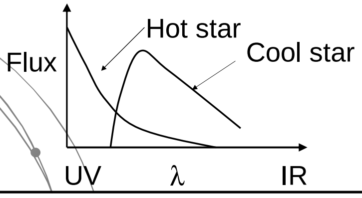
Why Do Starburst Galaxies Produce Lots of Lyman Continuum Photons?

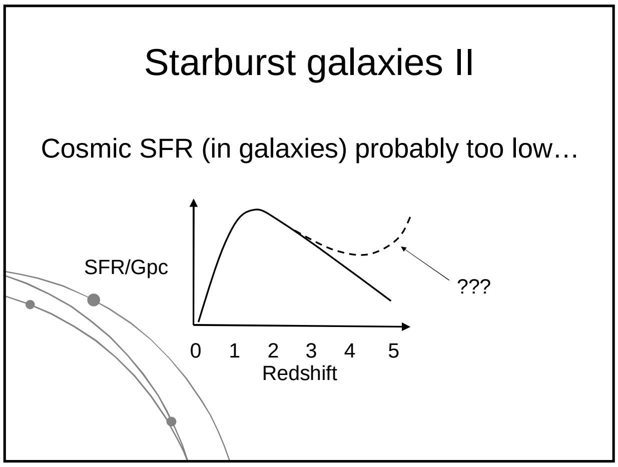
- Stars are born in the mass range ~ 0.08—120 solar masses
- The highest-mass stars have the shortest lifetimes (a few Myr)

→ Large numbers of high-mass stars are only found in galaxies that actively form stars

- High-mass stars are typically hotter than low-mass stars
- Hot stars emit more UV radiation (stars are almost black bodies)

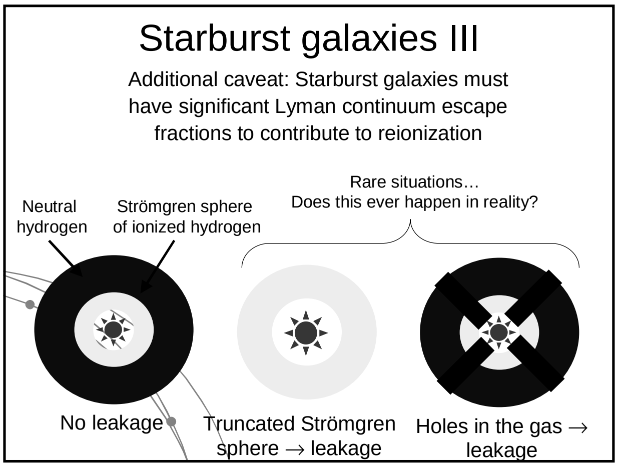
→ Starbursts emit more Lyman continuum radiation than galaxies with low star formation rates





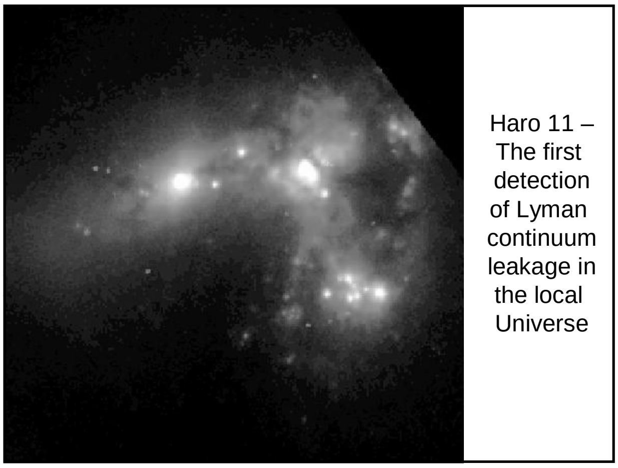
Starburst galaxies II

Cosmic SFR (in galaxies) probably too low…



Starburst galaxies III

Additional caveat: Starburst galaxies must have significant Lyman continuum escape fractions to contribute to reionization



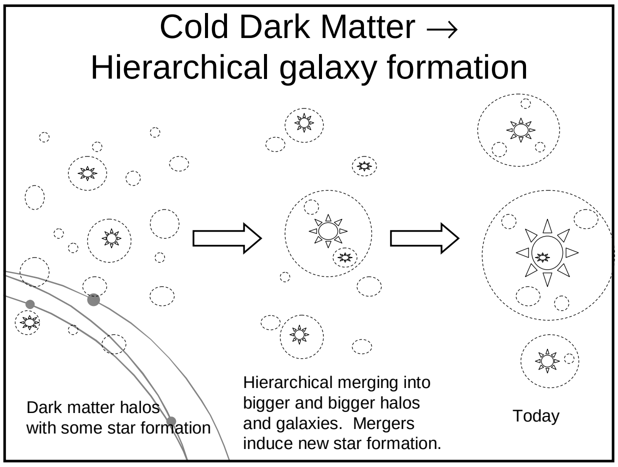
Haro 11 – The first detection of Lyman continuum leakage in the local Universe

Population III stars?

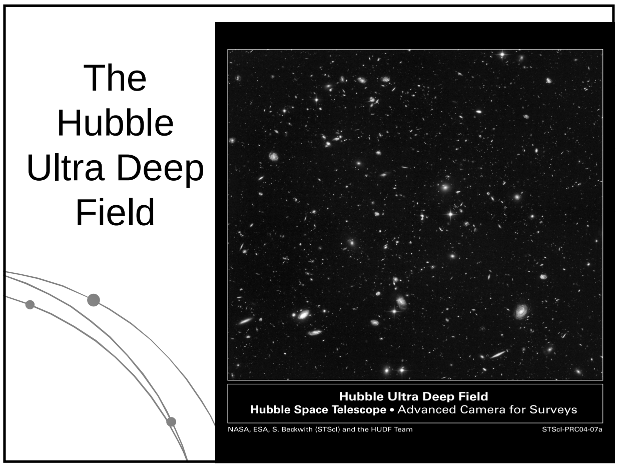
- Population I stars (young, metal-rich, disk)
- Population II stars (old, metal-poor, stellar halo)
- Population III stars (the oldest stars, metal-free)

Population III stars may have been very Massive (> 100 — 200 solar masses)

→ Short-lived, but produce a lot of Lyman continuum emission during their lifetimes

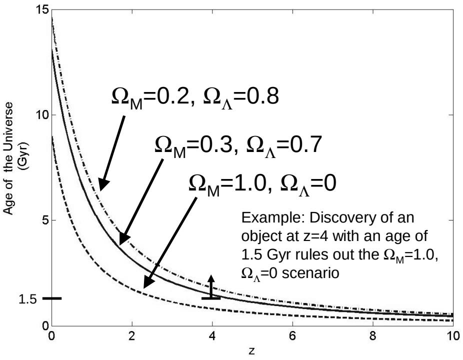


Cold Dark Matter → Hierarchical galaxy formation



The Hubble Ultra Deep Field

High-Redshift Objects as Probes of Cosmology: Ages



Ω_M=0.2, Ω_Λ=0.8

Ω_M=0.3, Ω_Λ=0.7

Ω_M=1.0, Ω_Λ=0

Example: Discovery of an object at z=4 with an age of 1.5 Gyr rules out the Ω_M=1.0, Ω_Λ=0 scenario

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Spectroscopic Age Determinations of Galaxies

Stars typically become redder when they grow older →
The shape of the spectrum of a galaxy (containing billions of stars) is indicative of the age