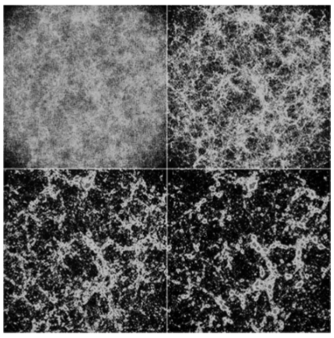
Cosmology 1FA209, 2015
Lecture 9: Structure formation

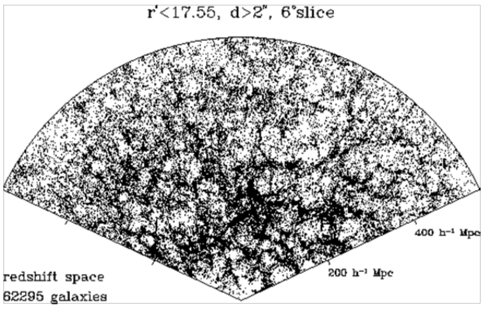


Outline

- Structure formation
  - Jeans length, Jeans mass
  - Structure formation with and without dark matter
  - Cold versus hot dark matter
  - Dissipation
  - The matter power spectrum
  - Baryon acoustic oscillations
- Reionization and high-z objects
  - What caused reionization?
  - The first stars and galaxies

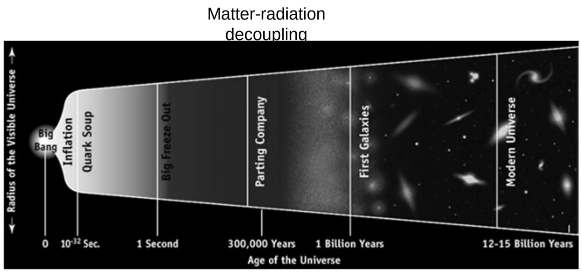
Covers chapter 12 in Ryden + extra stuff

Walls, Filaments, Voids



Voids ~ 70 Mpc

Cosmic epochs



Likely seeds of galaxy formation: Quantum fluctuations expanded to macroscopic scales by inflation

Jeans length I

Which baryonic objects will collapse under the force of gravity?

- Two time scales:
  - Dynamical collapse time, t_{dyn}
  - Characteristic time scale for pressure build-up, t_{pre}
- $t_{pre} > t_{dyn} \rightarrow$ Object collapses
- $t_{pre} < t_{dyn} \rightarrow$ Hydrostatic equilibrium attained; collapse prevented

Jeans length II

Jeans length λ_J: Size of overdense regions for which $t_{pre} = t_{dyn} \rightarrow$

Regions of size > λ_J will collapse
Regions of size < λ_J will not

$$\lambda_J = \sqrt{\frac{\pi c^2}{G\bar{\varepsilon}}}\, c_s$$

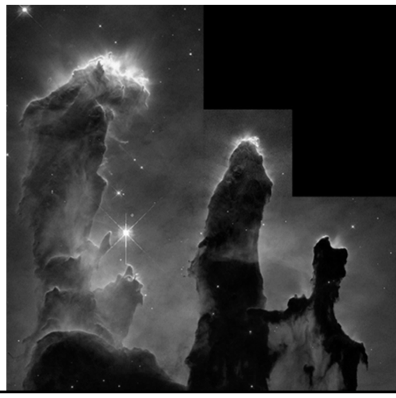
Sound speed in overdense region (c_s)

Mean density of overdense region ($\bar{\varepsilon}$)

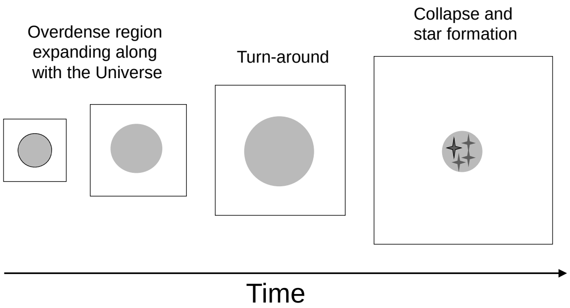
Jeans mass

- Jeans mass M_J: Mass of baryons inside sphere of radius λ_J
  - $M > M_J \rightarrow$ Collapse
- Before decoupling: photon-baryon fluid with very high M_J (~ 10^{19} M_{solar})
- After decoupling: M_J drops to (~10^4-10^5 M_{solar}) in baryon fluid $\rightarrow$ Baryons lose pressure support

Intermission: What is this?

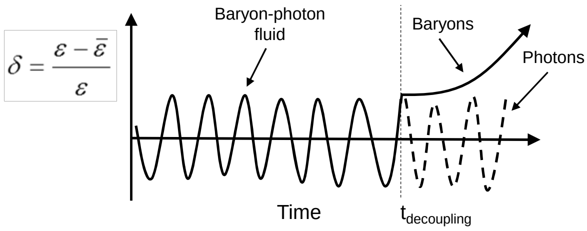


Collapse in an expanding Universe



Structure Formation *Without* Non-Baryonic Dark Matter

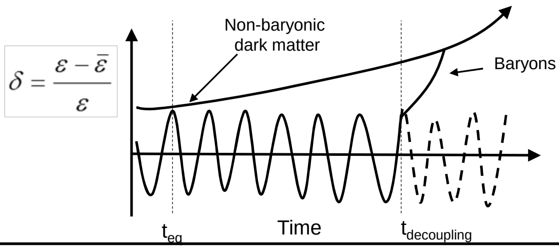
- Density perturbations that will eventually form galaxies and galaxy clusters cannot start to grow until after decoupling ($t \approx 0.35$ Myr)



Problem: Too slow structure formation – fails to explain the observed structures at high and low redshift

Structure Formation *With* Non-Baryonic Dark Matter

- Density perturbations will start to grow at the epoch of matter-radiation equality ($t \approx 0.047$ Myr)
- Baryons will fall into the potential wells already produced by the dark matter

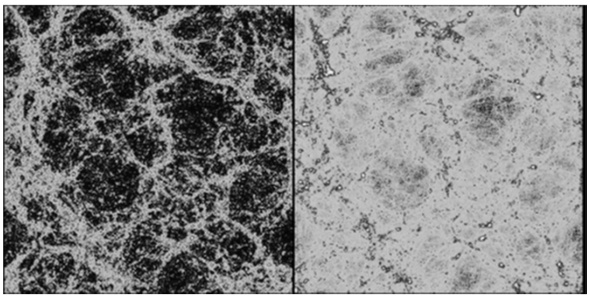


Hot & cold dark matter I

- Hot dark matter (HDM): Relativistic velocities at decoupling
- Cold dark matter (CDM): Non-relativistic velocities at decoupling
- Warm dark matter (WDM): Intermediate velocities at decoupling

Velocities of the dark matter particles regulate how massive the first collapsing objects are

Hot & cold dark matter II



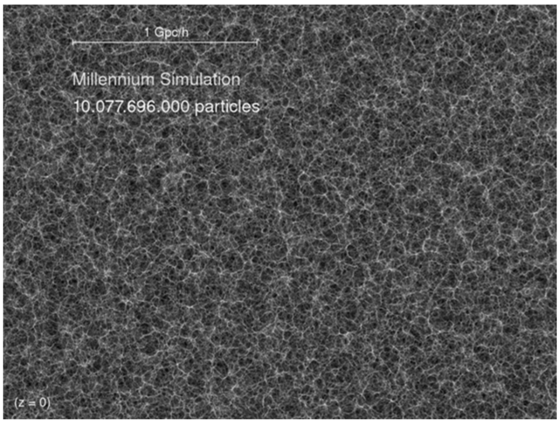
Cold dark matter   Cold + hot dark matter

HDM → Top-down structure formation

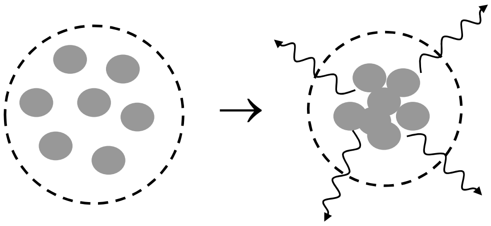
- Free-streaming wipes prevents growth of density perturbations on small scales
- Top-down: Big structures form first, small ones later
- Overdensitities of galaxy cluster mass collpase before the galaxies inside are formed
- Massive galaxies form before dwarf galaxies

CDM → Bottom-up structure formation

- Bottom-up = Small structures form first, big ones later
- Potential wells in non-baryonic CDM form before decoupling, into which baryons may fall after decoupling
- Small objects form first, galaxy clusters last (some are still collapsing)



Dissipation inside dark matter halos



Isolated dark matter halo with baryons (gas) inside

The baryons cool (dissipate) by radiating photons and sink towards the centre. The dark matter halo contracts slightly due to the changes in the gravitational field

Intermission: What is this?

The Matter Power Spectrum

Most inflation models predict an adiabatic, power-law spectrum of Gaussian perturbations

Amplitude → P; Wave number → k

$$P \propto k^n$$

Corresponding potential fluctuations :

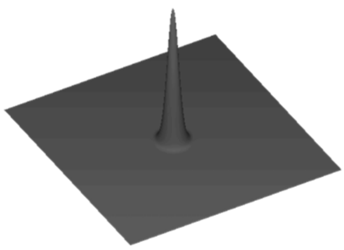
$$\delta\Phi \propto M^{(1-n)/6}$$

$n = 1 \Rightarrow$ Harrison - Zeldovich spectrum (scale - free)

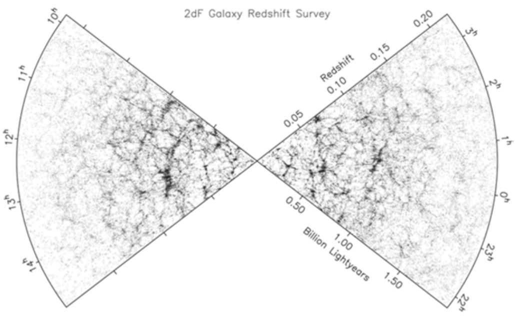
Baryon Acoustic Oscillations

- Overdensities (in baryons and dark matter), eject spherical sound waves
- Sound speed ~0.5 c
- Photons decouple → Sound speed drops
- Wave stalls at R~150 Mpc
- This overdensity of gas acts as seed for galaxy formation and can be detected in large galaxy surveys
- The 150 Mpc radius serves as a standard ruler

Baryon Acoustic Oscillations II

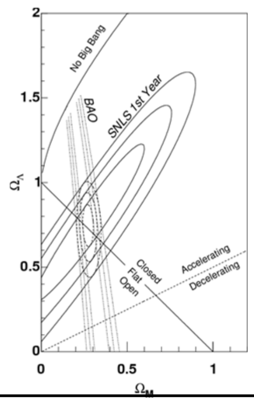


Baryon Acoustic Oscillations III

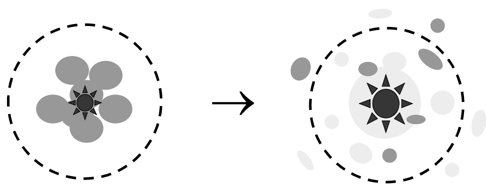


The 150 Mpc galaxy overdensity cannot be seen with the naked eye, but can be extracted using statistical analysis methods → Cosmological parameters can be derived

Baryon Acoustic Oscillations IV



Star Formation and Feedback



Stars form in the cooled gas

Feedback from star formation kicks in. Young, hot stars ionize part of the gas. Stellar winds and supernovae blow the gas outwards – some of it may leave the halo

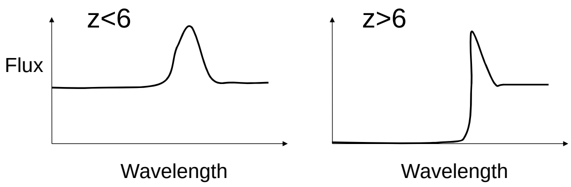
Intermission: What is this?

Reionization

- The Universe cooleds and becomes neutral at the epoch of recombination
- But most of the gas in the local Universe is ionized → Somewhere along the way the Universe must have experienced reionization
- Conjecture: Reionization is caused by the formation of astronomical objects (sources of Lyman continuum photons)
- The first astronomical light sources are expected to light up at around z = 30—15 (100—300 Myr after the Big Bang)

When did reionization take place?

Quasar Lyα spectra (The Gunn—Peterson test)

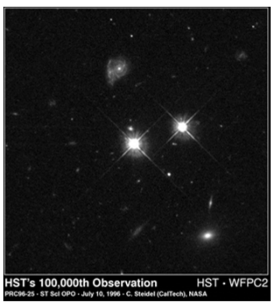


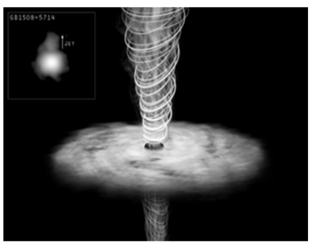
- Gunn-Peterson test → Reionization ended at $z \approx 6$
- CMBR analysis measures electron scattering optical depth τ → Reionization started at $z \approx 12$

What caused reionization?

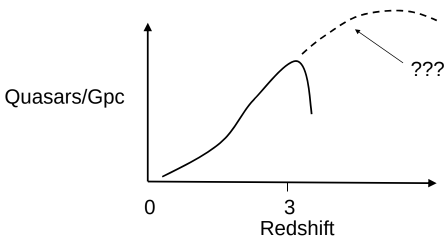
- Quasars?
- Starburst galaxies?
- Population III stars?
- Evaporating primordial black holes?
- Decaying dark matter?

Quasars



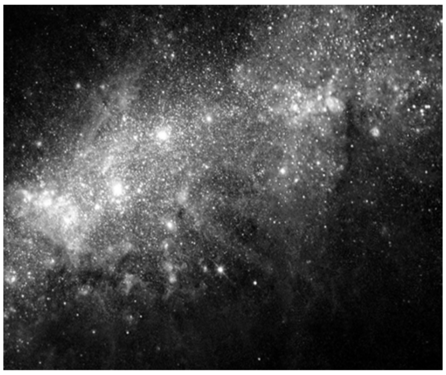


Quasars as Sources of Lyman Continuum Radiation at High Redshift



Probably too few…

Starburst Galaxies



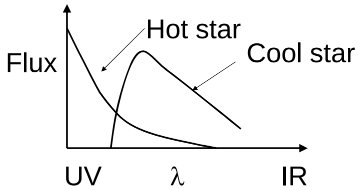
Why Do Starburst Galaxies Produce Lots of Lyman Continuum Photons?

- Stars are born in the mass range ~ 0.08—120 solar masses
- The highest-mass stars have the shortest lifetimes (a few Myr)

→ Large numbers of high-mass stars are only found in galaxies that actively form stars

- High-mass stars are typically hotter than low-mass stars
- Hot stars emit more UV radiation (stars are almost black bodies)

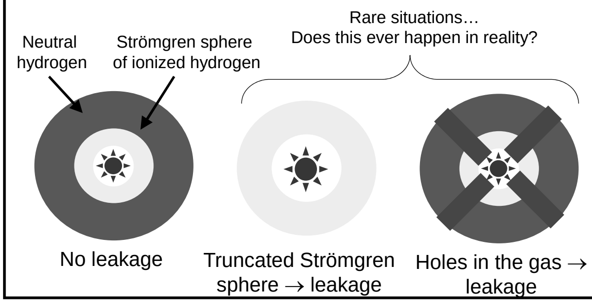
→Starbursts emit more Lyman continuum radiation than galaxies with low star formation rates

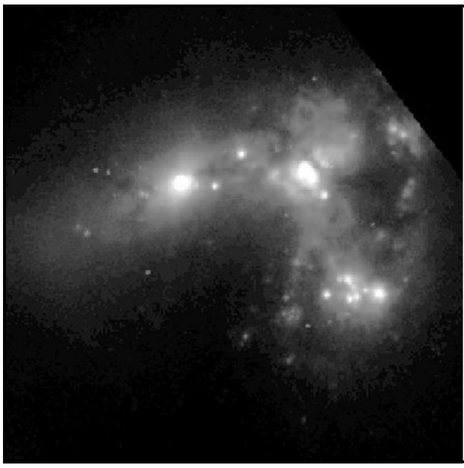


Black: Neutral hydrogen
Blue: Ionized hydrogen
Red/White: Partially ionized hydrogen
Yellow: Galaxies

Starburst galaxies

Additional caveat: Starburst galaxies must have significant Lyman continuum escape fractions to contribute to reionization





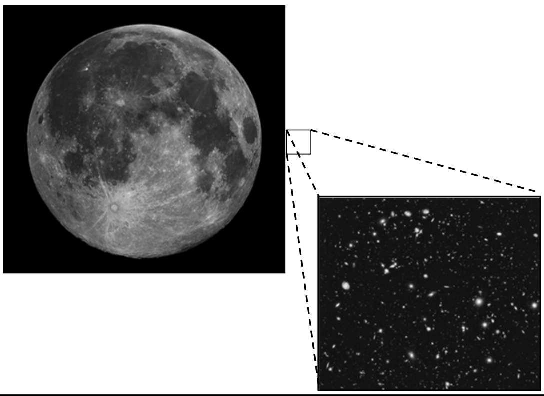
Haro 11 – The first detection of Lyman continuum leakage in the local Universe

Population III stars

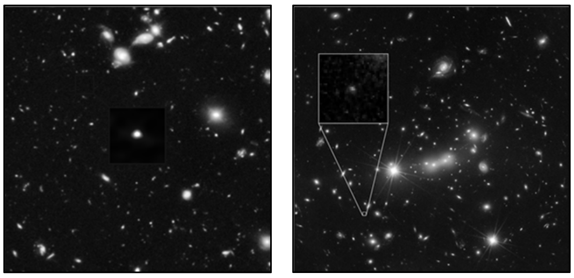
- Population I stars (young, metal-rich, disk)
- Population II stars (old, metal-poor, stellar halo)
- Population III stars (the oldest stars, metal-free)

Population III stars may have been very massive (~10 — 1000 solar masses)
→ Short-lived, but produce a lot of Lyman continuum emission during their lifetimes

Intermission: What is this?

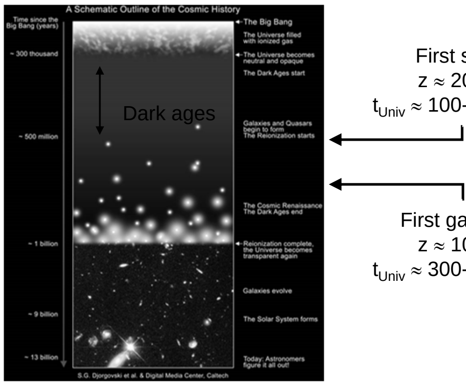


The most distant galaxies – so far



Snapshots of galaxies as they were about 500 million years after the Big Bang

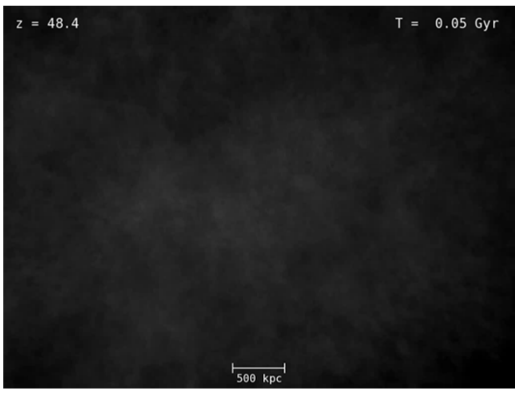
The end of the Dark Ages



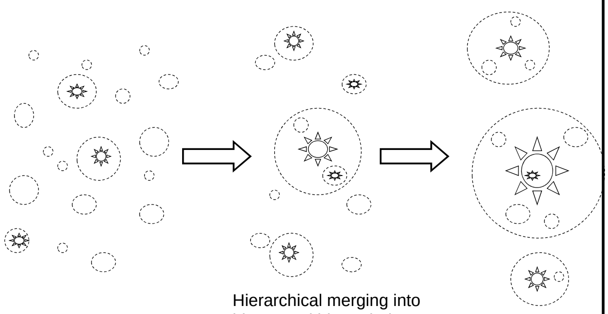
First stars
$z \approx 20\text{-}30$
$t_{Univ} \approx 100\text{-}200$ Myr

First galaxies
$z \approx 10\text{-}15$
$t_{Univ} \approx 300\text{-}500$ Myr

Intermission: What is this?



Cold Dark Matter → Hierarchical galaxy formation

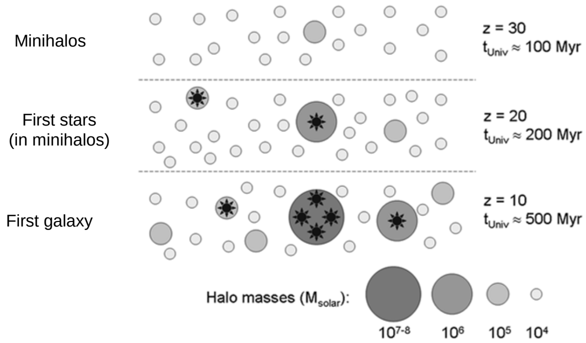


Dark matter halos with some star formation

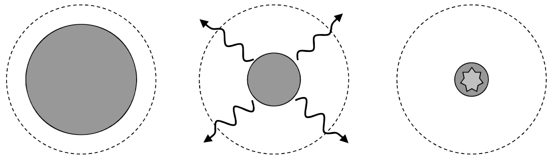
Hierarchical merging into bigger and bigger halos and galaxies. Mergers induce new star formation.

Today

The first stars and galaxies



Pop III stars forming in minihalos



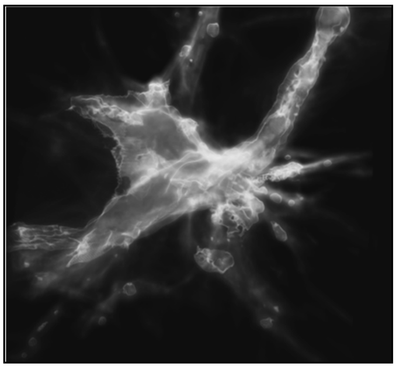
Metal-free gas in a 10^5-10^6 Msolar halo ("minihalo") at z=20-30

Radiative cooling → gas collapses towards the center

Small number of pop III stars formed (possibly just one)

Lack of metals at high z → Star formation relies on molecular cooling
Limited fragmentation → Pop III stars very massive (~10-100 Msolar)

Intermission: What is this?



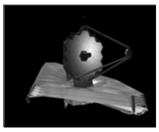
Hunting for the first stars and galaxies

Now

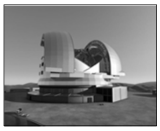
HST

VLT

Future

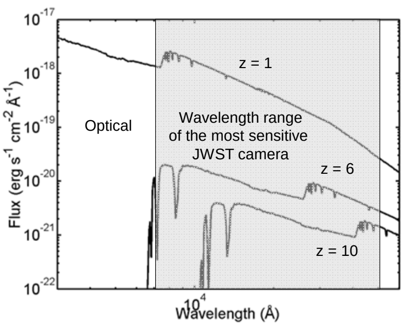


JWST (2018)

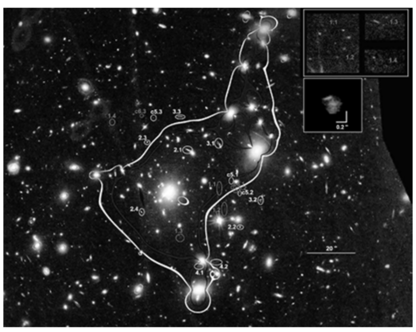


E-ELT (2021)

Infrared – the prime wavelength range for studying the first galaxies

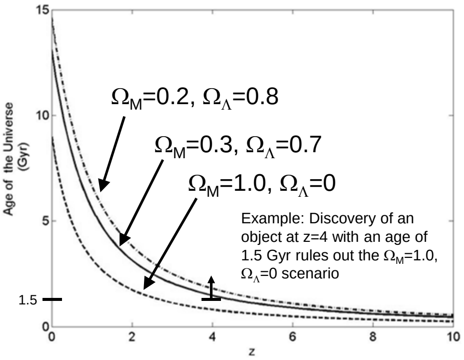


Graviational telescopes



Foreground galaxy cluster magnifies background objects → Objects otherwise too faint to be detected can be seen

High-Redshift Objects as Probes of Cosmology: Ages



Spectroscopic Age Determinations of Galaxies

Stars typically become redder when they grow older → The shape of the spectrum of a galaxy (containing billions of stars) is indicative of the age