

Physics of Galaxies 2016
10 credits
Lecture 4: Disks and ellipticals



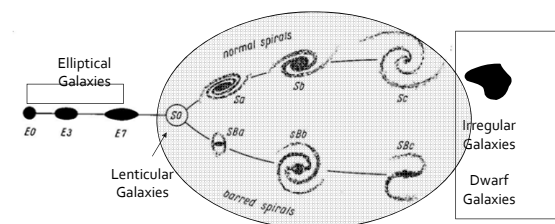
Outline I

- Disk galaxies
 - Surface brightness profiles
 - Stars and gas
 - Rotation curves
 - The Tully-Fisher relation
 - Spirals and bars

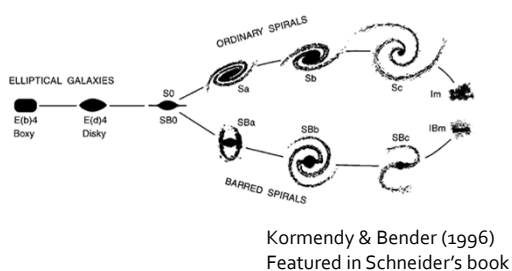
Outline II

- Elliptical galaxies
 - Surface Brightness Profiles
 - Stars
 - cD-Galaxies
 - Triaxiality
 - Stellar Motions
 - The Faber-Jackson Relation
 - Masses

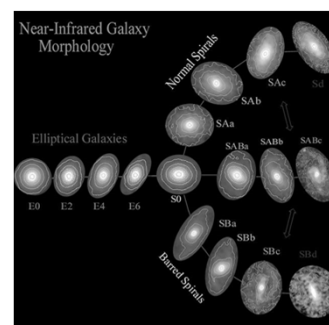
Recall the Hubble Tuning Fork



Alternative version: More elliptical subclasses



Alternative version: More spiral subclasses



Disk galaxies

- Sequence:

S0-Sa-Sb-Sc-Sd-Sm

SB0-SBa-SBb-SBc-SBd-SBm

Early-type disks

Late-type disks

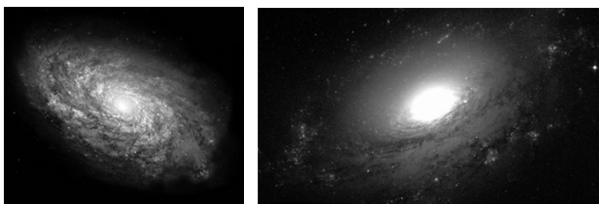
- Outside the original Hubble Tuning fork:

- Sd-galaxies: Bulgeless disks
- Sm-galaxies: Magellanic spirals (almost irregular, prototype LMC)

Disk galaxies

	S0-Sa	Sd-Sm
Spiral arms:	Absent or tight	Open spiral
Bulges:	Big	Small
Color (B-V):	Red (0.7-0.9)	Blue (0.4-0.8)
Young stars:	Few	Many
HII-regions:	Few, faint	Many, bright
Surface brightness:	High	Low
Mass:	High	Low
Rotation:	Fast rising	Slow rising

Intermission: Which of these disks is the most "early-type"?



Surface Brightness

$$\alpha \approx \frac{D}{d}$$

Size of object (D)
Distance to object (d)

$$I(r) = \frac{F}{\alpha^2} = \frac{L / 4\pi d^2}{D^2 / d^2} = \frac{L}{4\pi D^2}$$

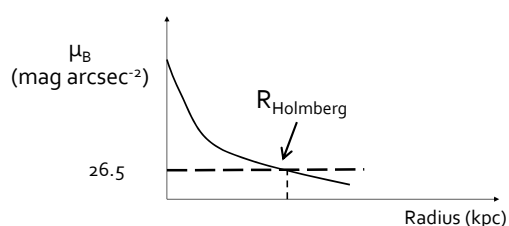
$$\mu(r) \propto -2.5 \log_{10} I(r)$$

- $I(r)$ usually $L_{\odot} \text{ kpc}^{-2}$, but $\mu(r)$ in mag arcsec^{-2}
- Determines observability of extended objects (e.g. galaxies)
- $I(x)$ independent of distance(!) in local universe...
- ... but subject to factor $(1+z)^{-4}$ of redshift dimming → One reason why high-redshift objects are extremely difficult to detect

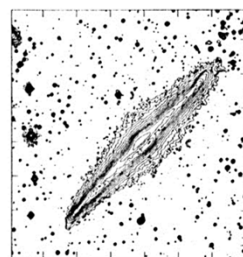
Surface Brightness

- Sizes of galaxies often given out to a specified isophote:

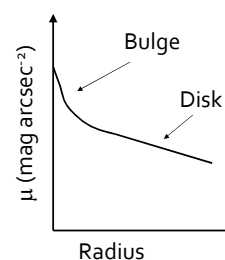
- R_{25} : Radius at 25 mag arcsec^{-2} in B-band
- Holmberg radius: Radius at 26.5 mag arcsec^{-2} in B-band



Surface Brightness Profiles I



Isophotes
(constant surface brightness)



Neutral hydrogen

- Flux in 21 cm line $\rightarrow$ HI mass:

$$\frac{M(\text{HI})}{M_{\text{solar}}} = 2.356 \times 10^5 D^2 \int F_{\nu} dV_r$$

Distance
In Mpc

Integration
over line profile

Molecular hydrogen

- H_2 most abundant molecule, but difficult to observe in emission
- 2.6 mm line of CO can be used as tracer:
 - $M(\text{H}_2)/F(\text{CO})=X$
 - However: the conversion factor X depends on metallicity; very uncertain in metal-poor galaxies

Gaseous and stellar motions

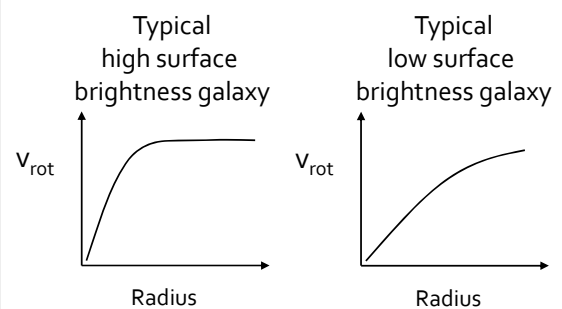
In disks:

Average
rotational
velocity

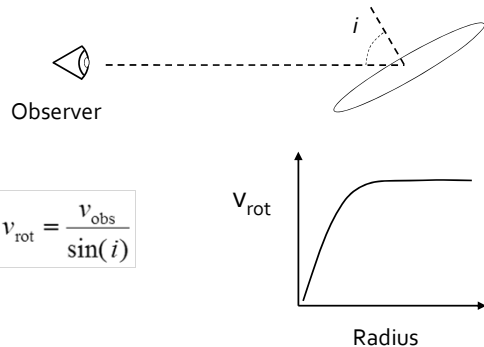
Typical
velocity
dispersion

$$\frac{v_{\text{rot}}}{\sigma_v} \sim 10$$

Rotation curves



Rotation curves



Rotation curves

Recall from lecture 3:

$$M(< R) = \frac{v_{\text{rot}}(R)^2 R}{G}$$

Rotation curve decomposition

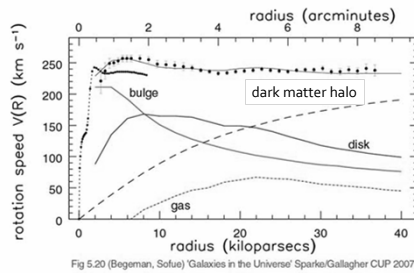


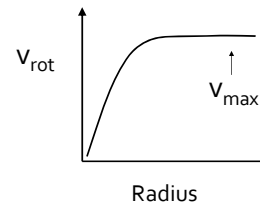
Fig 5.20 (Begeman, Sofue) 'Galaxies in the Universe' Sparks/Gallagher CUP 2007

Typical global $M/L \sim 10-100$

The Tully-Fisher relation

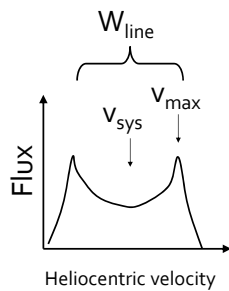
$L \propto V_{\max}^4$
Example:

$$\frac{L_H}{3 \times 10^{10} L_{H\text{solar}}} = \left(\frac{v_{\max}}{196 \text{ km/s}} \right)^{3.8}$$



The Tully-Fisher relation II

Don't need rotation curve — you can also use HI spectral line profile

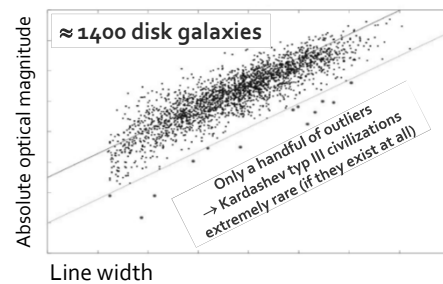


$$V_{\max} \approx \frac{W_{\text{line}}}{2}$$

In one of the exercises, we use the following form of the TF relation:

$$M_H \approx -9.50(\log_{10} W - 2.50) - 21.67,$$

Weird stuff: The Tully-Fisher relation as a tool to search for extraterrestrial intelligence



Shameless self-promotion: Zackrisson, E., Calissendorff, P., Asadi, S., Nyholm, A. 2015, *Astrophysical Journal*, 810, 23

Spiral patterns I: A "Grand Design" Spiral

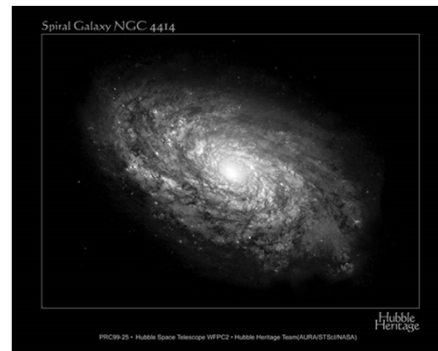


Spiral Galaxy NGC 2997 (VLT UT1 + FORS1)

ESO PR Photo 17a/99 (6 March 1999)

© European Southern Observatory

Spiral patterns II: A Flocculent Spiral



Spiral Galaxy NGC 4414

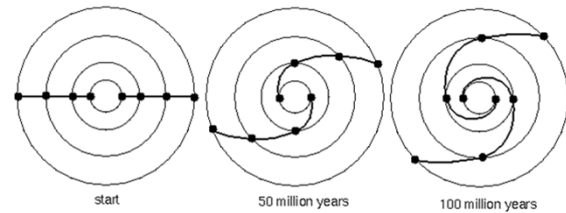
PRC925 - Hubble Space Telescope WFC3 - Hubble Heritage Team (AURA/STScI/NASA)

Hubble
Heritage

Intermission:
What type of spiral is this?

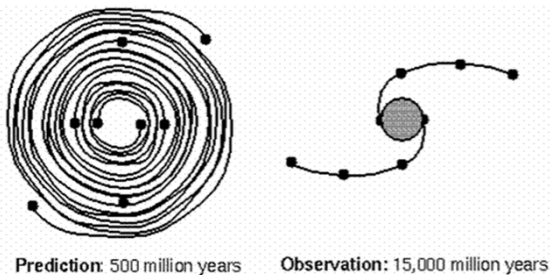


Spiral patterns III:
Differential rotation



Differential rotation: stars near the center take less time to orbit the center than those farther from the center. Differential rotation can create a spiral pattern in the disk in a short time.

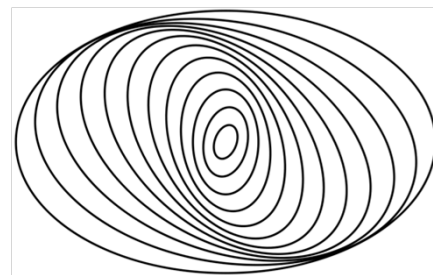
Spiral patterns IV:
The winding-up dilemma



Prediction: 500 million years

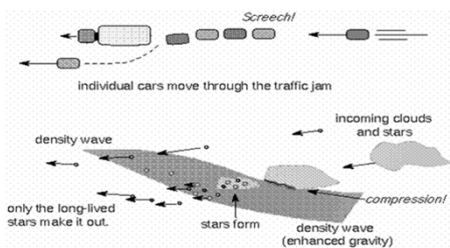
Observation: 15,000 million years

Spiral patterns V: Density waves



Stars on elliptical orbits with different orientations → stars in spiral arms continuously replaced

Spiral patterns VI: Density-wave theory



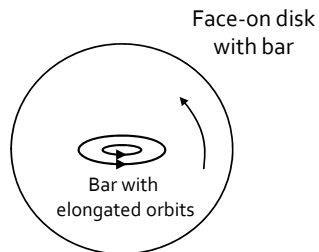
Spiral density waves are like traffic jams. Clouds and stars speed up to the density wave (are accelerated toward it) and are tugged backward as they leave, so they accumulate in the density wave (like cars bunching up behind a slower-moving vehicle). Clouds compress and form stars in the density wave, but only the fainter stars live long enough to make it out of the wave.

Spiral patterns VII:
Problems with density waves

- From where does the density wave get its energy?
 - From the rotation of the disk?
 - From a companion galaxy?
 - Internal forces from a central bar?
- Spiral patterns remain mysterious...

Bars

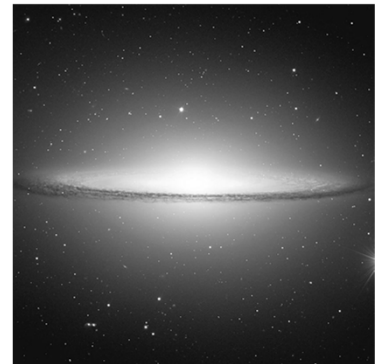
- At least 50% of all disk galaxies have bars
- Bars are not density waves!
- Elongated orbits



Bulges

- In bulges:

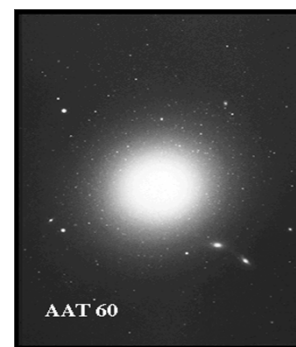
$$\frac{v_{\text{rot}}}{\sigma_v} \sim 1$$



Intermission: The Galaxy Zoo Project

<https://www.galaxyzoo.org/>

Elliptical Galaxies



Surface Brightness Profiles of Ellipticals I

$R^{1/4}$ or De Vaucouleurs law ($n \approx 4$)

$$I(R) = I(0) \exp \left(- (R / h_R)^{1/n} \right)$$

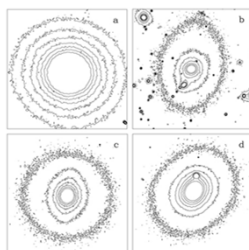
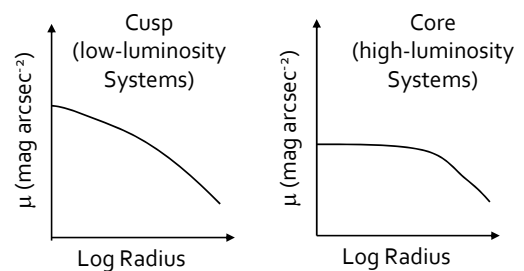


Fig 6.1 (R. de Jong) 'Galaxies in the Universe' Springer/Gallagher CUP 2007

Surface Brightness Profiles of Ellipticals II



Late-breaking news: The core is due to influence from the central supermassive black hole. The radius of the core correlates strongly with the black hole mass (Thomas et al. 2016, Nature)!

Mass Determinations for Ellipticals

- More difficult than for disk galaxies
- A few methods:
 - For gas-rich Es: HI rotation curves
 - X-ray gas: $M=f(\rho_{\text{gas}}, r, T)$
 - Virial theorem: $M=f(\sigma, r)$ with
 - Stellar $\sigma(r)$ from absorption lines
 - Stellar $\sigma(r)$ and v_{rot} from planetary nebula emission lines