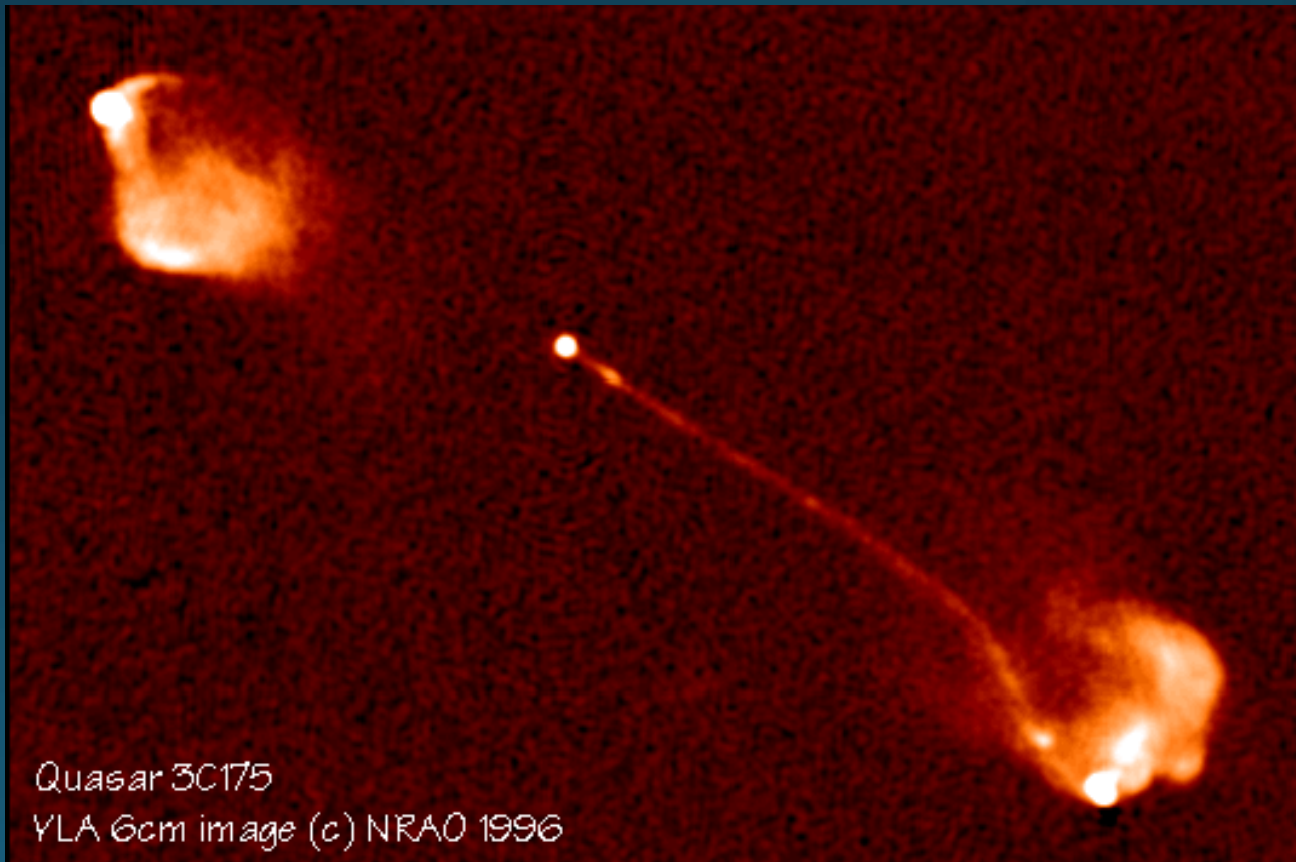


Physics of Galaxies 2016

10 credits

Lecture 6: Active galaxies and black holes



Outline

- Introduction to the database exercise
- Black holes in galaxies
 - How do we find them?
 - Relation between black hole mass and other galaxy properties
- Active galactic nuclei (AGN)
 - Basic AGN anatomy
 - Different types of AGN
 - Unification model
 - Quasar host galaxies
 - Quasar absorption systems

Database exercise



UPPSALA UNIVERSITET

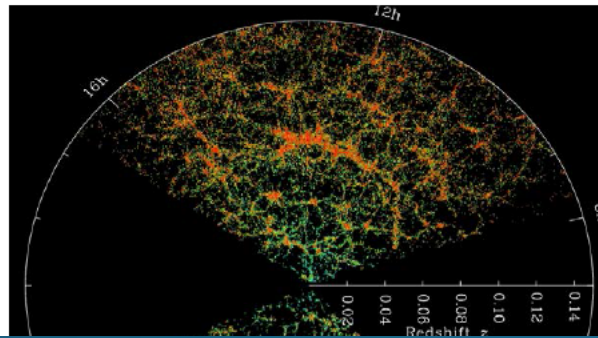


Studying galaxies with the Sloan Digital Sky Survey

Laboratory exercise, Physics of Galaxies, Spring 2016 (Uppsala Universitet)

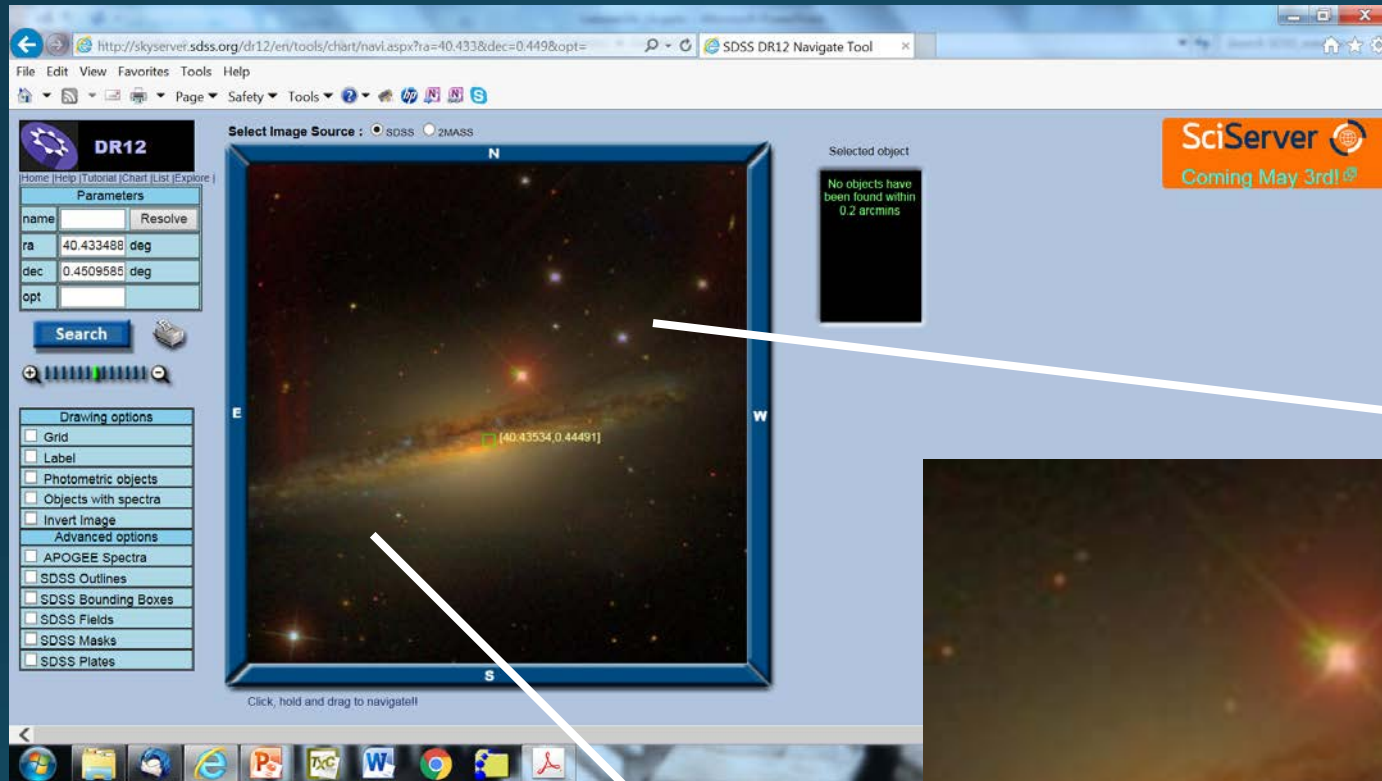
by

Beatriz Villarroel



Deadline June 7 – preferably no more than 5 pages

Obstacles: Astronomical coordinates



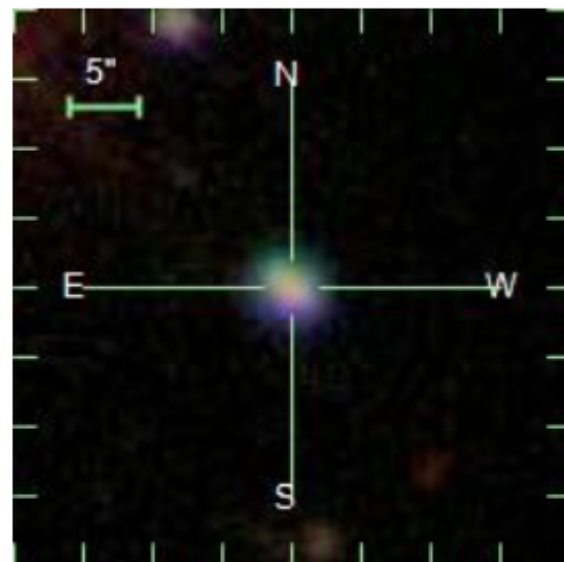
Obstacles: Bewildering photometric data

SDSS J113459.47+002509.1

GALAXY ra=173.74782136, dec=0.41921516, ObjId = 588848900446814264

Column names link to glossary entries. Move mouse over a column name to get its units.


mode	PRIMARY
status	TARGET PRIMARY OK_STRIPE OK_SCANLINE PSEGMENT RESOLVED OK_RUN GOOD SET
flags	STATIONARY MOVED BINNED1 CHILD
PrimTarget	TARGET_GALAXY
SecTarget	



<u>u</u>		<u>g</u>	<u>r</u>	<u>i</u>	<u>z</u>	
19.55		18.04	17.55	17.35	17.21	
<u>err_u</u>		<u>err_g</u>	<u>err_r</u>	<u>err_i</u>	<u>err_z</u>	
0.03		0.01	0.01	0.01	0.02	
<u>run</u>	<u>rerun</u>	<u>camcol</u>	<u>field</u>	<u>obj</u>	<u>rowc</u>	<u>colc</u>
756	44	4	387	56	549.4	1974.6
<u>fiberMag_r</u>	<u>petroMag_r</u>	<u>devMag_r</u>	<u>expMag_r</u>	<u>psfMag_r</u>	<u>modelMag_r</u>	
18.05	17.55	17.55	17.55	17.99	17.55	
<u>extinction_r</u>		<u>petroRad_r</u>	<u>parentId</u>			<u>nChild</u>
0.06		1.796	588848900446814263			0

Register an account!

A bit worrisome:
The interface is
getting updated
during the
course...



The screenshot shows a web browser window with the address bar displaying `http://skyserver.sdss.org/CasJobs/`. The browser's menu bar includes File, Edit, View, Favorites, Tools, and Help. Below the menu bar is a toolbar with various icons. The website header features the SDSS logo and the text "SDSS Query / CasJobs". A navigation bar contains links for Help, Tools, Create Account, and Login. A prominent orange box contains an announcement about the SciServer migration. Below this, a paragraph describes CasJobs as an online workbench. A list of features follows, and a footer section provides contact information and a link to the FAQ or guide.

On May 3rd 2016, we will roll out a new collaborative science framework called **SciServer**, which brings several enhancements to CasJobs. The primary new feature is a Single Sign-On system that will unify access to CasJobs, SkyServer, and many additional data resources. At that time, you will need to migrate your CasJobs account into the new SciServer system. For more information on all upcoming new features, please see our **New Features** section on the SciServer website.

CasJobs is an online workbench for large scientific catalogs, designed to emulate and enhance local free-form query access in a web environment.

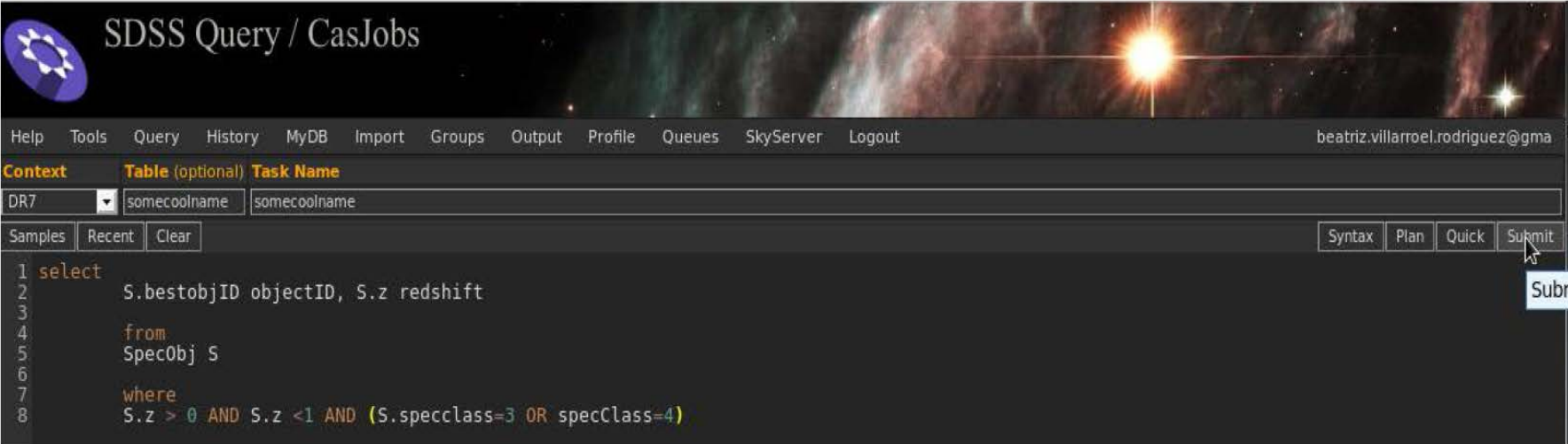
Some features of this application include...

- Both synchronous and asynchronous query execution, in the form of 'quick' and 'long' jobs.
- A query 'History' that records queries and their status.
- A server-side, personalized user database, called 'MyDB', enabling persistent table/function/procedure creation.
- Data sharing between users, via the 'Groups' mechanism.
- Data download, via MyDB table extraction, in various formats.
- Multiple interface options, including a browser client as well as a java-based command line tool.

For more information, try the CasJobs **FAQ**, or **guide**. Or, to get started right away, either **login** or **create** an account.

Contact
heads/master-0-q36fdc80

SQL (Structured Query Language)



The screenshot shows the SDSS Query / CasJobs web interface. At the top, there's a header with the SDSS logo and the text "SDSS Query / CasJobs". Below the header is a navigation bar with links: Help, Tools, Query, History, MyDB, Import, Groups, Output, Profile, Queues, SkyServer, and Logout. The user's email address, beatriz.villarroel.rodriguez@gmail.com, is displayed on the right. The main interface has a "Context" section with a dropdown menu set to "DR7" and two input fields for "Table (optional)" and "Task Name", both containing "somecoolname". Below this are buttons for "Samples", "Recent", and "Clear". On the right side of the interface are buttons for "Syntax", "Plan", "Quick", and "Submit". A mouse cursor is hovering over the "Submit" button, and a tooltip that says "Submits your query" is visible. The main area is a text editor with a dark background, showing a SQL query being typed. The query is as follows:

```
1 select
2   S.bestobjID objectID, S.z redshift
3
4   from
5   SpecObj S
6
7   where
8   S.z > 0 AND S.z < 1 AND (S.specclass=3 OR specClass=4)
```

Fig.6. The Casjobs Query interface. Write your SQL code for selecting your favourite objects here.

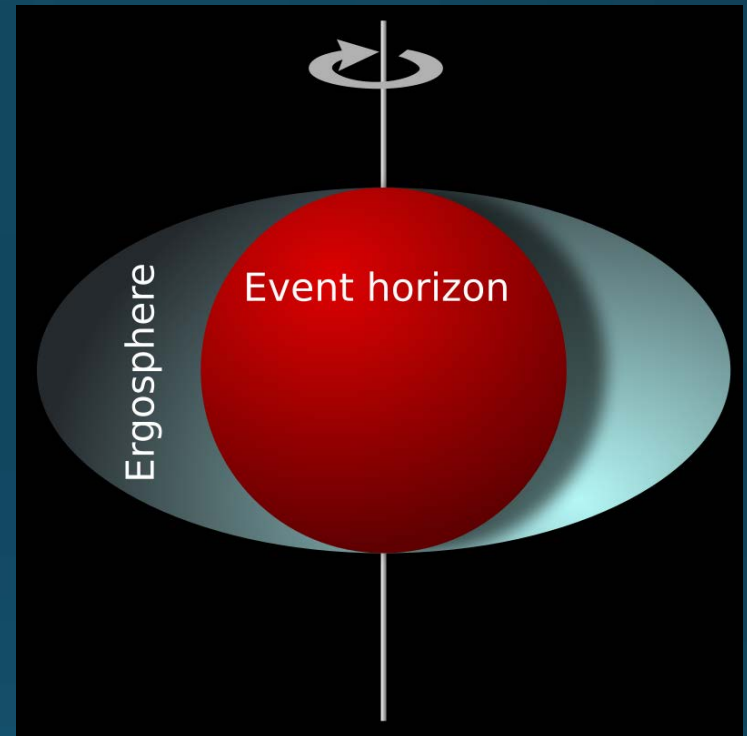
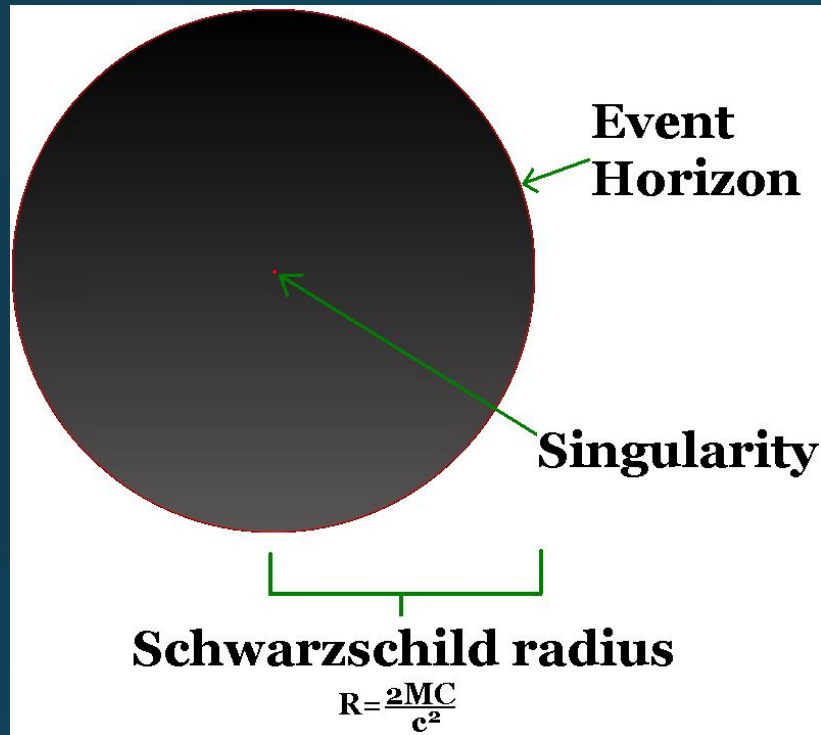
Connection to your essay

Finally, I'd like you to reconnect this exercise to the **topic of your written essay or a specific research problem described in it**. Many research problems can be solved using the SDSS and sometimes one only has to figure out a clever way to approach them. Other times, even if one cannot solve the problem using the SDSS directly, the survey is helpful for e.g. selecting interesting candidate objects.

Exercise 4: *(a) Describe the most interesting research problem from the essay very briefly.*

(b) How could you approach this fascinating research problem using the SDSS? If it is not possible to use the SDSS, explain what the solution requires that cannot be met by the SDSS data.

Black holes



Non-rotating black hole

Rotating black hole

Current observational evidence supports the existence of stellar-mass black holes ($\sim 10 M_{\odot}$) and supermassive black holes ($\sim 10^6$ - $10^{10} M_{\odot}$).

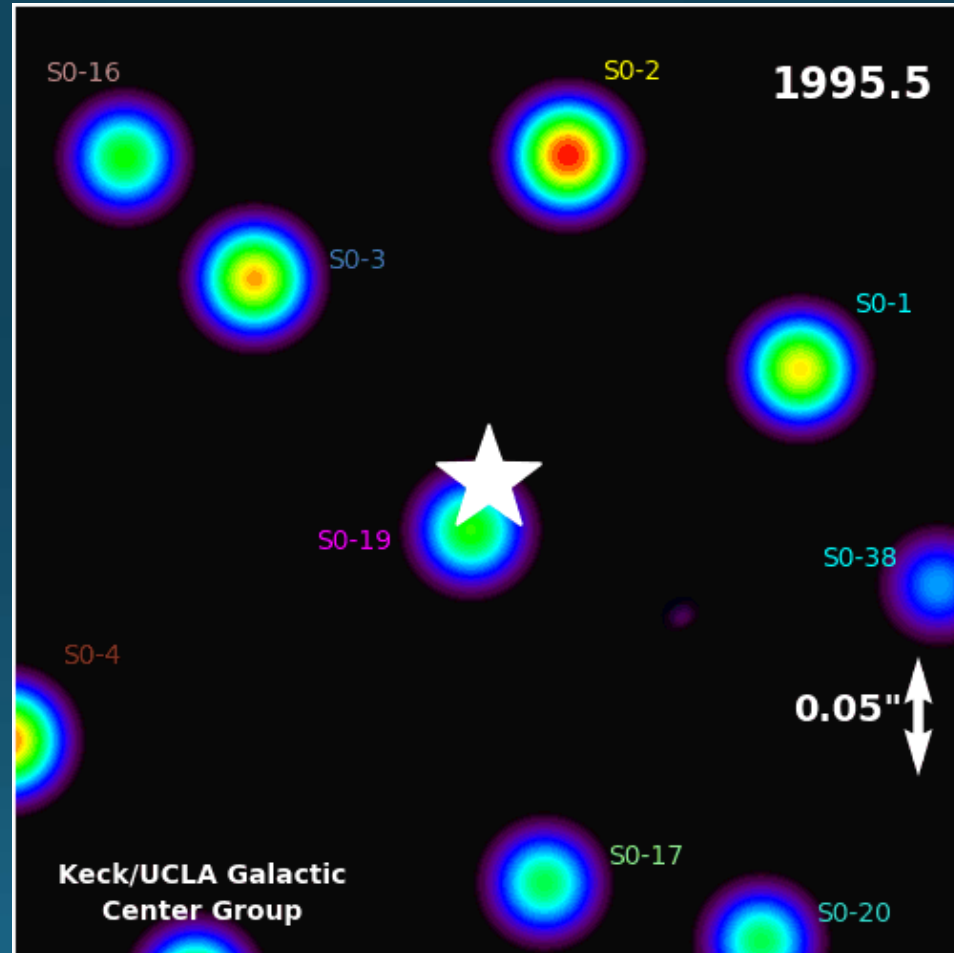
The evidence for intermediate-mass black holes ($\sim 10^2$ - $10^5 M_{\odot}$) remains scant

Hunting down black holes

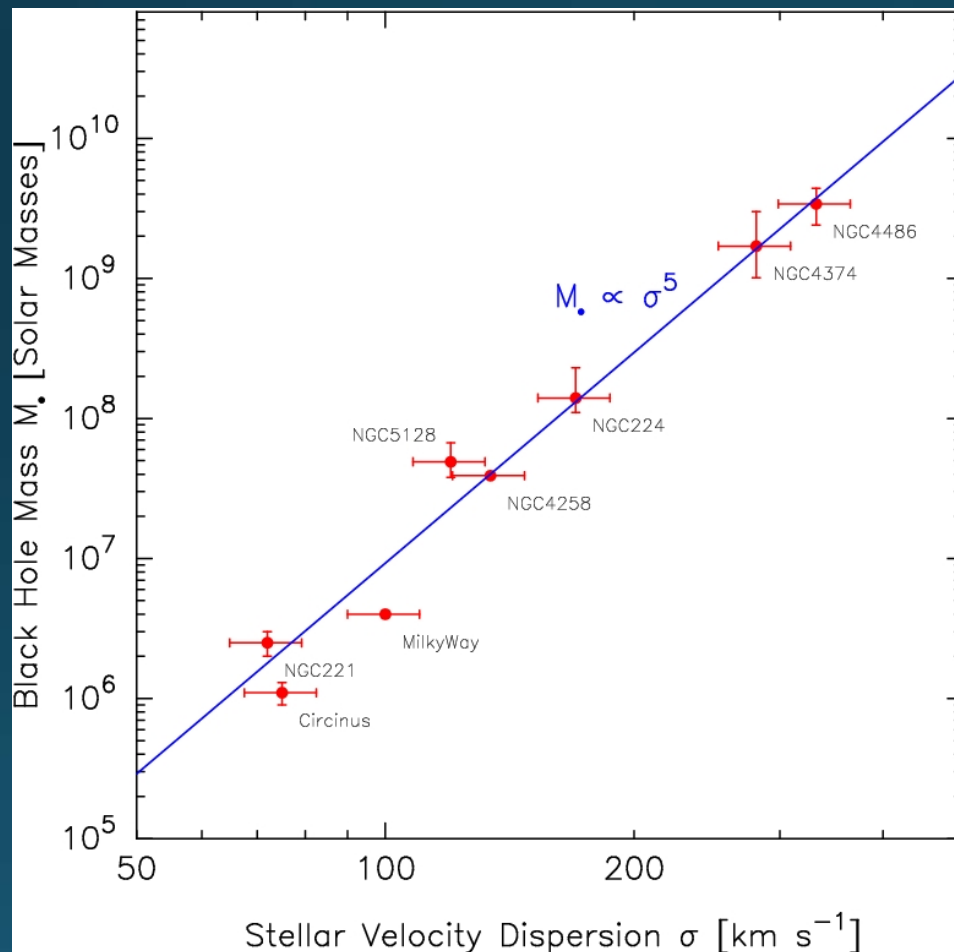
- Motion of stars
 - Milky Way: Proper motion of individual stars
 - Other galaxies – velocity distribution from integrated stellar population spectrum
- Active galactic nuclei
- Gravitational waves
- X-ray binaries
- Ultraluminous X-ray sources
- Gravitational lensing

The black hole at the centre of the Milky Way

Milky Way:
 $M_{\text{BH}} \approx 4 \times 10^6 M_{\odot}$



Relation between black hole mass and stellar velocity dispersion (or mass) of bulge



$$\frac{M_{\text{SMBH}}}{M_{\text{Bulge}}} \sim 0.002$$

Supermassive black holes in AGN

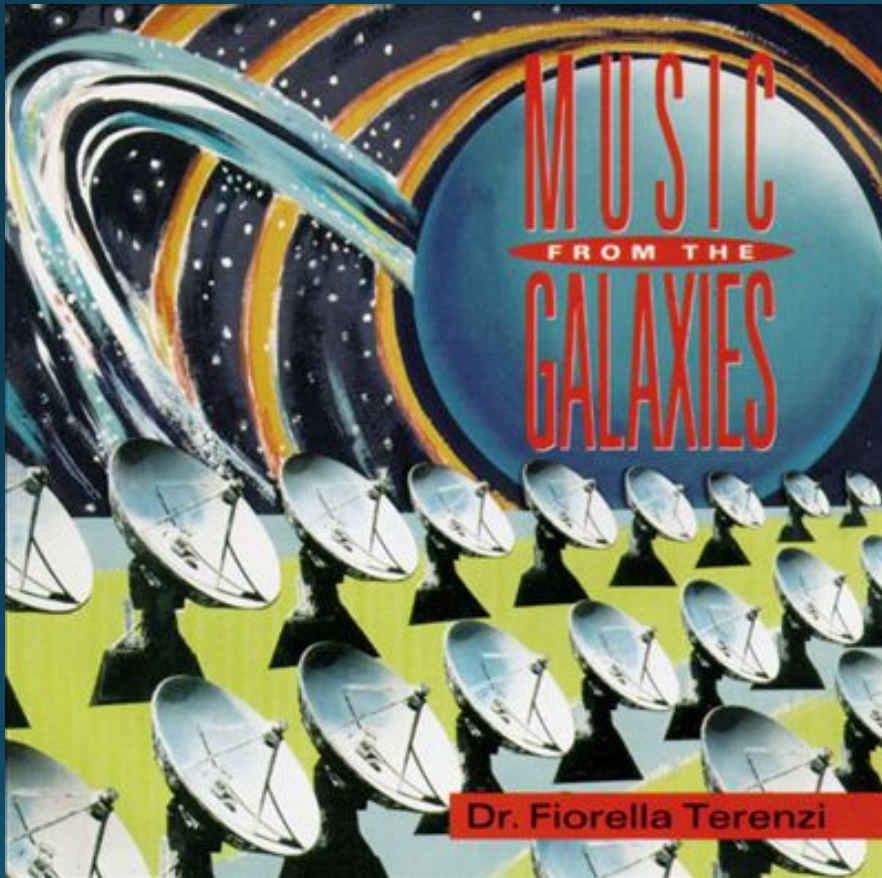
- Doppler broadened emission lines in AGN indicate gas velocities $\sim 10\,000$ km/s
- Line variability time scale (weeks) \rightarrow size of line-emitting region
- Velocity & size \rightarrow Mass(<size) & Density, indicating that the gas orbits a SMBH
- Schwarzschild radius:

$$R_s = \frac{2GM_{\text{BH}}}{c^2} \approx 3 \times \frac{M_{\text{BH}}}{M_{\text{solar}}} \text{ km}$$

Characteristics of Active Galactic Nuclei

- High luminosity produced in small region
- Fast variability
- High fraction of polarized light
- Non-thermal spectrum: Not stars!
 - Synchrotron radiation
 - Emission-line ratios → Ionization source more energetic than hottest known stars

Intermission: Music from AGN



Dr Fiorella Terenzi

Music from the Galaxies (1991):
Radio waves from the active
galaxy UGC 6697 converted
into music

Intermission: Music from AGN



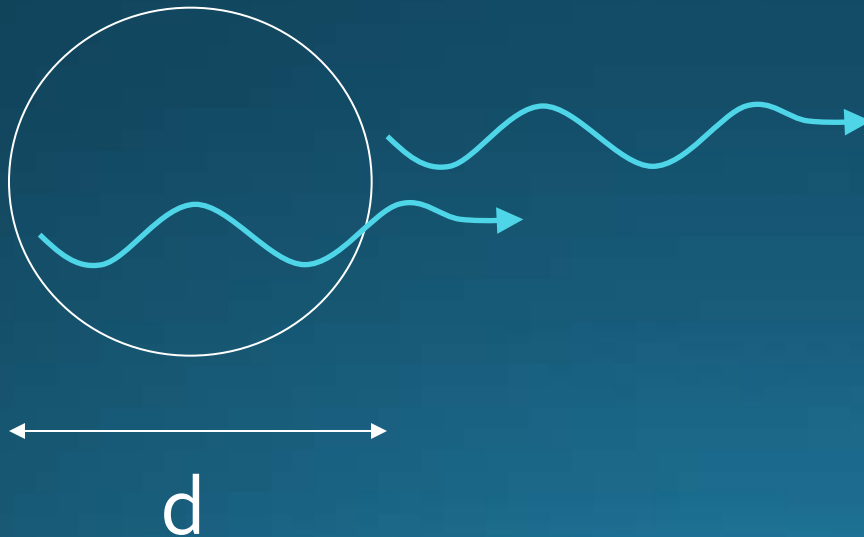
NGC 4151 (1993):

Rest-frame UV emission-line and continuum variability from the Seyfert galaxy NGC 4151 converted into music

Professor Nils Bergvall

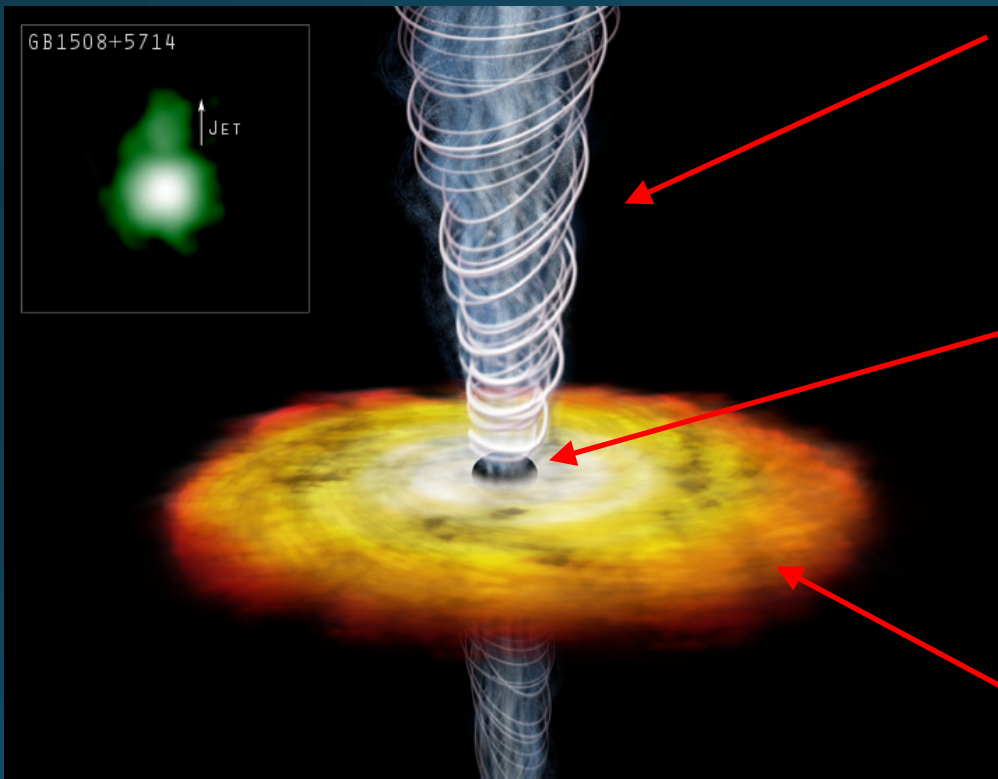
Variability-Size Relation

- Fast variability indicates that the luminosity is produced inside a small region
- Light variations on scales down to 1 hour → size smaller than the Solar system



$$\frac{d}{c} < \Delta t_{\text{var}}$$

Accretion Disks



Magnetic field channel
matter into relativistic
jets

SMBH

Angular momentum
of infalling material →
matter spirals inward
in an accretion disk

Eddington Luminosity

Too high radiation pressure of AGN may overcome inward gravitational force → upper limit on AGN luminosity which still allows material to fall inwards

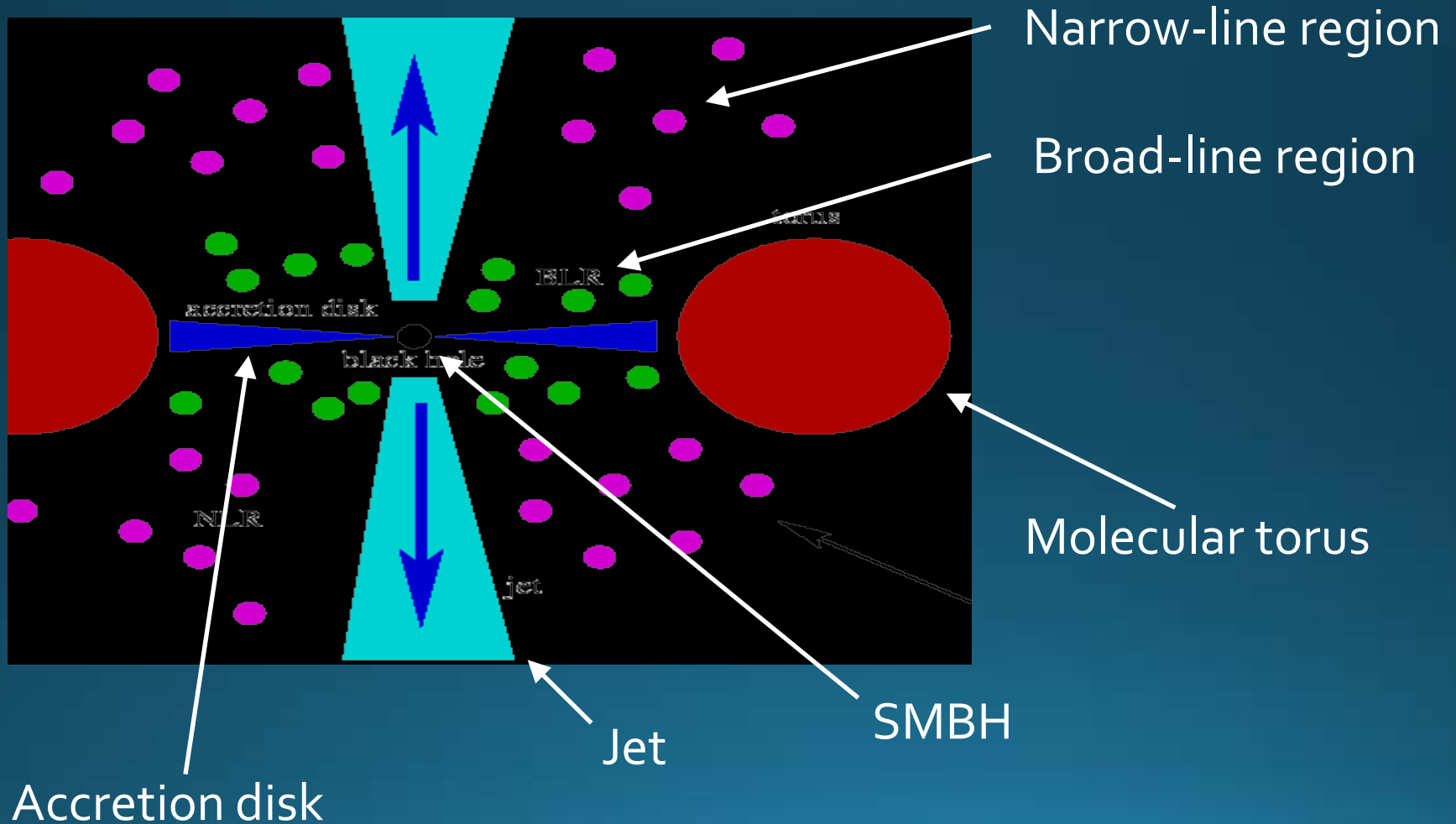
$$L_E \approx 300000 \frac{M}{M_{\text{solar}}} L_{\text{solar}}$$

Note: L_E assumes spherical accretion.
Super-Eddington luminosities (a few times L_E)
can be produced in accretion disks

Radiation Efficiency

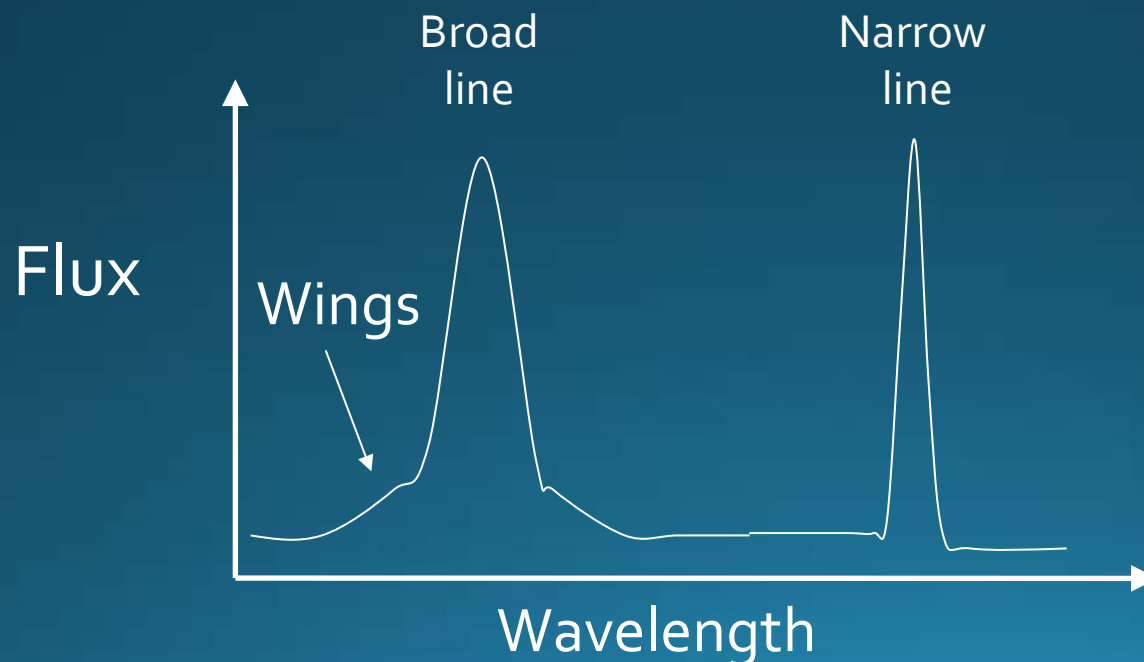
- Mass M falling into a SMBH \rightarrow
energy Mc^2 added
- Theoretical maximum:
42% of Mc^2 is converted into luminosity
The rest increases the SMBH mass
- But typically, $\leq 10\%$ of Mc^2 is converted into luminosity
- SMBHs in a typical quasar grows with $\geq 1 M_{\text{solar}}/\text{yr}$
- Activity is expected to last for $\sim 100 \text{ Myr} \rightarrow$
 $M_{\text{SMBH}} \geq 10^8 M_{\text{solar}}$ in faded quasars

The Central Power Source

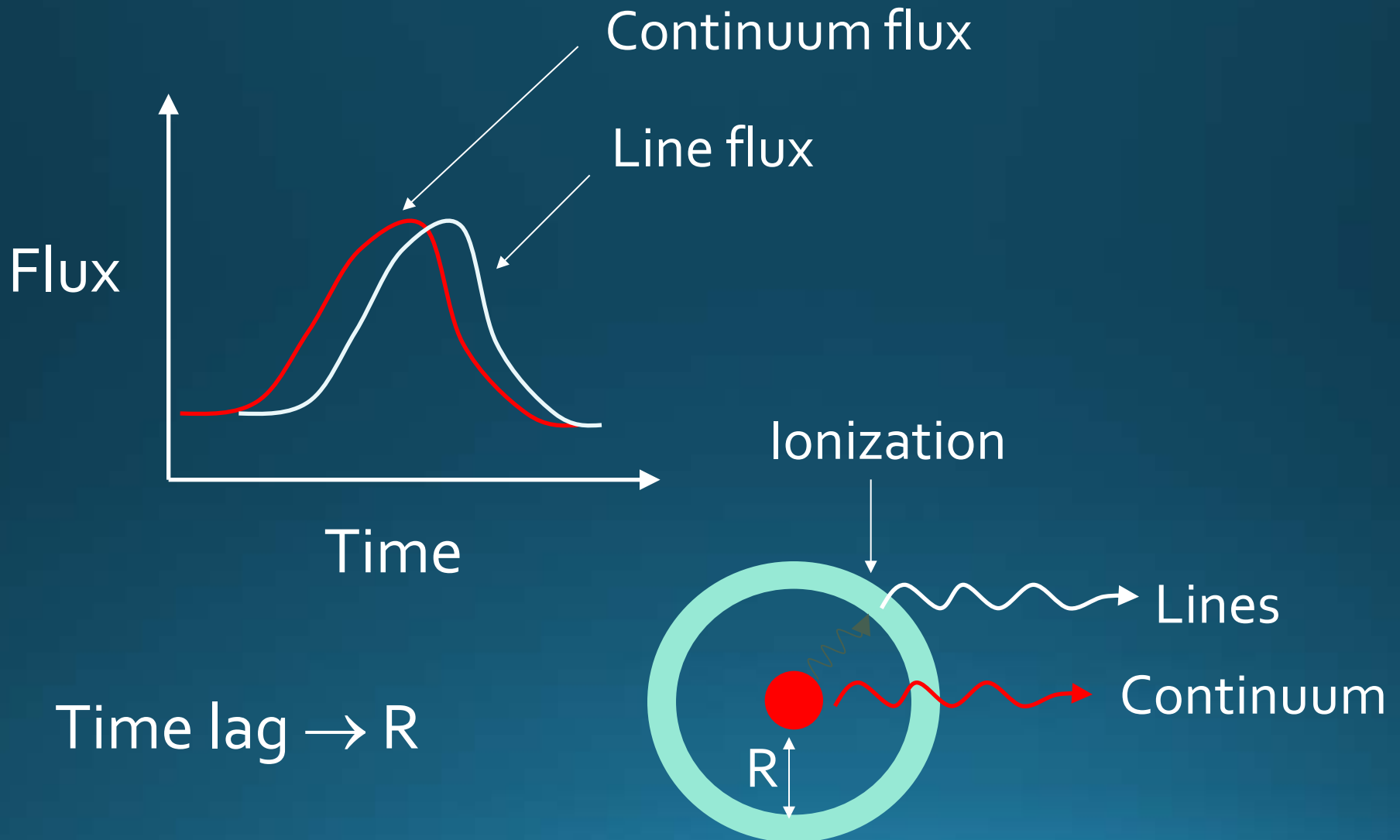


Transitions and Line Profiles

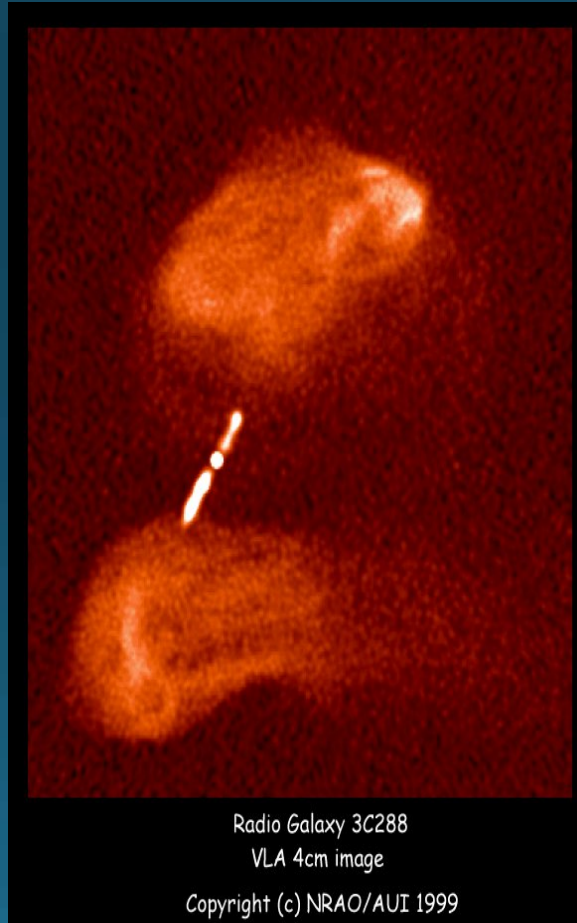
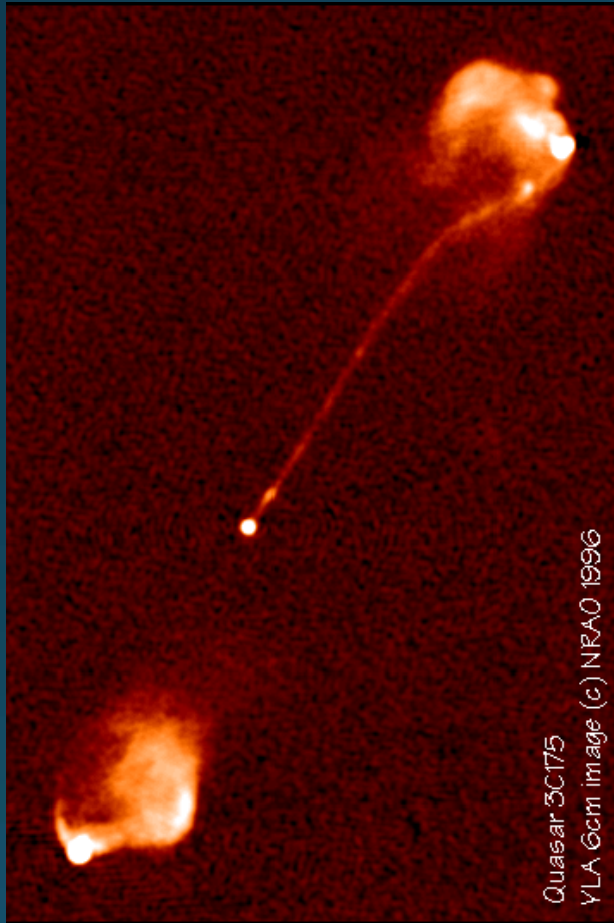
- Forbidden lines (e.g. [OII]3727Å, [OIII]5007) originate in low-density gas ($n_H \leq 10^8$ atoms/cm³)
- Allowed lines (e.g. H α , H β) originate in more dense gas
- Broad lines, $\Delta v \sim 10000$ km/s – Allowed (originate close to the nucleus, in dense medium)
- Narrow lines, $\Delta v < 1000$ km/s – Forbidden (originate further out, in less dense medium)



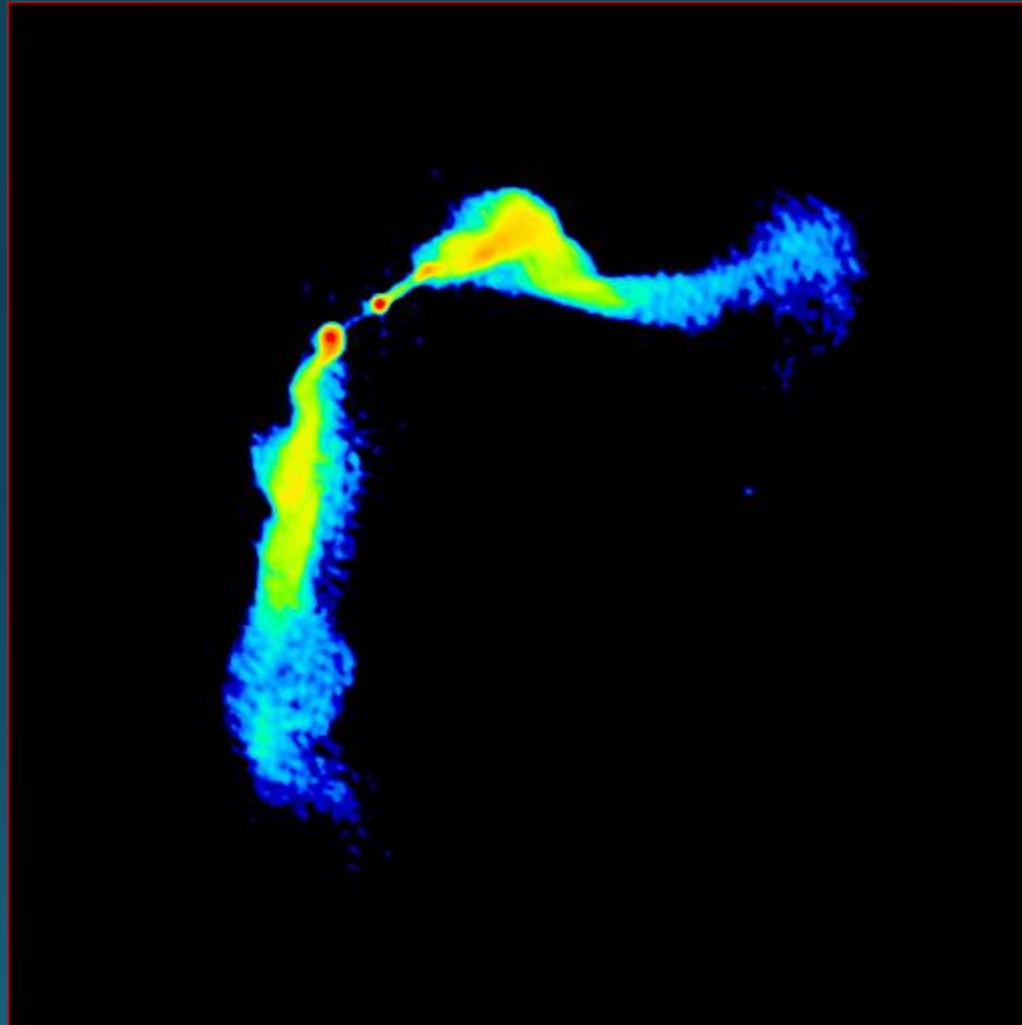
Reverberation Mapping



Jets and Lobes

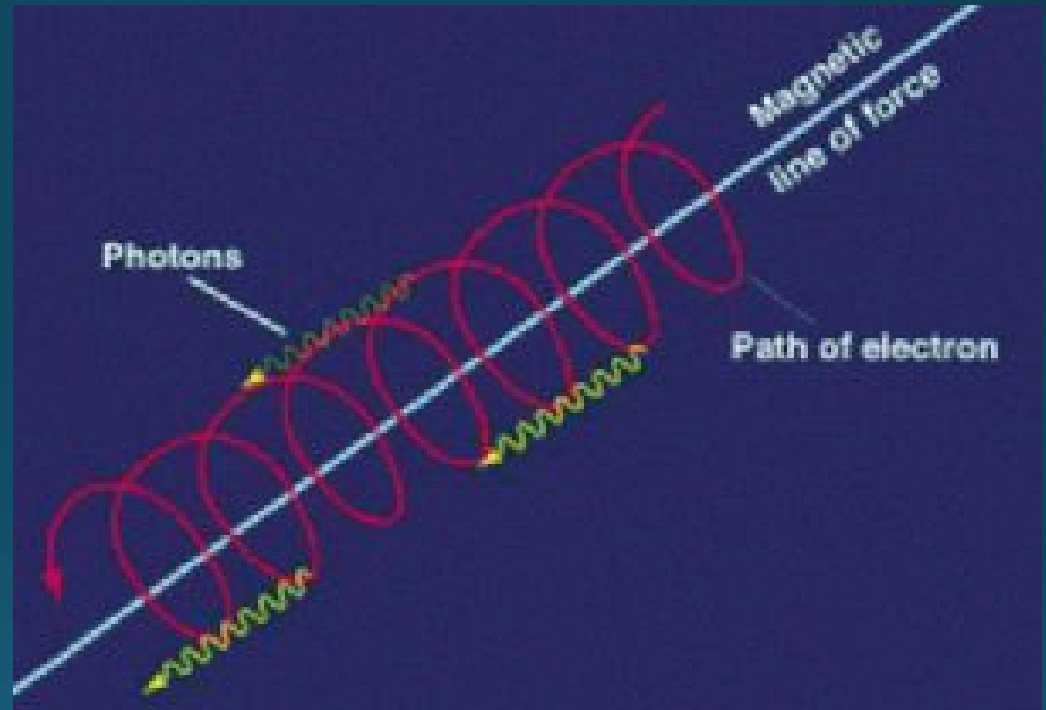


Intermission: What is happening here?



Synchrotron Radiation

- Relativistic electrons spiralling around magnetic field lines
- $B=1-10 \mu\text{G}$ in lobes (same as B close to the Sun)
- $B=0.1 \text{ G}$ in active cores



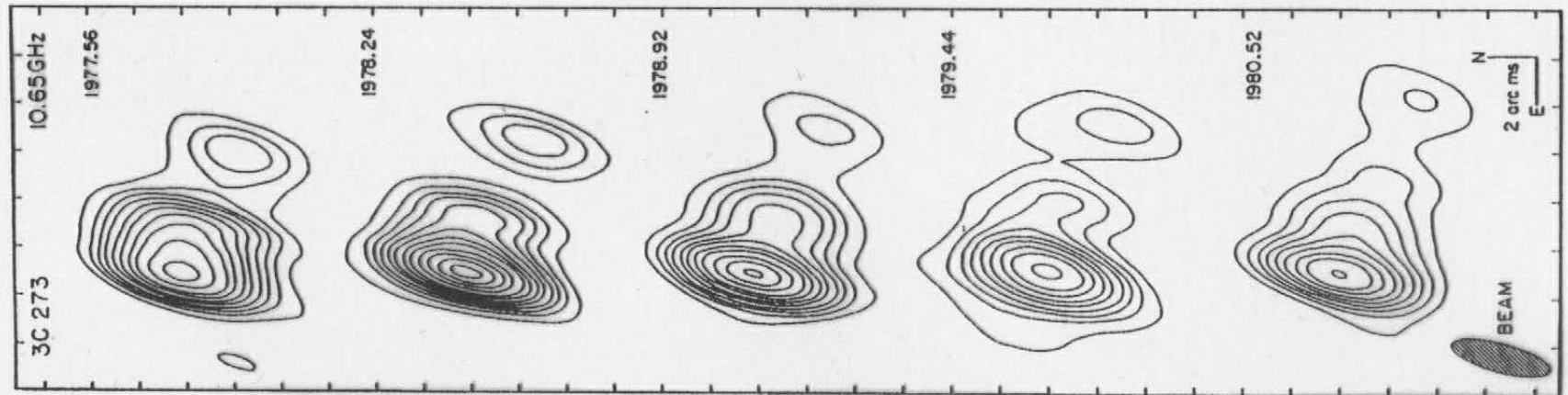
Power-Law Spectrum

$$f_\nu \propto \nu^{-\alpha}$$

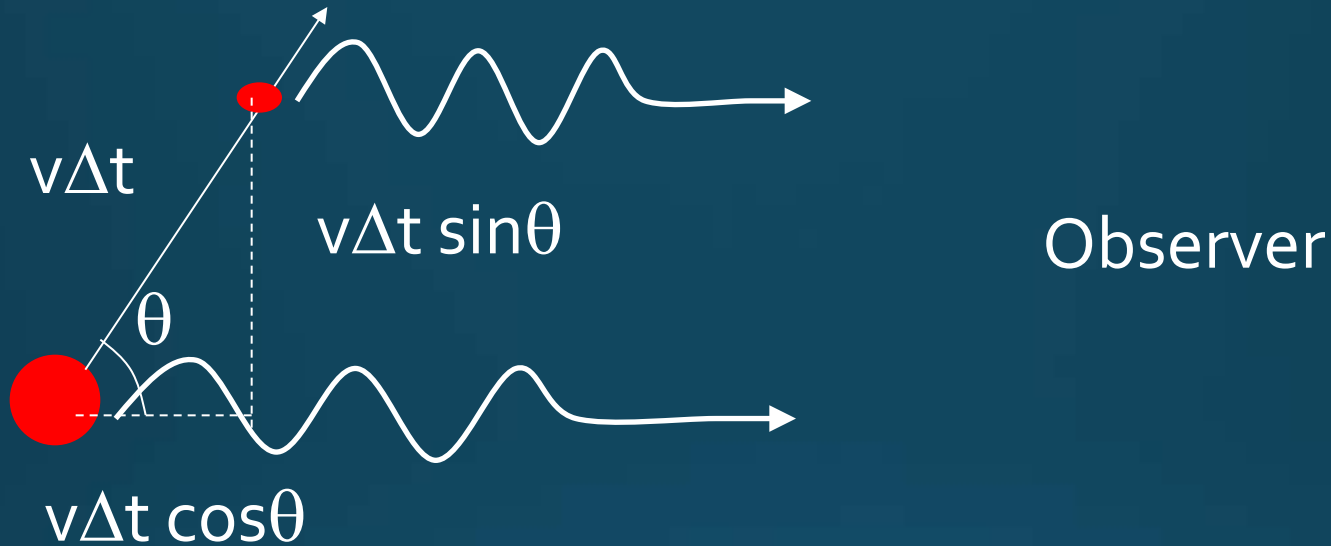
- In radio, due to synchrotron radiation:
 $\alpha=0.7\text{—}1.2$
- In optical, probably caused by thermal radiation from the accretion disk:
 $\alpha=0\text{—}2$

Superluminal motion

Blobs appear to move outwards at 5-50c
What's going on?



Superluminal motion



$$\Delta t_{\text{obs}} = \Delta t(1 - (v/c) \cos \theta)$$

In Δt_{obs} , the blob travels $v \Delta t \sin\theta$ across the sky,
With apparent velocity v_{obs} :

$$v_{\text{obs}} = \frac{v \sin \theta}{1 - (v/c) \cos \theta}$$

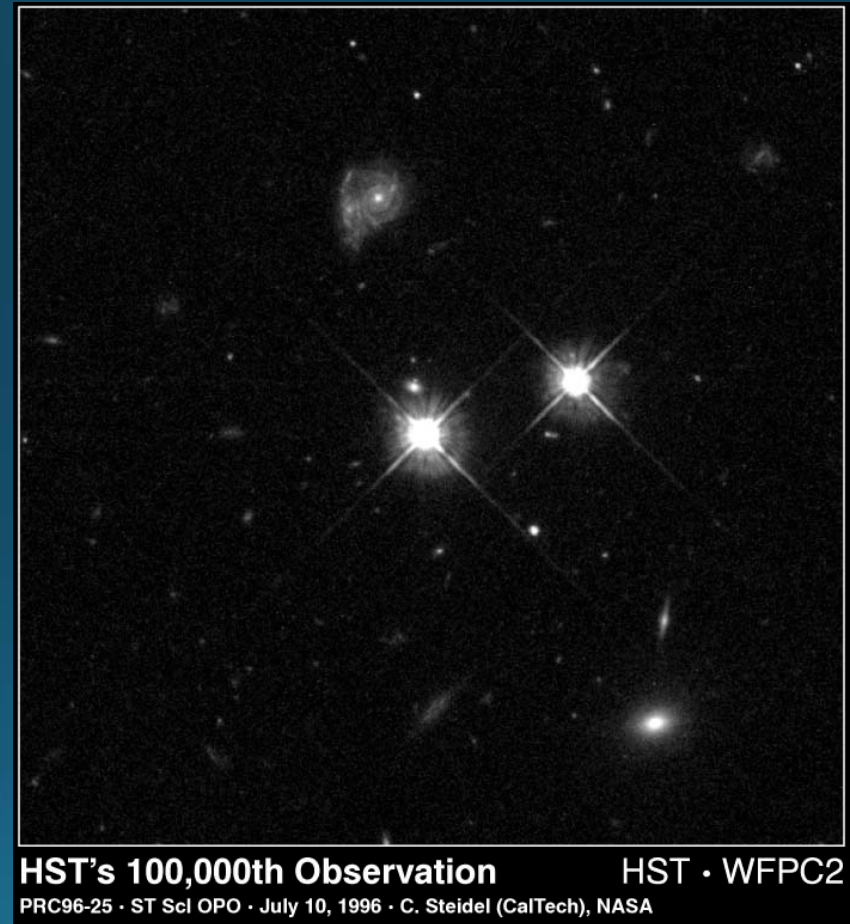
$$v \approx c \rightarrow v_{\text{obs}} > c$$

The number densities of AGN at $z=0$

Type	Number/Gpc ³
Spiral galaxies	$\sim 5 \times 10^6$
E/So galaxies	$\sim 1 \times 10^6$
Seyfert galaxies	$\sim 1 \times 10^5$
Radio galaxies	$\sim 3 \times 10^3$
Quasars	~ 100
Blazars	~ 80

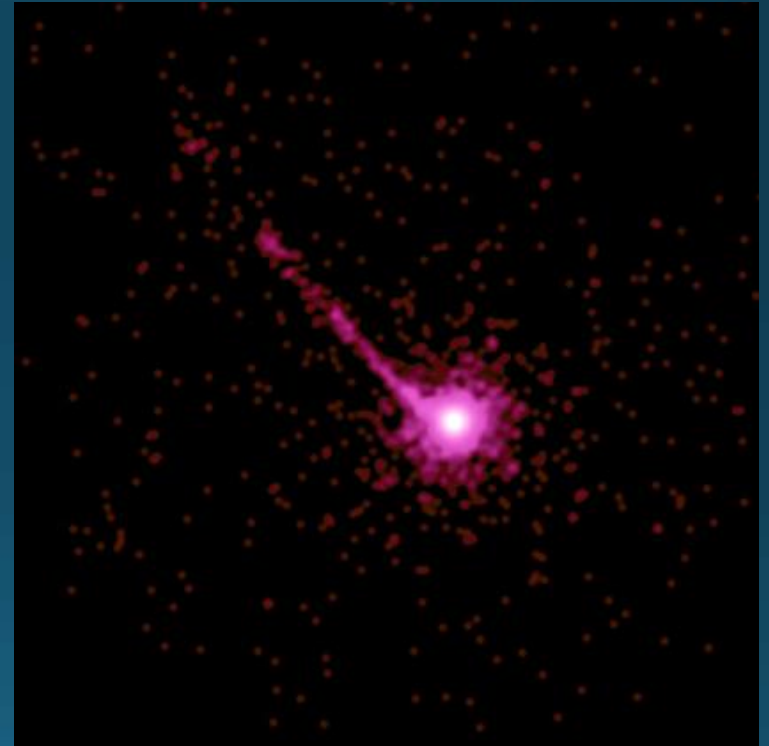
Quasars

- Originally:
 - Quasar = “Quasi-stellar radio source” (radio-loud)
 - QSO = “Quasi-stellar object” (radio-quiet)
- Today: Quasar = Both types



Quasars

- Most luminous of the non-transient objects in the Universe: $M_B < -23$
- Radio-quiet quasars >10 times more common than radio loud ones
- Both broad and narrow lines



X-ray quasar with jet

Seyfert Galaxies

Galaxy NGC 7742



Hubble
Heritage

Seyfert Galaxies

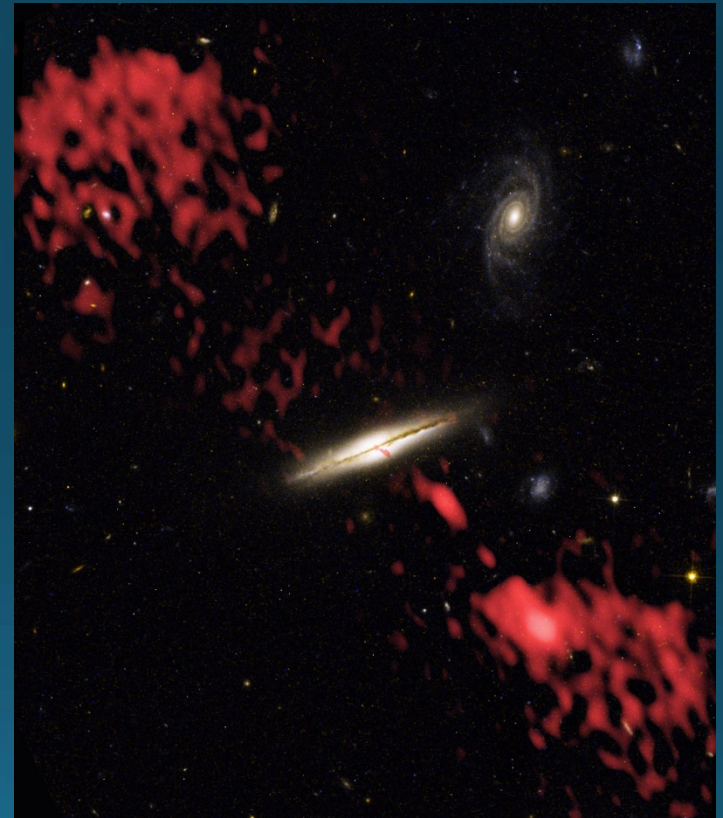
- “Low-luminosity quasars”
- Almost always in S- or So-galaxies
- Seyfert 1 nuclei
 - Broad lines (allowed) & Narrow lines (forbidden)
 - High optical luminosity
- Seyfert 2 nuclei
 - Narrow lines only, but with wings
 - Low optical luminosity

LINERs

- LINER = Low Ionization Nuclear Emission Line Region
- Low luminosities (lower than Seyfert 2)
- Exhibit lines which do not require very energetic power sources – hot stars sufficient
- Many LINERs are probably starbursts, not genuine AGN

Radio Galaxies

- Milky Way: 10^{30} W in radio
- Radio galaxies $\geq 10^{34}$ W in radio
- Lobes and hot spots
- Always elliptical galaxies



Blazars

- The most rapid and large variations among AGN
- Originally:
 - BL Lac (very weak emission lines)
 - OVV = Optically violent variable (strong emission lines)
- Today: Blazar = BL Lac & OVVs
- *Appear* to be the most luminous objects in the Universe, but this is due to beaming
- Often completely featureless spectrum
 - Emission-lines weak or absent

Intermission:

What sort of AGN is this?



Optical spectrum reveals lots of narrow emission lines

Intermission:

What sort of AGN is this?



The Unification Model

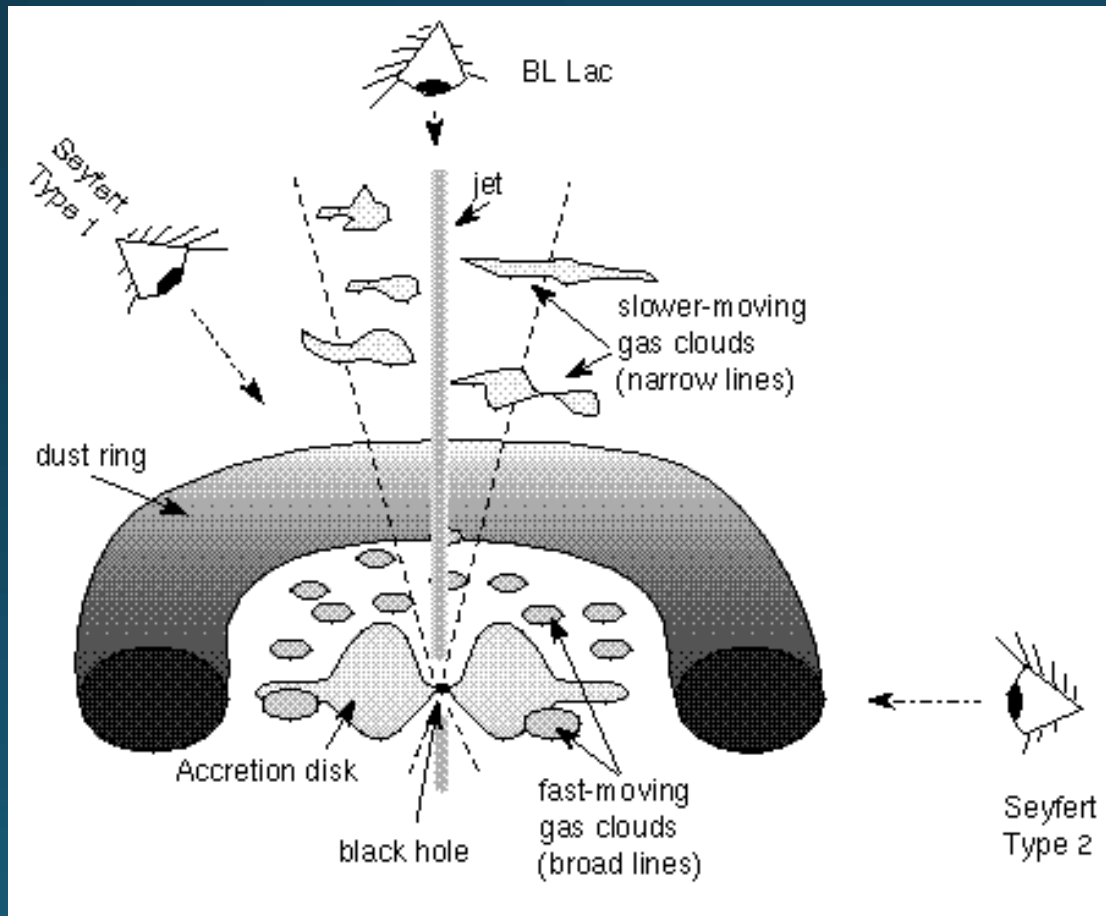
Quasar



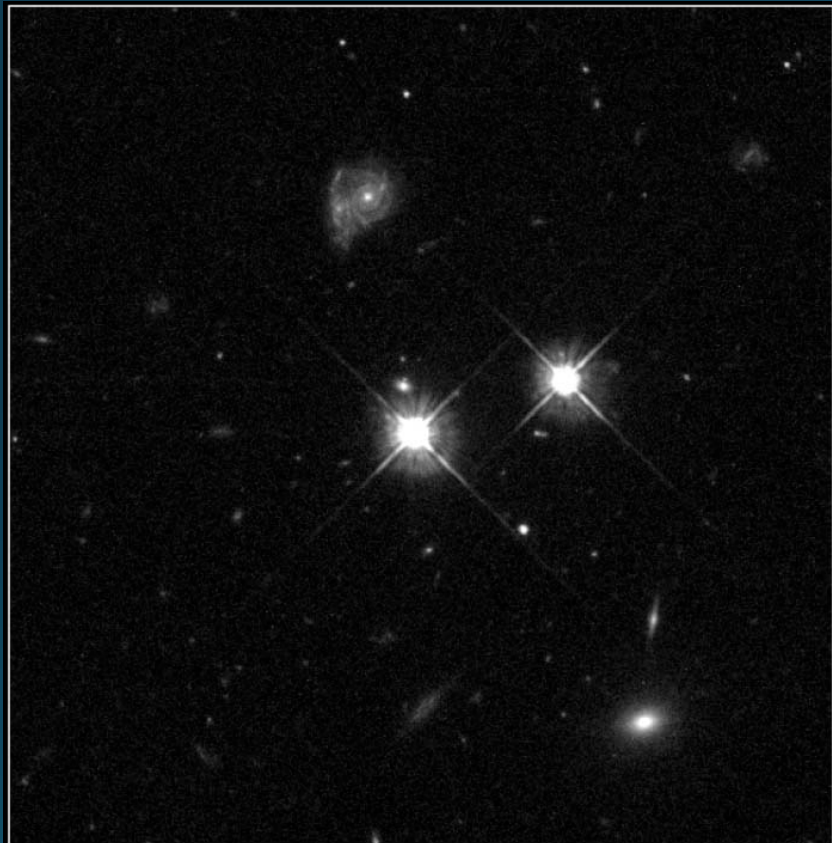
Radio galaxy



Radio galaxy



Quasar Host Galaxies

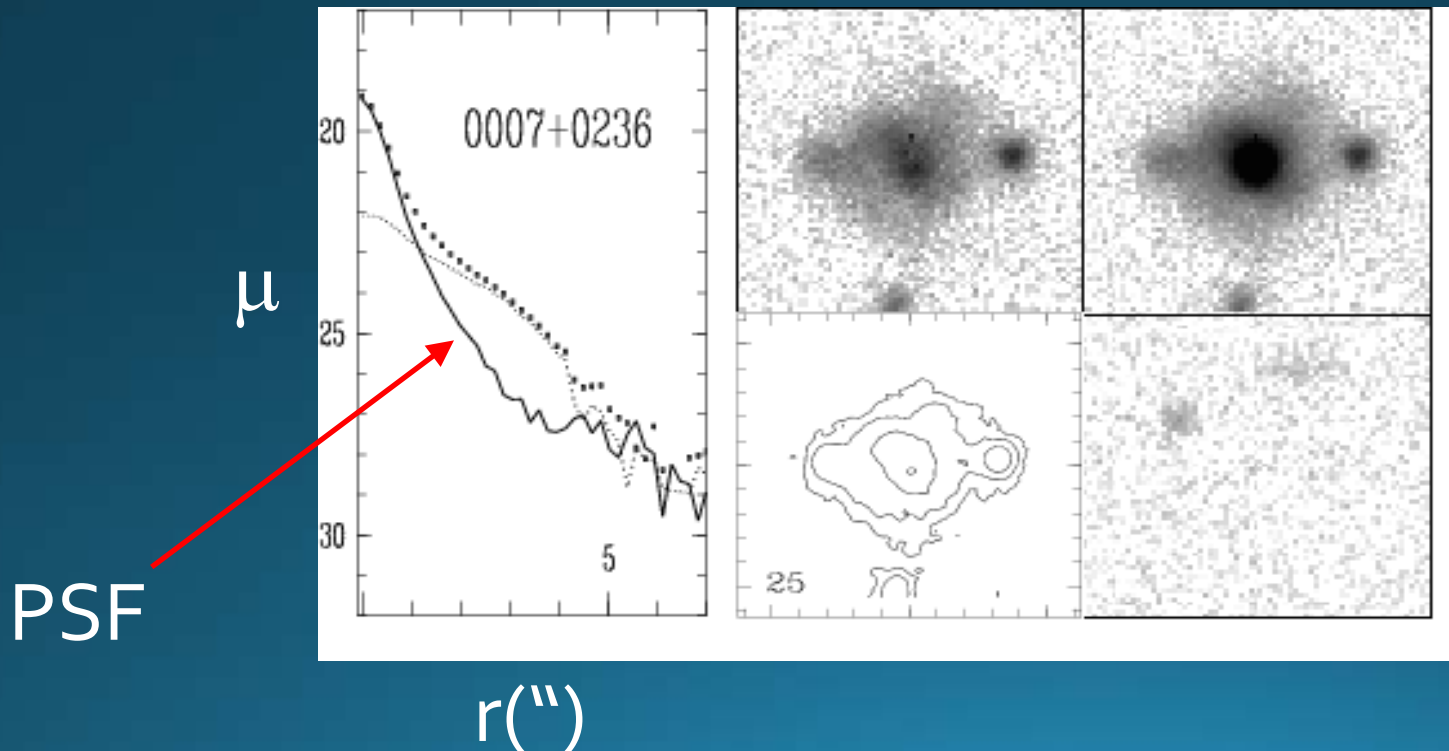


HST's 100,000th Observation HST • WFPC2
PRC96-25 • ST ScI OPO • July 10, 1996 • C. Steidel (CalTech), NASA

- The AGN of a quasars typically outshines its host galaxy
- To study the host galaxy, one utilizes the fact that the AGN is a point source whereas the host is an extended object

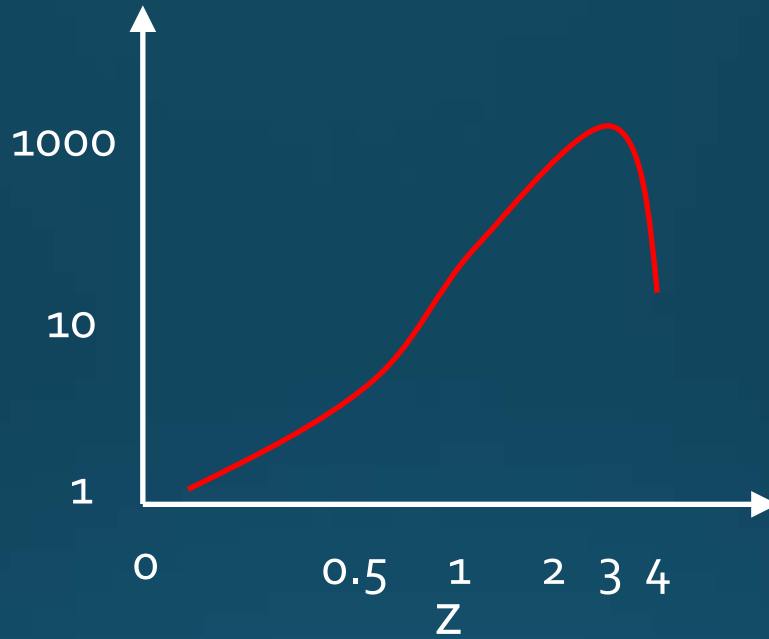
Quasar Host Galaxies

- Point spread function (PSF): Describes how the light of a perfect point source is distributed on the detector (CCD) because of telescope imperfections, diffraction etc.



Cosmological Evolution

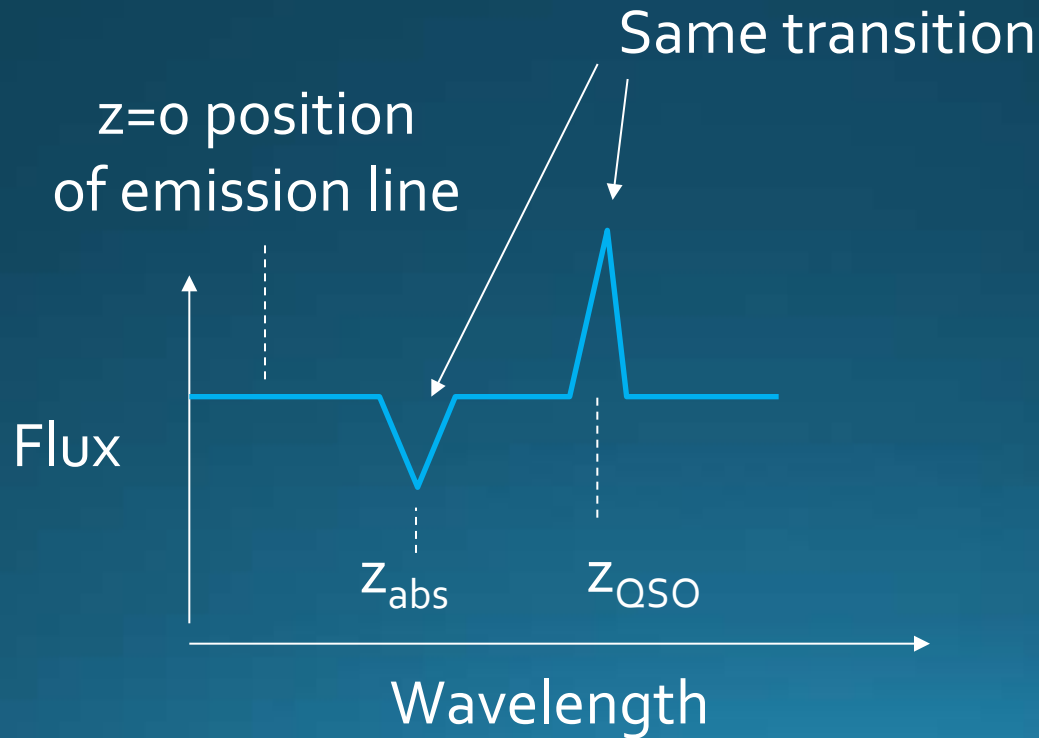
Quasars Gpc^{-3}



- At $z \approx 3$, bright quasars appear ~ 1000 times more common than today
- Quasar activity drops at even higher redshifts
- The highest-redshift quasar detected so far has $z \approx 7$
- Very challenging to build a SMBH by $z=7$ (less than 1 Gyr after the Big Bang) – exotic formation channel required?

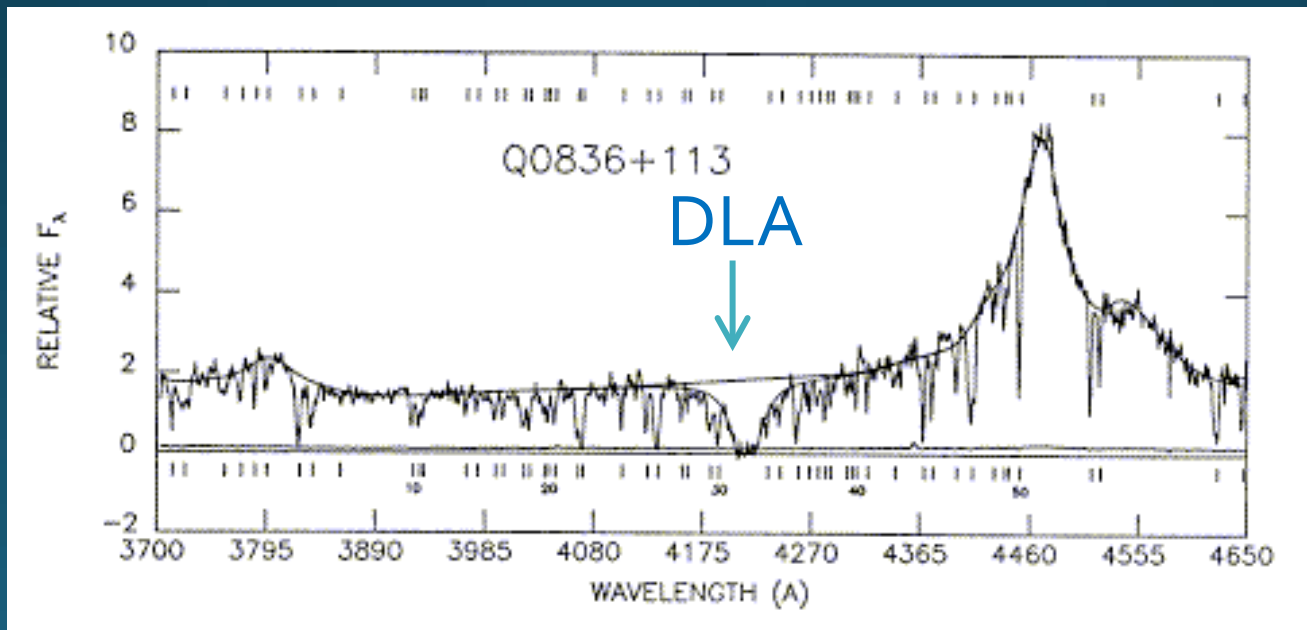
Quasar Absorption Systems

Quasar



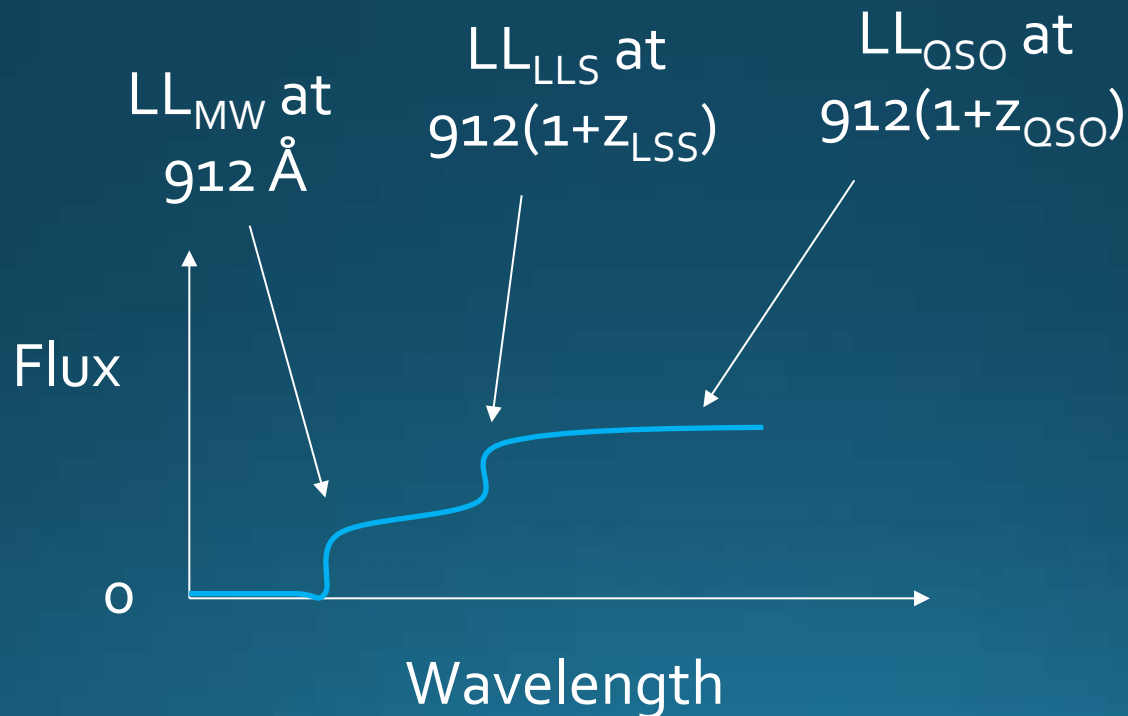
Damped Lyman-Alpha Clouds

- Dense, neutral gas $N(\text{HI}) \approx 2 \times 10^{20} \text{ cm}^{-2}$
- Often interpreted as gas associated with intervening galaxies
- Many DLAs are also Lyman-limit systems

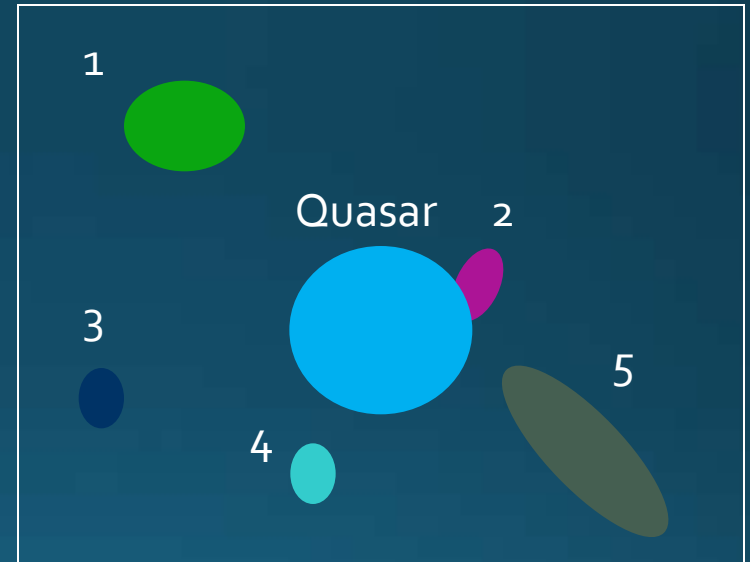
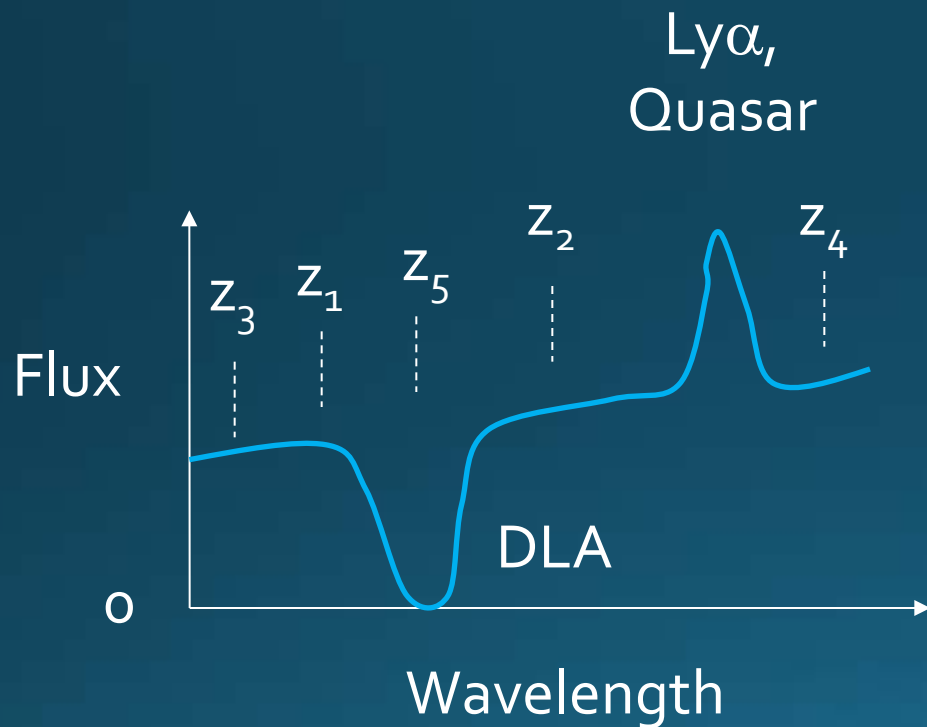


Lyman-Limit Systems

- Dense, neutral gas
- Absorbs hydrogen-ionizing radiation at $\lambda < 912(1+z_{\text{LLS}}) \text{ \AA}$

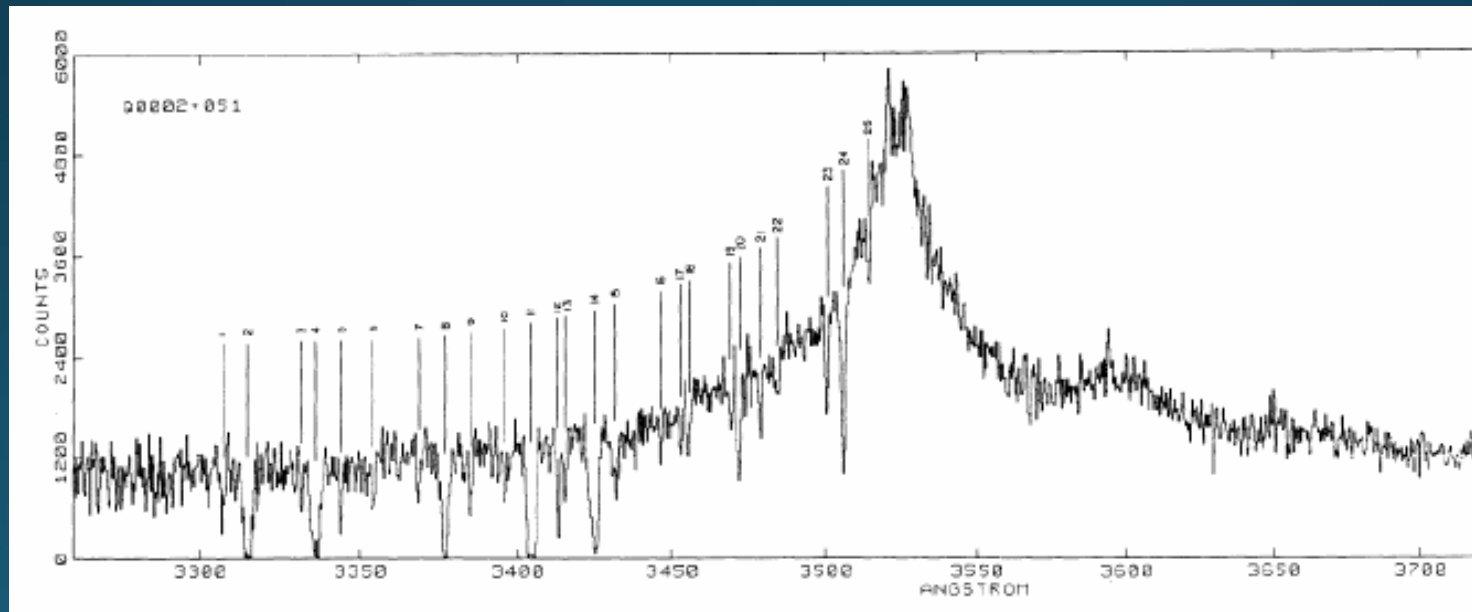


Identifying the Absorber



The Lyman-Alpha Forest

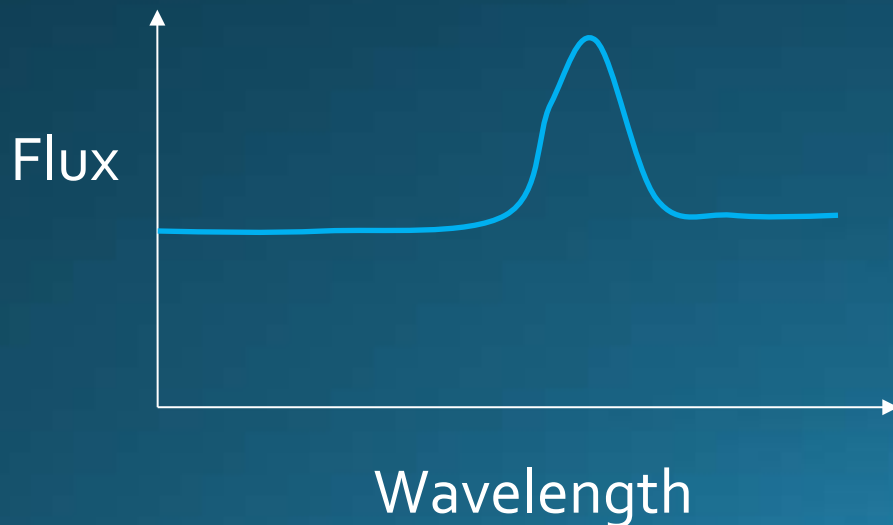
- Low-density, very extended clouds in the intergalactic medium
- Proximity effect:
 - Ly α -forest thinner at $z_{\text{abs}} \approx z_{\text{QSO}}$
 - Indicates that clouds close to the quasar are photoionized by it



The Gunn-Peterson Test

- If the Universe (the intergalactic medium, IGM) is neutral at z_{QSO} , then a strong absorption feature blueward of $\text{Ly}\alpha$ in quasars should appear – *the Gunn-Peterson trough*.
- This does indeed appear – at $z_{\text{QSO}} \approx 6$, indicating that the transition from an neutral to ionized IGM takes place at around this redshift

Universe ionized at z_{QSO}



Universe neutral at z_{QSO}

