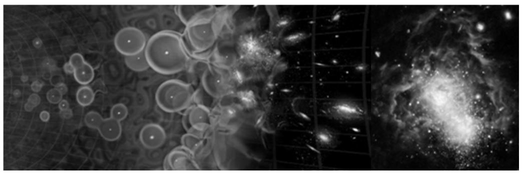
Physics of Galaxies 2017
10 credits
Lecture 8: The High-Redshift Universe



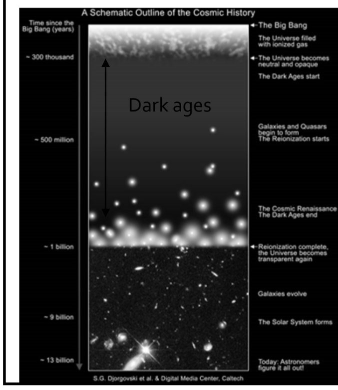
Outline: Part I

- The first stars and galaxies
  - End of the dark ages
  - Pop III stars
  - Dark stars
  - First galaxies

Outline: Part II

- Finding high-redshift objects
  - Deep fields
  - Gravitational lensing
  - Dropout techniques
  - Lyα searches
- Future prospects

The end of the dark ages

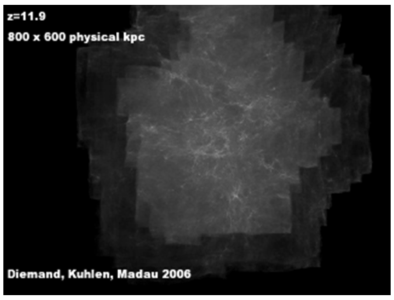


First stars
$z \approx 20\text{-}30$
$t_{Univ} \approx 100\text{-}200$ Myr

First galaxies
$z \approx 10\text{-}15$
$t_{Univ} \approx 300\text{-}500$ Myr

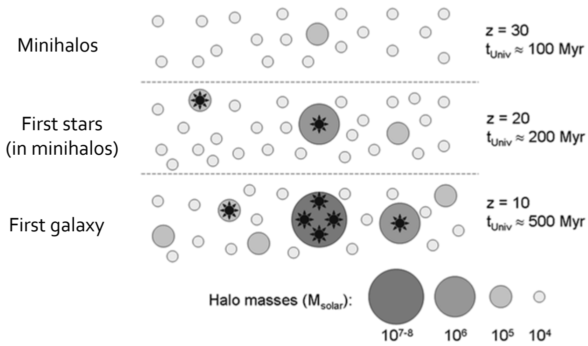
Current observational limit: HST and 8-10 m telescopes on the ground can detect light sources up to $z \approx 11$

Merging cold dark matter halos



Formation of a ~10^{12} M_{solar} dark matter halo
Simulation runs from $z \approx 12$ to 0 ($t_{Univ} \approx 0.25$ to 13.7 Gyr)

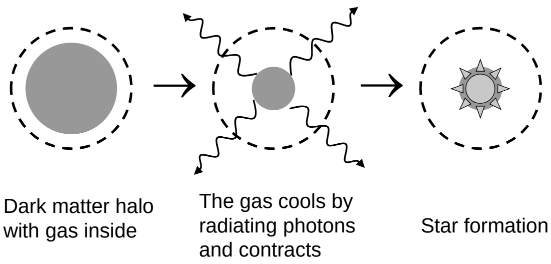
Structure formation



Population I, II and III

- Population I: Metal-rich stars
  Example: Stars in the Milky Way disk
- Population II: Metal-poor stars
  Example: Stars in the Stellar halo of the Milky Way
- Population III: (Almost) Metal-free stars
  Example: Stars forming in minihalos at $z \approx 20$

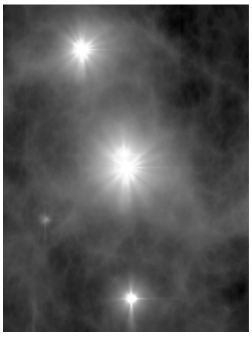
Star formation in dark matter halos



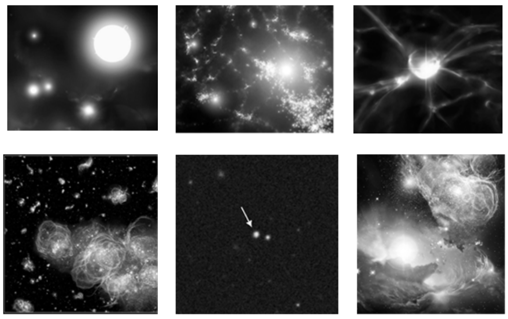
Problem: Low metallicity at high redshifts →
Lack of efficient coolants

Population III stars

- These stars will be *very massive*, hot and short-lived.
- Mass range 10^1-10^3 Msolar (but predictions still shaky)
- The first ones are expected in minihalos – prior to the formation of the first galaxies.
- Feedback → *Only a few stars (maybe just one) per minihalo*

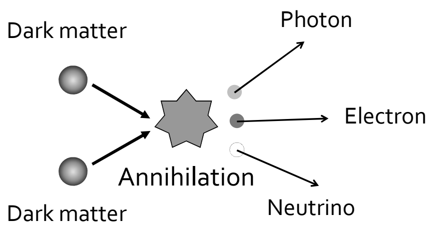


Intermission: The first stars(?)

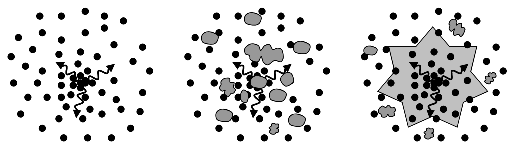


Normal star ≈ hydrogen bomb

Dark matter annihilation



Dark stars

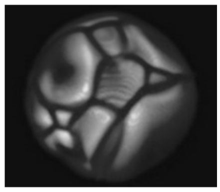


WIMP annihilation in centre of CDM halo

Gas cools and falls into the centre

Star fueled by WIMP annihilation rather than hydrogen fusion

Dark star properties



- Conventional Pop III stars
  - Teff ~ 50 000-100 000 K
  - M ~ 10^1-10^3 Msolar
  - Lifetime $\tau \sim 10^6$-10^7 yr
- Pop III dark stars
  - Teff ≈ 4000-50000 K   Cooler!
  - M ~ 10^2-10^7 Msolar   More massive???
  - Lifetime $\tau \sim 10^6$-10^{10} yr   More long-lived???

Problem: Still no consensus on likely masses or life times of dark stars

The sizes of primordial stars I

The Sun

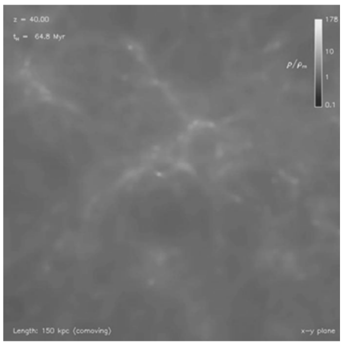
Vanilla population III star

The sizes of primordial stars II

The Sun

Supermassive dark star

Formation of the first galaxies



Greif et al. 08

Formation of a ~ 10^7 M_{solar} dark matter halo

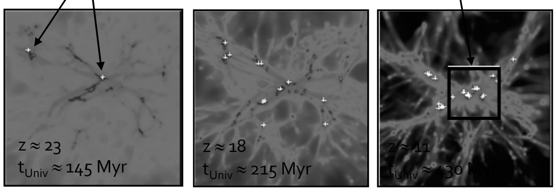
Simulation runs from z ≈ 40 to 11 (t_{Univ} ≈ 65 to 430 Myr)

Star formation inside and outside the first galaxies

Star formation in minihalos

Minihalo mergers and further star formation

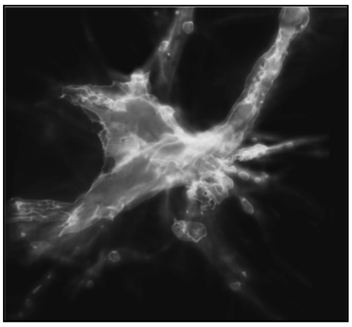
Object qualifies as a *galaxy*



Greif et al. 08

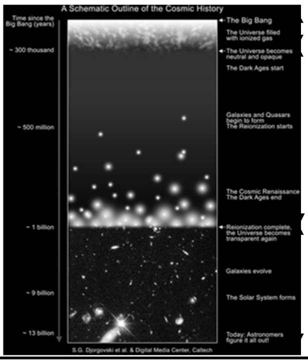
Gas density shapshots

A galaxy is born (at $z \approx 10$)



Greif et al. 08

Reionization



Intergalactic medium

Ionized

Neutral

CMBR (Planck) $\rightarrow z_{reion} \approx 8$

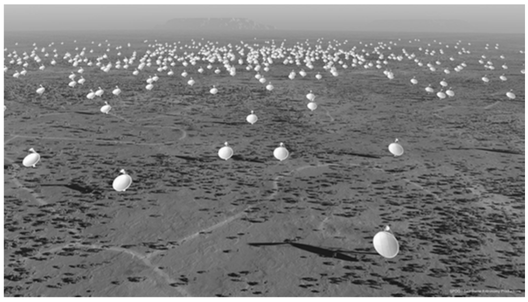
Lyα absorption in quasars $\rightarrow z_{reion} > 6$

Reionized

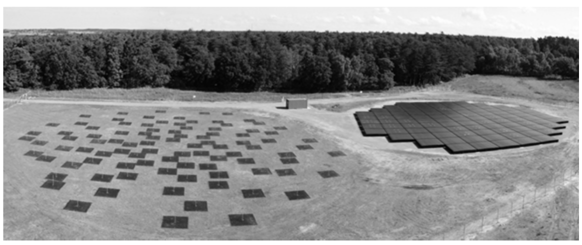
What caused reionization?

- Population III stars in minihalos?
- *High-redshift galaxies?* ← Popular scenario
- Accreting black holes?
- Decay of exotic particles?

Intermission: Name the telescope!



Intermission: Name the telescope!



Intermission: Name the telescope!

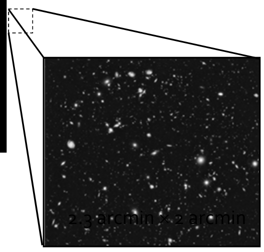
How to find and study high-redshift galaxies

Imaging strategies

- Deep field-style observations
  - Very long exposures of single patch (devoid of bright foreground objects) in the sky
- Cluster-lensing observations
  - Hunt for gravitationally lensed background objects in relatively short exposures (few hours per filter) of a low-z galaxy cluster

The Hubble Extreme Deep Field

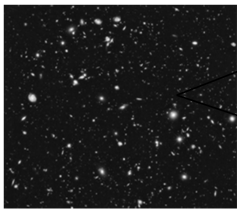
otal exposure time: 23 days
million seconds)

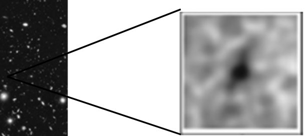


The Hubble Extreme Deep Field

Example of one of the most distant galaxy candidates so far

2.4 arcsec x 2.4 arcsec



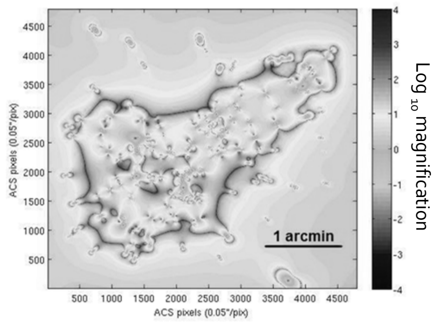


Bouwens et al. (2010)
z ≈ 10 candidate

Cluster lensing I

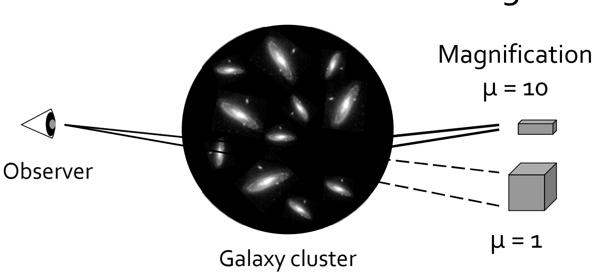
Galaxy cluster at z≈0.5

Cluster lensing II



Magnification map

Pros and Cons of Cluster Lensing



+ Background sources appear brighter by a factor μ
- The volume probed becomes smaller by a factor μ

Bottom line: Lensed survey fields can be superior for sources that are *very faint*, *not too rare* and *not too highly clustered*

Intermission: Why are redshift records important?

Notably distant objects [edit]

1 Gly = 1 billion light-years.

Most distant astronomical objects with spectroscopic redshift determinations

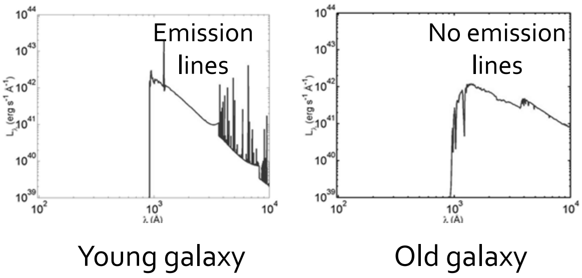
Name	Redshift (z)	Light travel distance§ (Gly)[1]	Type	Notes
GN-z11	z=11.09	13.39	Galaxy	Confirmed galaxy[2]
EGSY8p7	z=8.68	13.23	Galaxy	Confirmed galaxy[3]
GRB 090423	z=8.2	13.18	Gamma-ray burst	[4][5]
EGS-zs8-1	z=7.73	13.13	Galaxy	Confirmed galaxy[6]
z8 GND 5296	z=7.51	13.10	Galaxy	Confirmed galaxy[7][8]
A1689-zD1	z=7.5	13.10	Galaxy	Galaxy[9]
SXDF-NB1006-2	z=7.215	13.07	Galaxy	Galaxy[10][11]
GN-108036	z=7.213	13.07	Galaxy	Galaxy[11][12]
BDF-3299	z=7.109	13.05	Galaxy	[13]
ULAS J1120+0641	z=7.085	13.05	Quasar	[14]
A1703 zD6	z=7.045	13.04	Galaxy	[11]

Selecting high-z galaxy candidates

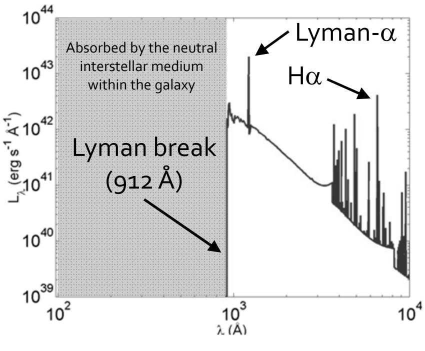
Two techniques:

- Dropout selection
  - Crude redshift estimator (Δz≈1.0)
  - But works well for all high-z, star-forming galaxies
- Lyman-alpha surveys
  - High-precision redshift estimation (Δz≈0.1)
  - But doesn't work well at z>6
  - And not all galaxies are Lyα-emitters

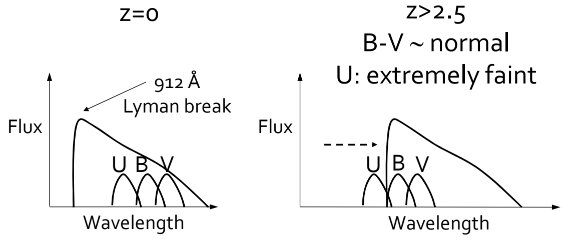
The UV/optical spectra of galaxies I



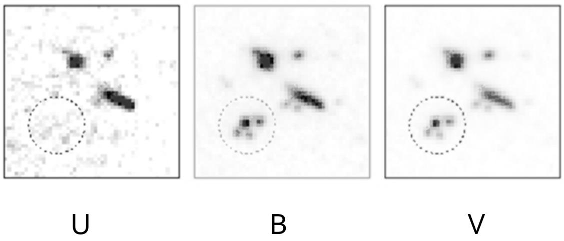
The UV/optical spectra of galaxies



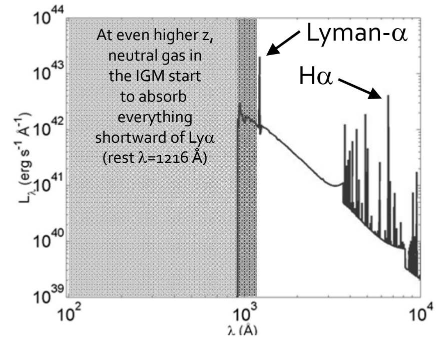
Drop-out techniques: Lyman-Break Galaxies



Drop-out techniques: Lyman-Break Galaxies

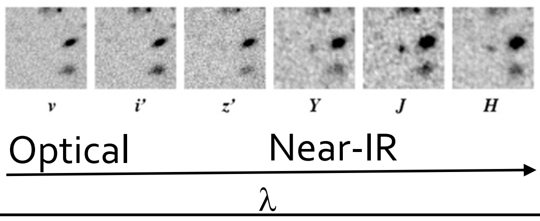


Reionization-epoch galaxies



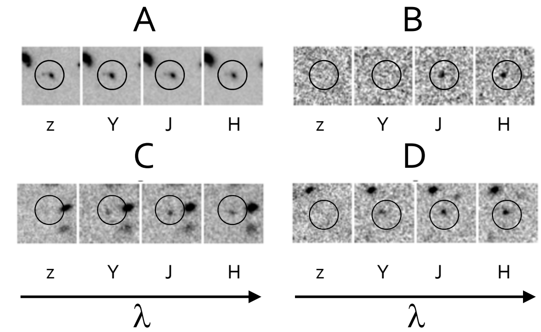
Drop-out techniques: z>6 objects

Eventually, the break shifts into the near-IR. Example: z-band dropout (z≈6.5)



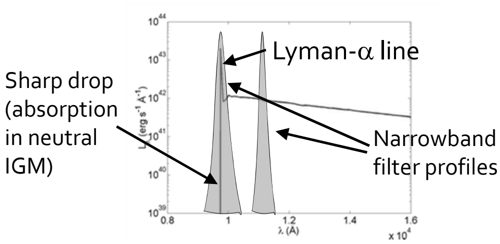
Intermission:

Which of these drop-out candidates is likely to have the highest redshift?



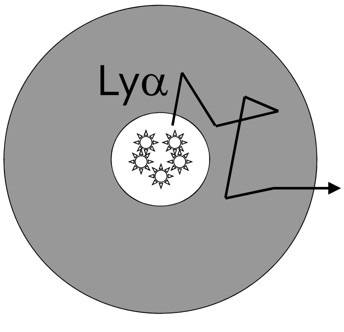
Lyman-alpha surveys

- Potentially the brightest line in rest frame UV/optical
- Two narrowband images (covering continuum and line) required for survey of redshift range (Δz~0.1)



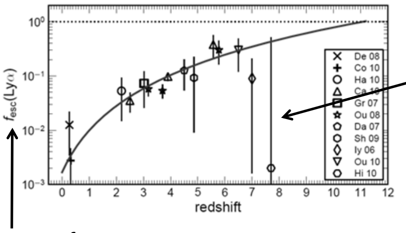
Lyman-α at z=7

Problem I: Lyman-α notoriously difficult to predict



- Lyα resonant line → random walk through neutral interstellar medium
- Many Lyα photons destroyed by dust before emerging
- Lyα flux ranges from low to very high

Problem II: Lyman-α largely absorbed in the neutral intergalactic medium at z>6



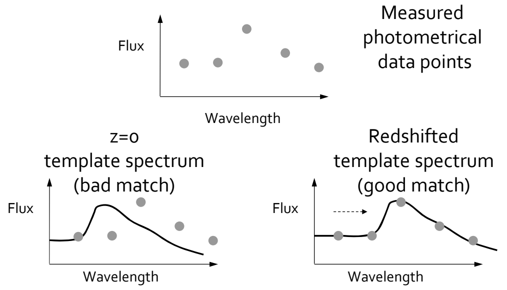
Abrupt drop → Lyα not good way to find z>6 galaxies (but may be good way to probe reionization)

Fraction of Lyα photons reaching the observer

Hayes et al. 11

Photometric redshifts

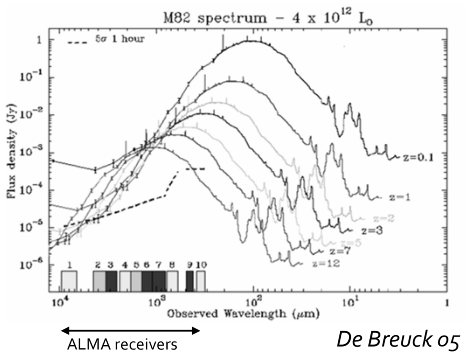
- Estimate the galaxy type (morphological) and assume that the galaxy is identical to some template (often an average over many galaxy spectra of similar type)



New telescope for high-z studies: ALMA

Atacama Large Millimeter/ submillimeter Array (ALMA): An array of seventy 12-m antennas operating @ 200-10000 μm (sub-mm)

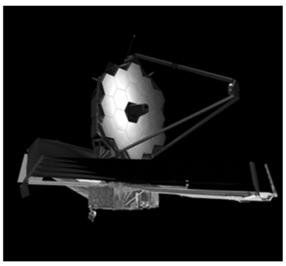
Can be used to search for dust emission and emission lines like [CII] @ 158 μm and [OIII] @88 μm (rest-frame) from z>6 galaxies



De Breuck 05

Dust continuum flux drops slowly with z (if no source evolution).

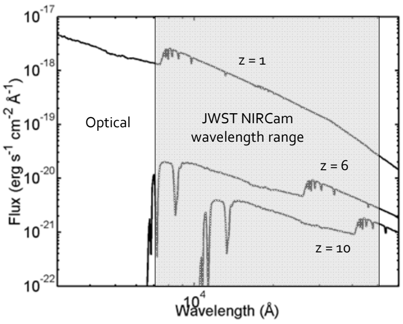
Future prospects: JWST



James Webb Space Telescope
'The first light machine'
To be launched by NASA / ESA / CSA in 2018

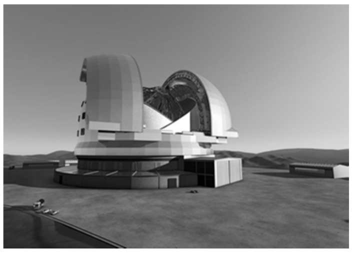
6.5 m mirror
Observations @ 0.6-29 μm
Useful for:
Galaxies up to z ≈ 15
Pop III supernovae

Why infrared?



Zackrisson et al. (2001) model

Future prospects: E-ELT



39 m European Extremely Large Telescope (E-ELT) estimated to be completed in 2024.