

# Stokes *IQUV* magnetic Doppler imaging of Ap stars – I. ESPaDOnS and NARVAL observations<sup>★</sup>

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## ABSTRACT

In this paper we describe and evaluate new spectral line polarization observations obtained with the goal of mapping the surfaces of magnetic Ap stars in great detail. One hundred complete or partial Stokes *IQUV* sequences, corresponding to 297 individual polarized spectra, have been obtained for seven bright Ap stars using the Échelle SpectroPolarimetric Device for the Observation of Stars (ESPaDOnS) and NARVAL high-resolution spectropolarimeters. The targets span a range of masses from approximately 1.8 to 3.4  $M_{\odot}$ , a range of rotation periods from 2.56 to 6.80 d and a range of maximum longitudinal magnetic field strengths from 0.3 to over 4 kG. For three of the seven stars, we have obtained dense phase coverage sampling the entire rotational cycle. These data sets are suitable for immediate magnetic and chemical abundance surface mapping using magnetic Doppler imaging. For the remaining four stars, partial phase coverage has been obtained, and additional observations will be required in order to map the surfaces of these stars. The median signal-to-noise ratio of the reduced observations is over 700 per  $1.8 \text{ km s}^{-1}$  pixel. Spectra of all stars show Stokes *V* Zeeman signatures in essentially all individual lines, and most stars show clear Stokes *QU* signatures in many individual spectral lines. The observations provide a vastly improved data set compared to previous generations of observations in terms of signal-to-noise ratio, resolving power and measurement uncertainties. Measurement of the longitudinal magnetic field demonstrates that the data are internally consistent within computed uncertainties typically at the 50–100 $\sigma$  level. Data are also shown to be in excellent agreement with published observations and in qualitative agreement with the predictions of published surface structure models. In addition to providing the foundation for the next generation of surface maps of Ap stars, this study establishes the performance and stability of the ESPaDOnS and NARVAL high-resolution spectropolarimeters during the period 2006–2010.

**Key words:** stars: chemically peculiar – stars: magnetic field.

<sup>★</sup> Based on observations obtained at the Canada–France–Hawaii Telescope (CFHT) which is operated by the National Research Council of Canada, the Institut National des Sciences de l'Univers of the Centre National de la Recherche Scientifique of France and the University of Hawaii. Also based on observations obtained at the Bernard Lyot Telescope (TBL, Pic du Midi, France) of the Midi-Pyrénées Observatory, which is operated by the Institut National des Sciences de l'Univers of the Centre National de la Recherche Scientifique of France.

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## 1 INTRODUCTION

The classification Ap identifies a (main-sequence) A- or B-type star which displays peculiar chemical abundances, usually combined with an observable magnetic field. Although other classes of chemically peculiar stars exist (e.g. Am stars, Hg-Mn stars, He-weak stars), these stars have been demonstrated to lack strong, organized magnetic fields at their surfaces (e.g. Shorlin et al. 2002; Wade et al. 2006; Makaganiuk et al. 2011). Ap stars appear to be the only class of middle main-sequence stars for which, in all cases, an observable magnetic field is present (Aurière et al. 2007).

Since their discovery by Babcock in 1947, the magnetic fields of Ap stars have been established through observation to have important global dipole components with polar strengths ranging from hundreds to tens of thousands of gauss. The symmetry axis of the dipole component is almost always significantly tilted relative to the stellar rotation axis. In addition, Ap stars generally spin much more slowly than non-peculiar stars of similar masses (Stepień 2000), and as they spin they exhibit line profile variations attributed to rotational modulation of patchy, non-axisymmetric lateral and vertical distributions of chemical abundance in their photospheres. The distributions of abundance vary significantly from element to element: some are distributed relatively uniformly, while others show strong contrast; some are distributed in relatively simple patterns, while others show complex distributions. While it is generally accepted that the fundamental mechanism responsible for the chemical peculiarities is microscopic chemical diffusion (as described by Michaud 1970), the origin of chemical patchiness, and the relationship to the magnetic field, is poorly understood.

The earliest studies of the magnetic field geometries of Ap stars interpreted the rotational variations of their longitudinal magnetic fields in the context of Stibbs' oblique rotator model assuming a simple magnetic dipole field (e.g. Babcock 1947, 1951; Stibbs 1950). However, with the acquisition of increasingly sophisticated diagnostic data (the mean surface field, or mean field modulus, and high-resolution line profiles), it became clear that the large-scale field topologies exhibited important departures from the simple dipolar model (e.g. Preston & Sturch 1967; Preston 1969a,b, 1970; Landstreet 1970, 1988; Landstreet et al. 1989).

Leroy and collaborators (Landolfi et al. 1993; Leroy, Landolfi & Landi Degl'Innocenti 1993; Leroy, Landstreet & Bagnulo 1994; Bagnulo et al. 1995; Leroy et al. 1995a,b; Leroy, Landolfi & Landi degl'Innocenti 1996; Wade et al. 1996) systematically studied Ap stars using broad-band linear polarization measurements and models, constraining the transverse component of the magnetic field. Importantly, they found that differences between the observed linear polarization variations and those predicted by the simple dipole model could not be fully explained by abundance inhomogeneities alone (Leroy et al. 1996). With this work, they established a modified dipolar model with a trend towards an outward expansion of the field lines over some parts of the magnetic equator and showed the potential of linear polarization for diagnosing small-scale structure of the magnetic fields of Ap stars. Thus, the observations and modelling undertaken during the latter half of the 20th century allowed progress from a simple view of the magnetic fields of Ap stars to a relatively sophisticated picture in which fields were known to show both global-scale and local-scale departures from a simple dipole.

Leroy et al. (1996) commented that high-resolution spectropolarimetry represented the next step in furthering the study of the magnetic field geometry of Ap stars. Four years later, Wade et al. (2000a) published the first compendium of phase-resolved high-resolution spectropolarimetric observations of Ap stars in both circular and linear polarizations. Using the MuSiCoS spectropolarimeter,  $R = 35\,000$  Stokes  $IQUV$  spectra with a median signal-to-noise ratio (S/N) of 300 (per  $4.6\text{ km s}^{-1}$  pixel) were obtained for 14 Ap stars. While the quality of the spectra was sufficiently good to show the shape and phase variation of all Stokes parameters in mean least-squares deconvolved (LSD) line profiles, measurement in individual spectral lines was restricted to a few particularly strong lines, principally those of the Fe II multiplet 42. Nevertheless, the Stokes profiles of 53 Cam were used by Bagnulo et al. (2001) and Kochukhov et al. (2004) to evaluate published magnetic models developed based on less sophisticated data. These authors found

that models based on so-called magnetic observables (e.g. Bagnulo et al. 2000) led to the derivation of surface magnetic field characteristics that were not consistent with the detailed Stokes profiles, and that both circular and especially linear polarization profiles were required for realistic reconstruction of the field.

Following these conclusions, Kochukhov et al. (2002) for  $\alpha^2$  CVn and Luftinger et al. (2010) for HD 24712 employed the new magnetic Doppler imaging (MDI) technique, described by Piskunov & Kochukhov (2002) and Kochukhov & Piskunov (2002), to construct high-resolution maps of the surface vector magnetic field using Stokes  $IV$  observations and by preferring a global low-order multipolar field structure. Maps using linear polarization profiles (Stokes  $Q$  and  $U$ ) were constructed by Kochukhov et al. (2004) and Kochukhov & Wade (2010) for the Ap stars 53 Cam and  $\alpha^2$  CVn. These maps were distinguished from earlier models in that they were computed directly from the observed polarized line profiles, making no a priori assumptions regarding the large-scale or small-scale topology of the field. The MDI surface magnetic field maps of both stars revealed that their magnetic topologies depart significantly from low-order multipoles. In particular, both studies concluded that while the global topology of the magnetic field was reasonably smooth, the strength of the field was quite patchy, indicating complex structure on relatively small scales. Simultaneous mapping of the distributions of the surface chemical abundances of several elements was also performed, allowing a comparison between the local field properties and local photospheric chemistry.

It is important to note that the observational material used in the MDI studies of 53 Cam and  $\alpha^2$  CVn represented the best data sets obtained from several years of MuSiCoS observations. In those spectra the uniquely valuable Stokes  $Q$  and  $U$  Zeeman signatures were only clearly detectable in three strong lines, with a significance (i.e. amplitude divided by error bar) of 5 or less. The relatively low S/N and resolving power achievable with the MuSiCoS instrument led to some ambiguity in the field reconstruction and limited the useful sample of stars to those with bright apparent magnitudes, strong fields and sharp lines. Because MDI exploits the indirect resolution of the stellar disc due to stellar rotation, those stars best suited to reconstruction (relatively rapidly rotating stars, with consequentially weaker Stokes profiles) were inaccessible to MuSiCoS. As a result, only an extremely limited range of stellar properties (rotation, mass, temperature, magnetic field, etc.) which may influence the phenomena of interest could be studied using the MuSiCoS data.

To address outstanding questions surrounding the detailed magnetic structure of Ap stars and the effect of the magnetic field on atmospheric chemical transport processes, we have acquired new higher resolution and S/N Stokes  $IQUV$  spectra of a small sample of well-studied magnetic Ap stars using the new generation of high-resolution spectropolarimeters. In this paper we describe the observations obtained. We demonstrate the stability of the instrumentation during the five years of observation by evaluating the internal and external agreement of the data. We illustrate the quality of the observed Stokes profiles, comparing with MuSiCoS results and demonstrating that they represent a qualitative step forward in our ability to diagnose the magnetic structure of Ap stars.

## 2 TARGETS

Targets selected for this study are bright Ap stars demonstrated to exhibit strong Stokes profiles by Wade et al. (2000a). We attempted to select targets spanning a large range of stellar physical properties as well as field strengths and geometries. Target stars were generally

**Table 1.** Stars discussed in this paper, along with ancillary data: HD designation, other name,  $V$  magnitude, spectral type, projected rotational velocity, rotational period, radius, inclination, the number of longitudinal field measurements obtained, the median longitudinal field uncertainty obtained, the number of net linear polarization measurements obtained. Projected rotational velocities ( $v \sin i$ ) have been measured in this study, as discussed in Section 5. The periods were obtained from various sources as described in Section 6. Masses are those reported by Kochukhov & Bagnulo (2006), and the average longitudinal magnetic fields are taken from the catalogue of stellar effective magnetic fields (Bychkov, Bychkova & Madej 2003), using values obtained from LSD profiles if available. Temperatures are those reported by Kochukhov & Bagnulo (2006), and stellar radii as reported in Leone, Catanzaro & Catalano (2000) for HD 32633, in Pasinetti Fracassini et al. (2001) for HD 40312, in Kochukhov & Wade (2010) for HD 112413 and in Wade (1997) for the remaining stars. Inclinations are taken from Stepień (1989) for HD 32633, from Rice, Holmgren & Bohlender (2004) for HD 40312, from Kochukhov & Wade (2010) for HD 112413 and from Leroy et al. (1996) for the remaining stars.

HD	Name	$V$	Spectral type	$T_{\text{eff}}$ (K)	$M$ ( $M_{\odot}$ )	$v \sin i$ ( $\text{km s}^{-1}$ )	$P_{\text{rot}}$ (d)	$R$ ( $R_{\odot}$ )	$i$ ( $^{\circ}$ )	# $B_{\ell}$	# net $Q/U$	Median $\sigma_B$ (G)	$\langle B_{\ell} \rangle^{\text{rms}}$ (kG)
4778		6.1	A0p	10 000	2.29	$36 \pm 2$	$2^{\text{d}}561\,71$	1.83	70	7	13	20	1.02
32633		7.1	B9p	12 800	3.10	$19 \pm 2$	$6^{\text{d}}430\,00$	2.4	76	21	40	30	2.85
40312	$\theta$ Aur	2.6	A0p	10 200	3.41	$53 \pm 1$	$3^{\text{d}}618\,60$	2.2	54	7	14	14	0.21
62140	49 Cam	6.5	F0p	7700	1.77	$24 \pm 2$	$4^{\text{d}}286\,79$	1.95	90	19	38	21	1.11
71866		6.8	A0p	8800	2.24	$15 \pm 2$	$6^{\text{d}}800\,22$	2.0	110	14	28	27	1.54
112413	$\alpha^2$ CVn	2.9	A0p	11 500	2.98	$17 \pm 1$	$5^{\text{d}}469\,39$	2.49	120	27	48	55	0.53
118022	78 Vir	4.9	A1p	9100	2.16	$13 \pm 1$	$3^{\text{d}}722\,0$	2.1	25	5	10	12	0.63

required to have the following characteristics to be suitable for this study.

(i) Well-determined rotation periods – This study requires phase-resolved time-series observations. All targets must therefore have well-determined, unambiguous rotation periods so that each phase is observed correctly.

(ii) Suitable projected rotational velocity – The projected rotational velocity must be neither too rapid (rapid rotators  $>50 \text{ km s}^{-1}$  typically have shallow Stokes profiles which are challenging to detect and interpret) nor too slow ( $<2 \text{ km s}^{-1}$  whilst MDI can be applied to stars with such small rotational velocities, the advantage of Doppler tomography is lost).

(iii) Strong magnetic fields and variability – Because MDI relies on both the shape and variation of line profiles to determine the geographic location of magnetic and chemical features, target stars should display strong and variable Stokes profiles.

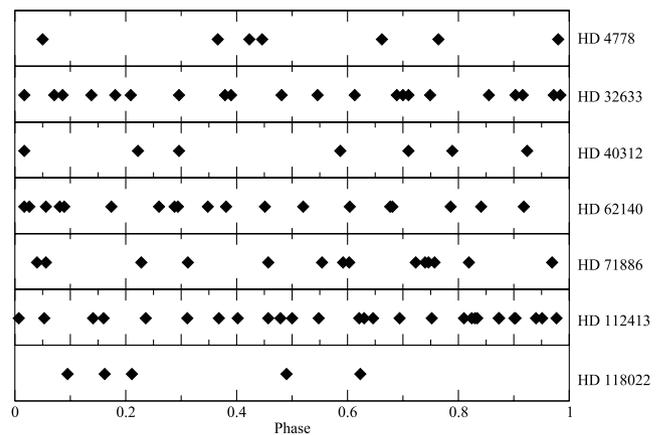
Ultimately, seven targets were selected for monitoring. The target list is shown in Table 1.

### 3 OBSERVATIONS OBTAINED WITH ESPaDOnS AND NARVAL

In total 48 sequences were obtained with ESPaDOnS and 52 with NARVAL, the achieved phase coverage of the stellar targets is illustrated in Fig. 1.

Both the Échelle Spectropolarimetric Device for the Observation of Stars (ESPaDOnS) and NARVAL instruments consist of a bench-mounted cross-dispersed échelle spectrograph, fibre-fed from a Cassegrain-mounted polarimeter unit. These instruments are designed to overcome the limitations encountered with MuSi-CoS, with improved resolution ( $R = \lambda/\Delta\lambda \simeq 65\,000$ ), sensitivity (approximately 15–20 per cent throughput) and wavelength coverage from 369 to 1048 nm (with gaps at 922.4–923.4, 960.8–963.6 and 1002.6–1007.4 nm). The ESPaDOnS spectropolarimeter is installed at the 3.6 m CFHT, and the NARVAL spectropolarimeter is installed at the 2 m TBL at Pic du Midi Observatory. ESPaDOnS and NARVAL are essentially identical instruments, with NARVAL constructed based on the experience of ESPaDOnS.

The polarimetric unit (at the Cassegrain focus) allows two orthogonal states of a given polarization (circular or linear) to be recorded throughout the entire spectral range. The polarimeters of

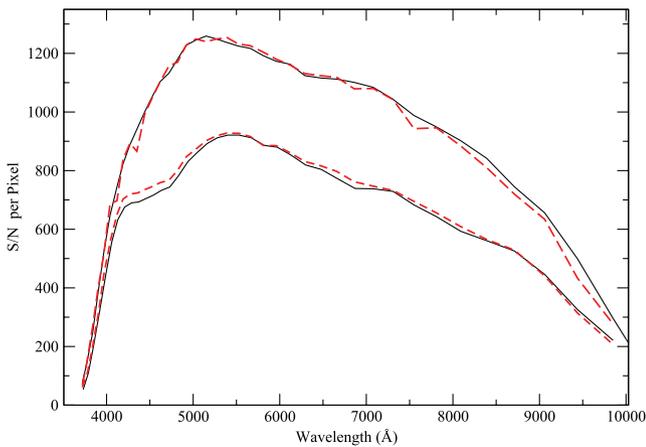


**Figure 1.** Phase coverage obtained for the target stars. The symbols denote rotational phases at which observations were obtained.

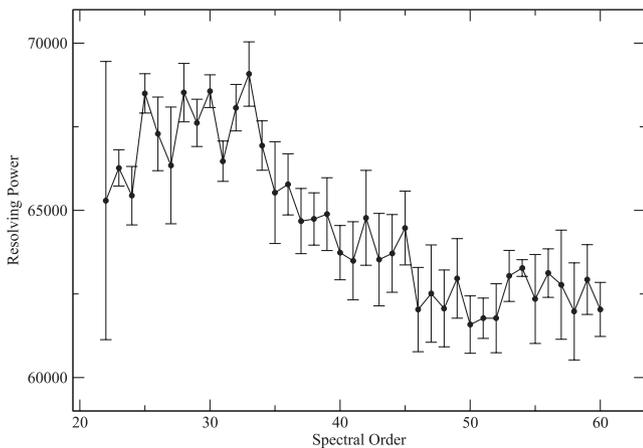
ESPaDOnS and NARVAL are of a similar design to the ‘Sempol’ visitor polarimeter (Donati et al. 2003) used on the Anglo-Australian Telescope. The polarimeter is split into two parts: the upper part is for guiding and calibration, containing the guiding camera (a commercial FLI MaxCam series CCD camera is used in NARVAL, and as of semester 2011A a QSI CCD camera is used in ESPaDOnS), an atmospheric dispersion corrector (ADC) and a calibration wheel. The lower part contains the Fresnel rhomb retarders which are used to perform the polarimetric analysis. These optics are in four different drawers which contain the following (in order): a half-wave rhomb (consisting of a pair of quarter-wave rhombs), a quarter-wave rhomb, a second half-wave rhomb and finally a Fabry–Perot wheel on one side and a Wollaston/wedge plate slide on the other side. To enhance the achromaticity of the phase delay, and to keep it at  $90^{\circ} \pm 0.5$  across the whole optical domain, the rhombs are coated with a thin layer of  $\text{MgF}_2$ , yielding a performance significantly better than achromatic crystalline plates which can vary by about  $20^{\circ}$  from the required quarter-wave retardance. Another advantage of Fresnel rhombs is that they do not produce detectable spectral ripples due to the fact that fringe spacing is of the order of the pixel size. The half-wave rhombs can rotate about the optical axis by a specified angle. The final component, the Wollaston prism, consisting of two orthogonal calcite prisms that are cemented

together, acts as a polarizing beamsplitter. The two beams of light from the beamsplitter are transmitted by about 30 m of optical fibre to the spectrograph. ESPaDOnS includes a fibre agitator, which shakes the optical fibre to remove modal noise that may be present.

The ESPaDOnS spectrograph unit consists of a double set of high-reflectance collimators cut from a single 680 mm parabolic mirror, with a focal length of 1500 mm. The grating is a 79 gr mm<sup>-1</sup> monolithic grating with dimensions of 200 mm × 400 mm. The camera lens is a fully dioptric  $f/2$ , 388 mm focal length lens, with a 210 mm free diameter (seven lenses in four blocks, one block being a 220 mm quadruplet). For cross-dispersing, a high-dispersion prism made up of a train of two identical PBL25Y prisms with 35° apex and 220 mm cross-section is used. Up until the 2011A semester the detector used with ESPaDOnS was a grade 1 EEV detector with 2 K × 4.5 K 0.0135 mm square pixels (known as EEV1 at CFHT). This was replaced in 2011A with a new E2V detector (named Olapa at CFHT).



**Figure 2.**  $S/N$  (per 1.8 km s<sup>-1</sup> spectral pixel) as a function of the wavelength for two observations of  $\alpha^2$  CVn obtained with ESPaDOnS [top curves; 2007 March 02 (air mass = 1.104) and 2009 January 13 (air mass = 1.058), both total exposures of 120 s] and two obtained with NARVAL (lower curves; 2006 December 20 (air mass = 1.250) and 2009 January 11 (air mass = 1.005), both total exposures of 240 s). The solid versus dashed lines indicate observations taken on respective nights.

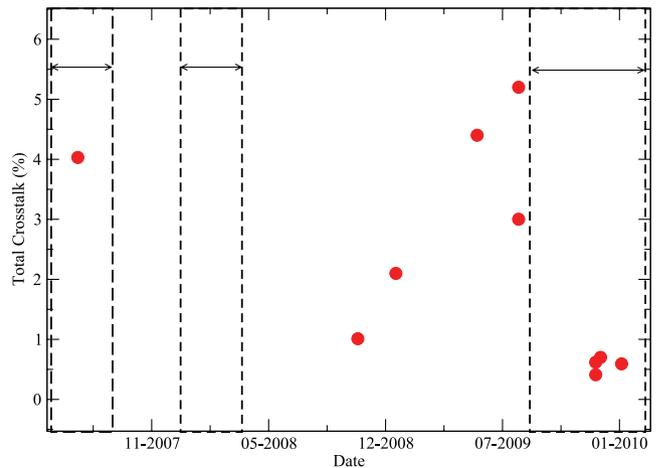


**Figure 3.** The mean spectral resolving power as a function of the spectral order for both ESPaDOnS and NARVAL for a selection of observing nights (three observations from each instrument spanning 12 months). The error bars represent the standard deviation of the resolving power of the various observations within each spectral order.

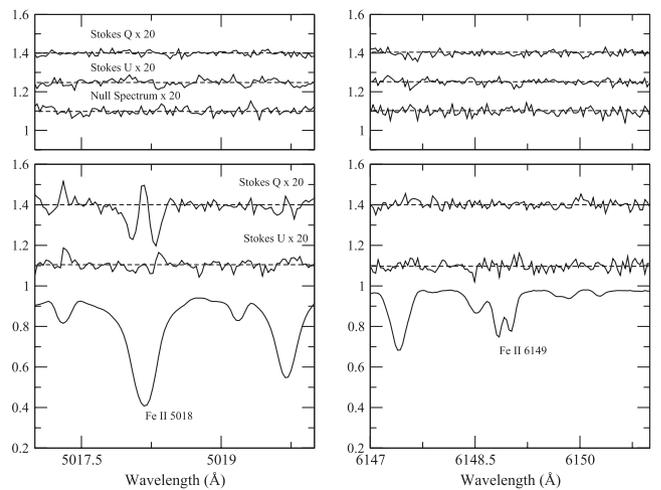
The spectrograph unit is mounted on an optical bench, which is housed in a thermal enclosure found in the inner Coudé room at CFHT.

This configuration yields full spectral coverage of the optical domain (from grating order 61 centred at 372 nm to grating order 22 centred at 1029 nm) in a single exposure. In polarimetric mode this should, in principle, achieve a resolution in excess of 65 000, but due to a charge transfer efficiency issue with the EEV1 CCD detector, the true resolution varies from approximately 68 000 in the blue to 61 000 in the red. The peak throughput of the spectrograph (with CCD detector) is about 40–45 per cent, bringing the total instrument peak efficiency to a level of about 15–20 per cent.

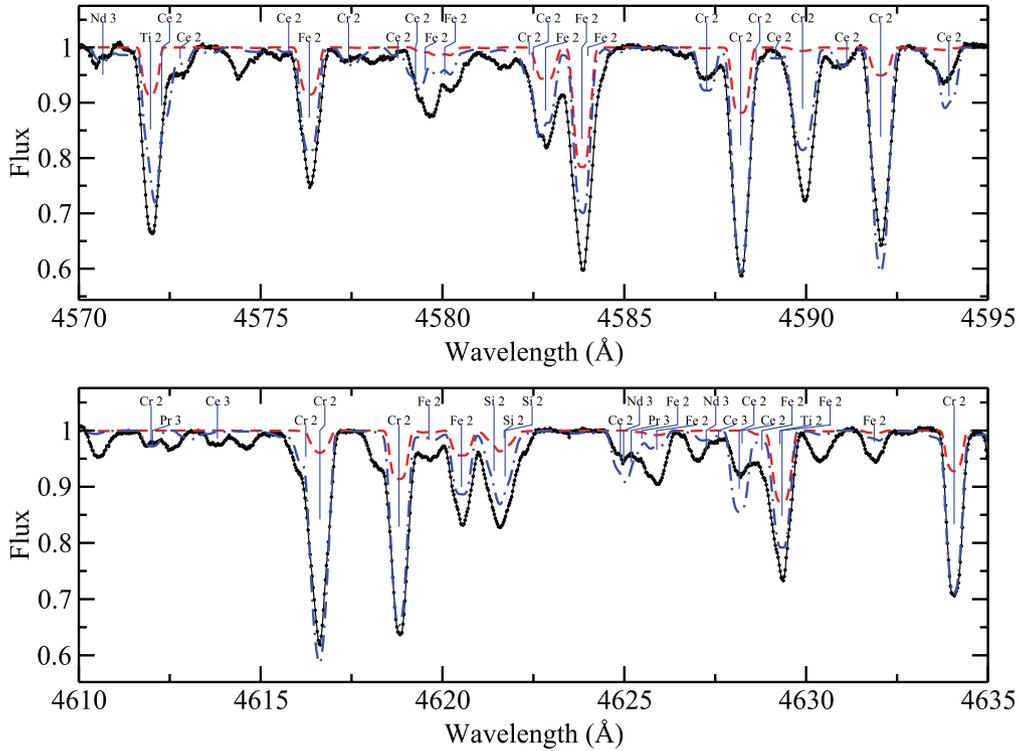
The configuration of NARVAL is much the same with the exception of a 2.8 arcsec aperture pupil (versus 1.6 arcsec for ESPaDOnS). NARVAL does however benefit from better spectrograph thermal stability (by approximately a factor of 10) than ESPaDOnS due to the use of a double thermal layer enclosure.



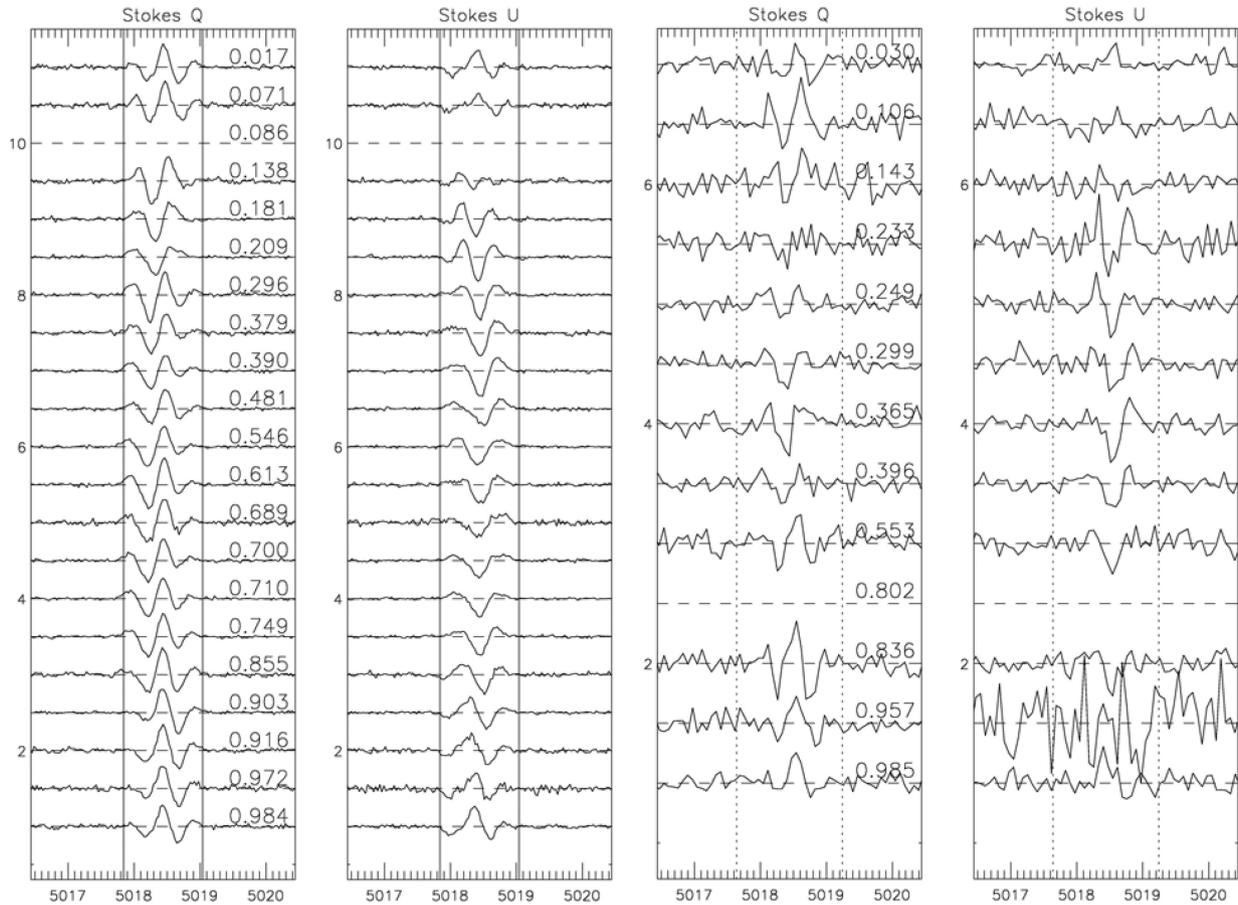
**Figure 4.** The total cross-talk of Stokes  $V$  into Stokes  $Q$  and  $U$  of ESPaDOnS as a function of time, as reported by Barrick et al. (in preparation). The dashed boxes indicate periods during which observations were obtained as part of this investigation.



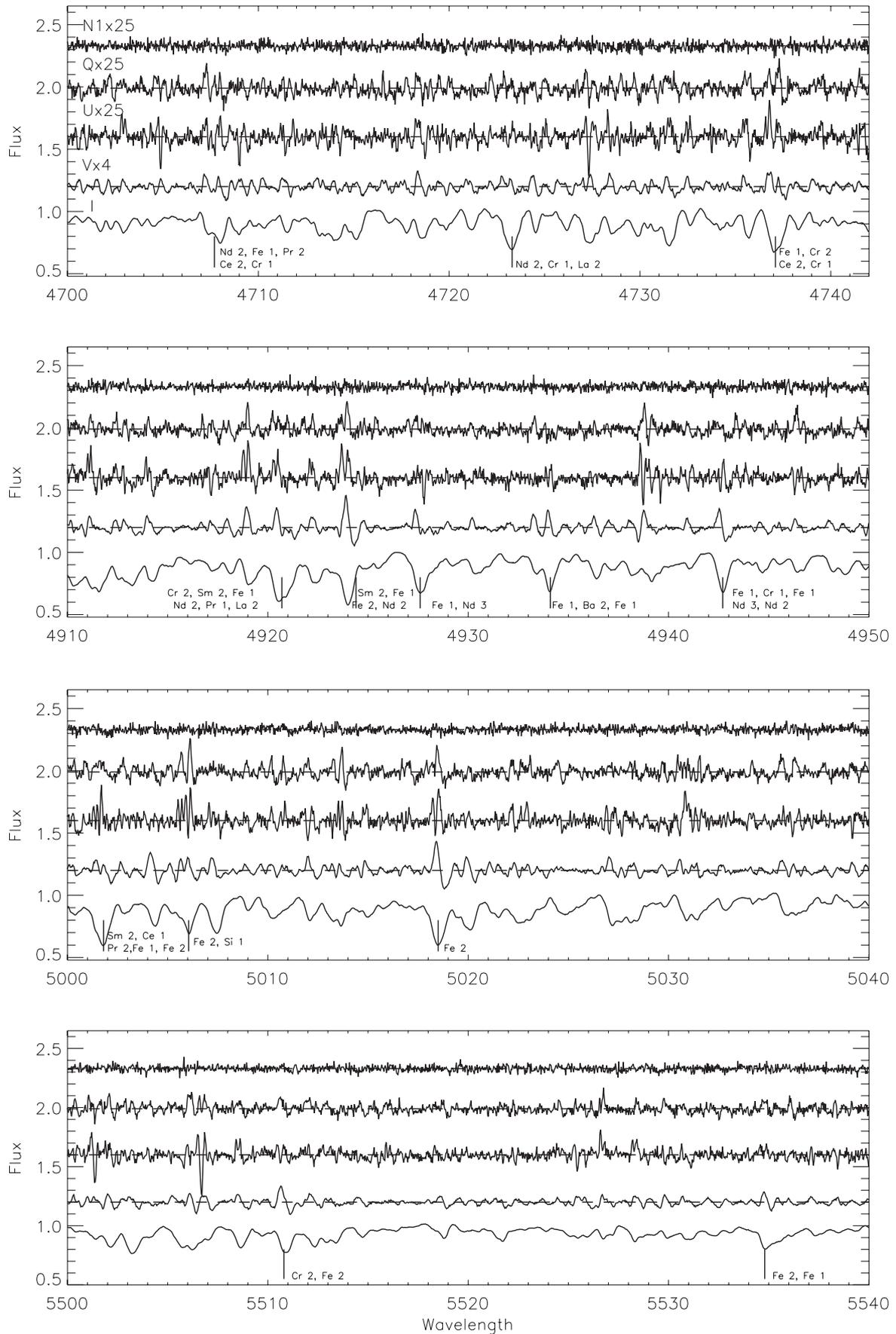
**Figure 5.** A comparison between the measured cross-talk in Stokes  $Q$  and  $U$  and the null spectrum for observations of  $\gamma$  Equ taken with ESPaDOnS (2009 July 16) with the actual observations shown below (lower frame). In the strong, magnetically sensitive Fe II 5018 line, the cross-talk in Stokes  $Q$  is below the noise, whereas the cross-talk in Stokes  $U$  is slightly above the noise.



**Figure 6.** An example of abundance fit for HD 32633, with a synthetic spectrum computed using solar abundance shown with a dashed line and the spectrum computed with the inferred abundances used in the LSD mask with a dot-dashed line.



**Figure 7.** Comparison between Stokes  $Q$  and  $U$  profiles of HD 32633 in the 5018 Å line for ESPaDOnS/NARVAL on the left-hand side and MuSiCoS on the right-hand side. Rotation phase for each observation is indicated.



**Figure 8.** Selected regions of the Stokes *IVQU* and the *N* spectra of HD 62140, with strongly contributing atomic lines labelled for reference.

The resolving power and S/N of both instruments vary with wavelength in a predictable manner, as illustrated in Fig. 2, which shows the S/N as a function of wavelength for four observations of  $\alpha^2$  CVn (two acquired with ESPaDOnS and two with NARVAL). The variation in the spectral resolving power  $R$  as a function of the spectral order for a selection of observing nights is shown in Fig. 3. This figure illustrates that the characteristics of the instruments do not vary significantly on the time-scales relevant for this project (nights to years). Variability in atmospheric conditions (e.g. seeing) may well be a dominant contributor to the scatter in resolving power.

In terms of observations the following steps occur; the two output beams from the Wollaston prism, which have been analysed into the two components of circular polarization (using the quarter-wave rhomb) or linear polarization (using the half-wave rhomb), are then carried by the pair of optical fibres to the spectrograph where two interleaved spectra are formed. The  $I$  component of the stellar Stokes vector is formed by adding the two corresponding spectra, while the  $V$ ,  $Q$  or  $U$  polarization component is obtained from the ratio method as discussed by Bagnulo et al. (2009). To minimize systematic errors due to small misalignments, differences in transmission, effects of seeing, etc., one complete observation of a star consists of four successive sub-exposures; for the second and third, the waveplate settings are changed so as to exchange the positions of the two analysed spectra on the CCD. This same procedure is used for all polarization spectra, Stokes  $V$ ,  $Q$  and  $U$ .

Calibration of the instrument uses a combination of thorium/argon and thorium/neon lamps, with the lamp calibrations taken at the beginning and at the end of each night for a primary wavelength calibration. Telluric lines are then later used to perform a second wavelength calibration during the reduction process using LIBRE-ESPRIT. Filters are used to minimize the blooming on the chip at the red end of the spectrum. Two tungsten lamps are utilized for the flat-field frames, with one low-intensity lamp being used with a red filter and the other higher intensity lamp used with a blue filter.

The reduction of observations is carried out at the observatories using the dedicated software package LIBRE-ESPRIT, which yields both the  $I$  spectrum and the  $V$  circular polarization spectrum, and/or  $QU$  linear polarization spectra of each star observed. It is important to note that LIBRE-ESPRIT automatically finds and removes continuum polarization. In this work each reduced spectrum is normalized order by order using a FORTRAN code specifically optimized to fit the continuum of these stars.

A diagnostic null spectrum called the  $N$  spectrum, computed by combining the four sub-exposures in such a way as to have real polarization cancel out, is also calculated by LIBRE-ESPRIT (again, see Bagnulo et al. 2009 for the definition of the null spectrum). The  $N$  spectrum tests the system for spurious polarization signals. The results of the reduction and normalization procedure are continuum-normalized, one-dimensional spectra in the form of wavelength,  $I/I_c$ ,  $V/I_c$ ,  $Q/I_c$ ,  $U/I_c$ , two independent normalized  $N$  spectra  $N_1/I_c$  and  $N_2/I_c$ , and an error bar (computed by propagating photon uncertainties through the reduction procedure), tabulated pixel by pixel.

Weak signatures are visible in LSD  $N$  profiles associated with some Stokes  $Q$  and  $U$  spectra. Experiments conducted in order to diagnose and mitigate polarization cross-talk (see Barrick et al., in preparation) indicate that these signatures are related to this phenomenon. However, data acquired during early NARVAL runs in 2006 exhibit substantially stronger signatures. These signatures are reported by TBL staff to result from problems with the coatings on the  $\lambda/2$  rhombs used in 2006. The rhombs were subsequently

replaced, and no similar strong signatures are detected in later observations.

Based on our comparisons of Stokes  $Q$  and  $U$  profiles obtained with the two instruments during the course of our observing programme, it is clear that these  $N$  signatures are not diagnostic of any detectable contamination of the associated  $Q$  and  $U$  profiles.

In this study we have obtained 100 polarimetric sequences corresponding to 297 individual polarized spectra. The observations were initially obtained in classical observing mode and later in service mode at both telescopes.

#### 4 CROSS-TALK

During the commissioning of ESPaDOnS in 2004 it was found that the instrument exhibited cross-talk between linear polarization and circular polarization (and vice versa). Due to the relatively strong circular polarization in the spectral lines of our targets, contamination of the significantly weaker linearly polarized profiles is potentially a serious problem.

The contributing optical component to this cross-talk identified initially was the collimating triplet within the polarimeter unit. This component was replaced in 2006 June, resulting in a reduction of the cross-talk to the 5 per cent level (from an initial level of 10–15 per cent). After replacing the triplet lens again in 2008 October, the cross-talk appeared to have been reduced to 2–3 per cent, but still exhibited strong temporal changes. It was discovered in 2009 October that the ADC was an important and previously unrecognized source of cross-talk within the polarimetric unit. The ADC was replaced in the fall of 2009, and since that time the cross-talk has been small and stable, with cross-talk from Stokes  $V$  into Stokes  $U$  at the 0.5 per cent level, and no measurable cross-talk from Stokes  $V$  into Stokes  $Q$  (i.e. below  $\sim 0.1 - 0.2$  per cent). The procedure and results of the investigations of the cross-talk are reported by Barrick et al. (in preparation). The evolution of the ESPaDOnS cross-talk with time, based on the results of Barrick et al. (in preparation), is illustrated in Fig. 4.

Less extensive cross-talk monitoring has been performed with NARVAL. Results from 2009 September indicated that with the ADC in place, the cross-talk was 3.1 per cent from Stokes  $V$  into Stokes  $Q$  and below 0.1–0.2 per cent from Stokes  $V$  into Stokes  $U$ . Without the ADC in place, the cross-talk was reduced to 2.1 per cent in Stokes  $Q$ , but increased to 1 per cent in Stokes  $U$  (illustrating that the NARVAL ADC also introduces cross-talk, but again that it is not the sole source). It is important to note that at the current time, no tests of how this cross-talk changes with time have been made with NARVAL.

**Table 2.** Reduced  $\chi^2$  values of harmonic fits to the longitudinal magnetic field phase curves determined from NARVAL and ESPaDOnS data.

Star	Stokes $V$ Fit		Null $N$ $\chi^2$
	Degree	$\chi^2$	
HD 4778	2	1.41	2.68
HD 32633	3	1.95	5.42
HD 40312	2	1.96	0.51
HD 62140	3	1.29	4.78
HD 71866	4	1.02	3.70
HD 112413	3	0.83	1.06
HD 118022	1	1.92	1.12

**Table 3.** Log of spectropolarimetric observations, where in the second to last column, an observation is denoted by a star and a missing observation by a dash. In the final column E = ESPaDOoS and N = NARVAL.

Object	Date (UT)	JD (245 0000 +)	Phase	$t_{\text{exp}}$ (s)	S/N (spectral pixel <sup>-1</sup> )	Stokes $VQU$	Instrument
HD 4778							
HD 4778	2006 November 30	4070.746	0.423	3200/3200/3200	152/709/638	***	E
	2006 December 04	4074.734	0.980	2600/2600/2600	281/667/858	***	E
	2006 December 06	4076.744	0.764	2000/2000/2000	163/647/708	***	E
	2006 December 19	4089.251	0.662	3200/3200/3200	129/641/552	***	N
	2006 December 20	4090.244	0.050	3200/3200/3200	220/507/598	***	N
	2006 December 21	4091.260	0.446	3200/3200/3200	158/585/639	***	N
	2008 January 09	4475.394	0.366	2800/-/2800	137/-/585	* - *	N
$\theta$ Aur (HD 40132)							
HD 40312	2006 December 04	4074.925	0.587	140/140/140	1231/1303/1350	***	E
	2006 December 05	4076.147	0.924	52/52/52	778/781/787	***	E
	2008 January 07	4473.417	0.710	240/240/240	783/835/836	***	N
	2008 January 09	4475.268	0.222	240/240/240	757/865/1056	***	N
	2008 January 22	4489.002	0.017	240/240/240	1377/1225/1601	***	E
	2008 January 23	4490.012	0.296	160/160/160	1296/906/1234	***	E
	2008 January 25	4491.795	0.789	160/160/160	1285/1501/1468	***	E
HD 32633							
HD 32633	2006 November 29	4070.098	0.972	3200/3200/3200	588/408/247	***	E
	2006 November 30	4070.830	0.086	3200/-/-	646/-/-	* - -	E
	2006 December 04	4074.848	0.710	3200/3200/3200	789/806/776	***	E
	2006 December 05	4075.094	0.749	3200/3200/3200	823/754/637	***	E
	2006 December 06	4076.085	0.903	3200/3200/3200	820/755/696	***	E
	2006 December 12	4082.577	0.916	1600/1600/1600	417/430/427	***	N
	2006 December 13	4083.558	0.071	3200/3200/3200	379/398/426	***	N
	2006 December 15	4085.537	0.379	3200/3200/3200	554/418/520	***	N
	2006 December 17	4087.528	0.689	3200/3200/3200	359/414/360	***	N
	2006 December 19	4089.430	0.984	3200/3200/3200	556/553/564	***	N
	2006 December 20	4090.419	0.138	3200/3200/3200	508/455/471	***	N
	2006 December 21	4091.434	0.296	3200/3200/3200	578/579/593	***	N
	2008 January 09	4475.482	0.017	3200/3200/3200	550/538/527	***	N
	2008 January 24	4490.736	0.390	2400/2400/2400	763/788/785	***	E
	2008 January 25	4491.741	0.546	2600/2600/2600	755/766/745	***	E
	2008 January 26	4492.729	0.700	2000/2000/2000	672/752/726	***	E
	2008 January 27	4493.730	0.855	2000/2000/2000	520/463/524	***	E
	2009 January 17	4849.432	0.181	3200/3200/3200	578/611/618	***	N
	2009 October 10	5115.055	0.481	2000/2000/2000	659/661/695	***	E
	2010 February 27	5254.743	0.209	2000/2000/2000	720/738/737	***	E
	2010 March 08	5263.774	0.613	2000/2000/2000	612/600/580	***	E
49 Cam (HD 62140)							
HD 62140	2006 December 10	4080.681	0.381	2800/2800/2800	685/709/735	***	N
	2006 December 12	4082.653	0.841	3200/3200/3200	673/757/764	***	N
	2006 December 13	4083.684	0.081	3200/3200/3200	518/618/599	***	N
	2006 December 14	4084.595	0.294	3200/3200/3200	666/505/586	***	N
	2006 December 17	4087.715	0.017	1600/1600/1600	148/200/182	***	N
	2006 December 19	4089.557	0.451	3200/3200/3200	668/628/605	***	N
	2006 December 20	4090.542	0.681	3200/3200/3200	656/680/664	***	N
	2006 December 21	4091.557	0.918	3200/3200/3200	774/764/743	***	N
	2008 January 09	4475.706	0.520	2400/2000/2000	612/571/588	***	N
	2008 January 24	4491.003	0.089	2600/2600/2600	957/1041/1029	***	E
	2008 January 27	4493.989	0.786	2600/2600/2600	709/836/692	***	E
	2008 March 24	4550.724	0.026	2000/2000/2000	951/929/956	***	E
	2008 March 25	4551.723	0.260	2000/2000/2000	939/944/936	***	E
	2008 April 24	4581.349	0.174	3200/3200/3200	605/740/655	***	N
	2009 January 14	4846.625	0.056	3200/3200/3200	465/363/308	***	N
	2009 January 15	4847.620	0.288	3200/3200/3200	709/723/735	***	N
	2009 September 10	5085.125	0.677	2000/2000/2000	758/853/817	***	E
	2010 February 23	5250.873	0.348	2000/2000/2000	860/782/809	***	E
	2010 February 24	5251.972	0.604	2000/2000/2000	837/800/841	***	E

Table 3 – continued

Object	Date (UT)	JD (245 0000 +)	Phase	$t_{\text{exp}}$ (s)	S/N (spectral pixel <sup>-1</sup> )	Stokes $VQU$	Instrument
HD 71866							
HD 71866	2006 December 15	4085.756	0.723	2400/2400/2400	527/544/534	***	N
	2006 December 19	4089.756	0.312	2000/2200/2200	507/477/498	***	N
	2006 December 20	4090.751	0.457	2600/2600/2600	503/466/513	***	N
	2006 December 21	4091.748	0.603	2800/2800/2800	604/649/645	***	N
	2007 March 07	4167.957	0.819	2000/2000/2000	738/748/748	***	E
	2007 March 08	4168.975	0.969	2000/2000/2000	778/779/549	***	E
	2007 May 06	4227.424	0.554	3200/3200/3200	549/537/518	***	N
	2008 January 07	4473.530	0.746	3000/3000/3000	495/423/418	***	N
	2008 January 09	4475.566	0.056	2800/2800/2800	522/612/633	***	N
	2008 January 26	4492.827	0.592	2000/2000/2000	699/734/651	***	E
	2008 January 27	4493.829	0.740	2000/2000/2000	448/453/530	***	E
	2008 April 25	4582.411	0.757	3000/3000/3000	600/444/524	***	N
	2009 January 10	4842.734	0.040	2400/2400/2400	556/579/570	***	N
2010 January 28	5224.799	0.228	2000/2000/2000	685/713/708	***	E	
HD 112413							
$\alpha^2$ CVn	2006 December 04	4075.158	0.977	240/240/240	1536/1566/1546	***	E
	2006 December 05	4076.159	0.160	120/120/120	1187/1190/1184	***	E
	2006 December 12	4082.760	0.368	240/240/240	967/910/867	***	N
	2006 December 19	4089.663	0.630	240/240/240	811/867/828	***	N
	2006 December 20	4090.647	0.810	240/240/240	772/811/812	***	N
	2007 March 02	4163.087	0.053	120/120/120	1020/1038/1049	***	E
	2007 April 23	4214.514	0.457	240/240/240	515/572/580	***	N
	2007 April 23	4214.635	0.479	240/240/240	528/512/503	***	N
	2007 April 25	4216.523	0.824	240/240/240	796/781/841	***	N
	2007 May 06	4227.516	0.834	240/240/240	918/764/875	***	N
	2007 May 07	4228.467	0.007	400/400/400	589/494/795	***	N
	2008 January 22	4489.024	0.646	240/240/240	1482/1555/1557	***	E
	2008 January 23	4490.031	0.830	240/240/240	1564/1470/1569	***	E
	2008 March 23	4549.044	0.621	120/120/120	1086/1133/1088	***	E
	2008 March 24	4550.850	0.951	120/120/120	1011/947/1035	***	E
	2008 March 25	4551.891	0.141	120/120/120	1190/1181/1196	***	E
	2008 April 24	4581.460	0.548	240/240/240	931/896/906	***	N
	2008 April 25	4582.574	0.752	240/240/240	1082/1067/1077	***	N
	2008 July 26	4672.734	0.236	120/120/120	1010/877/1015	***	E
	2008 August 24	4703.724	0.901	120/120/120	370/226/357	***	E
	2008 August 24	4703.736	0.903	120/120/120	361/338/464	***	E
	2009 January 11	4843.728	0.500	240/240/240	767/797/825	***	N
	2009 January 13	4846.143	0.940	120/120/120	1026/1085/1039	***	E
	2010 January 26	5223.163	0.873	240/240/240	1448/1458/1393	***	E
	2010 June 04	5351.852	0.402	240/240/240	1665/1637/1589	***	E
	2010 June 20	5367.761	0.311	240/240/240	1570/1557/1613	***	E
	2010 June 22	5369.856	0.694	240/240/240	1072/983/972	***	E
78 Vir (HD 118022)							
78 Vir	2007 April 23	4214.570	0.623	800/800/800	425/369/395	***	N
	2007 April 25	4216.580	0.162	1000/1000/1000	327/378/346	***	N
	2007 May 06	4227.503	0.095	1400/1400/1400	681/821/741	***	N
	2008 April 24	4581.520	0.211	1200/1200/1200	587/634/588	***	N
	2008 April 25	4582.562	0.490	1200/1200/1200	891/928/944	***	N

Although the cross-talk is now below 1 per cent in ESPaDOnS and probably around 2 per cent in NARVAL, observations for this project have been acquired over the last four years, and during some of this time the cross-talk may have been higher (and almost certainly was for ESPaDOnS). It is important to understand how cross-talk could affect the Stokes  $Q$  and  $U$  signatures measured in

spectral lines. A series of tests were therefore performed to evaluate the importance of this cross-talk.

The first test was to examine the Stokes profiles of the Fe II  $\lambda 6149$  line. This line has a relatively large Landé factor, but in the linear regime of the Zeeman effect it is predicted to be purely circularly polarized as a consequence of the  $\sigma$  and  $\pi$  components of this line

having the same strength and identical splitting. In this case we would interpret any signal in Stokes  $Q$  and  $U$  as due to cross-talk from Stokes  $V$ . We have examined  $\lambda 6149$  in our spectra, as well as spectra discussed by Barrick et al. (in preparation). In no case do we observe any significant signal in this line.

Then, we carefully examined one of the cross-talk diagnostic observations of the cool magnetic star  $\gamma$  Equ obtained during CFHT engineering time. As described by Barrick et al. (in preparation), on-sky cross-talk diagnosis employs observations of slowly rotating magnetic Ap stars observed in all Stokes parameters at two positions of the CFHT's Cassegrain bonnet to unambiguously measure the cross-talk into Stokes  $Q$  and  $U$ . The diagnostic observation of  $\gamma$  Equ (obtained in 2009 July, with a peak S/N of over 1000) yielded relatively high cross-talk levels of 2.3 per cent in Stokes  $Q$  and 5.1 per cent in Stokes  $U$ . By comparing the observations with and without cross-talk, it was found that the cross-talk contribution to Stokes  $Q$  is within the noise, whilst the cross-talk contribution to Stokes  $U$  appears to be slightly above the noise, as illustrated by Fig. 5. It is important to note that the removal of the cross-talk using techniques as described in Barrick et al. (in preparation) requires the acquisition of a second series of observations following a rotation of the Cassegrain bonnet, a procedure which is not a standard practice during regular observations.

$\gamma$  Equ is an extremely sharp lined star, which means that the cross-talk contribution will have a greater effect than that in the broader lined stars studied in this paper. Nevertheless, the potential influence of the cross-talk on MDI will be investigated and discussed in a future paper (Silvester et al., in preparation).

These comparisons demonstrate that the contribution of cross-talk to the Stokes profiles is below, or at most just above, the level of the noise of the best-quality observations of Ap stars acquired with these instruments.

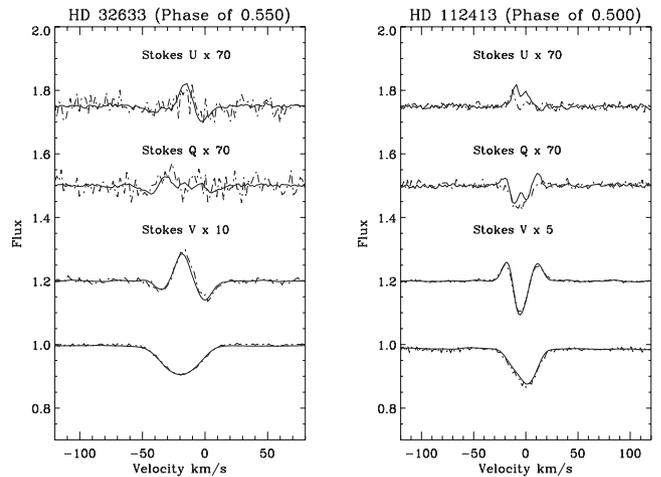
## 5 LONGITUDINAL MAGNETIC FIELD AND NET LINEAR POLARIZATION

To examine the self-consistency of the new polarization spectra and to evaluate their consistency with published magnetic data for our targets, we have measured the mean longitudinal magnetic field and net linear polarization using LSD. LSD (Donati et al. 1997; Kochukhov, Makaganiuk & Piskunov 2010) is a multiline analysis method that produces mean Stokes  $I$  and  $V$  profiles using essentially all metallic lines in the stellar spectrum. It assumes that the observed spectrum can be represented as the convolution of a single mean line profile with an underlying spectrum of unbroadened metal and helium lines of appropriate wavelength, depth and Landé factor (the 'line mask' computed using spectrum synthesis; e.g. Wade et al. 2000a).

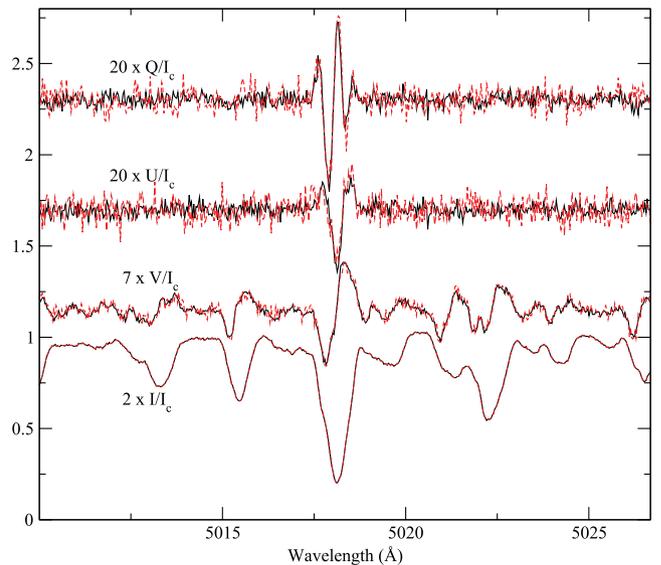
The LSD model allows the computation of single, average Stokes  $I$  and  $V$  line profiles, usually characterized by an S/N significantly higher than that of individual spectral lines, scaling roughly as the square root of the number of lines used.

The projected rotational velocity  $v \sin i$  of each star was derived by fitting a selection of spectral lines (typically in the 4500 Å region) with a synthesized spectrum created using the SYNTH3 spectrum synthesis code (Kochukhov 2007). Each line in the selection had  $v \sin i$  determined by using a  $\chi^2$  fitting function which is included as part of BINMAG, the spectral visualization tool. The resulting mean (and standard deviation) values are reported in Table 1. Line masks for this study were compiled using Vienna Atomic Line Database (Kupka et al. 1999) 'extract stellar' requests, with effective temperatures (adopted based on the literature) of each target. It should be noted

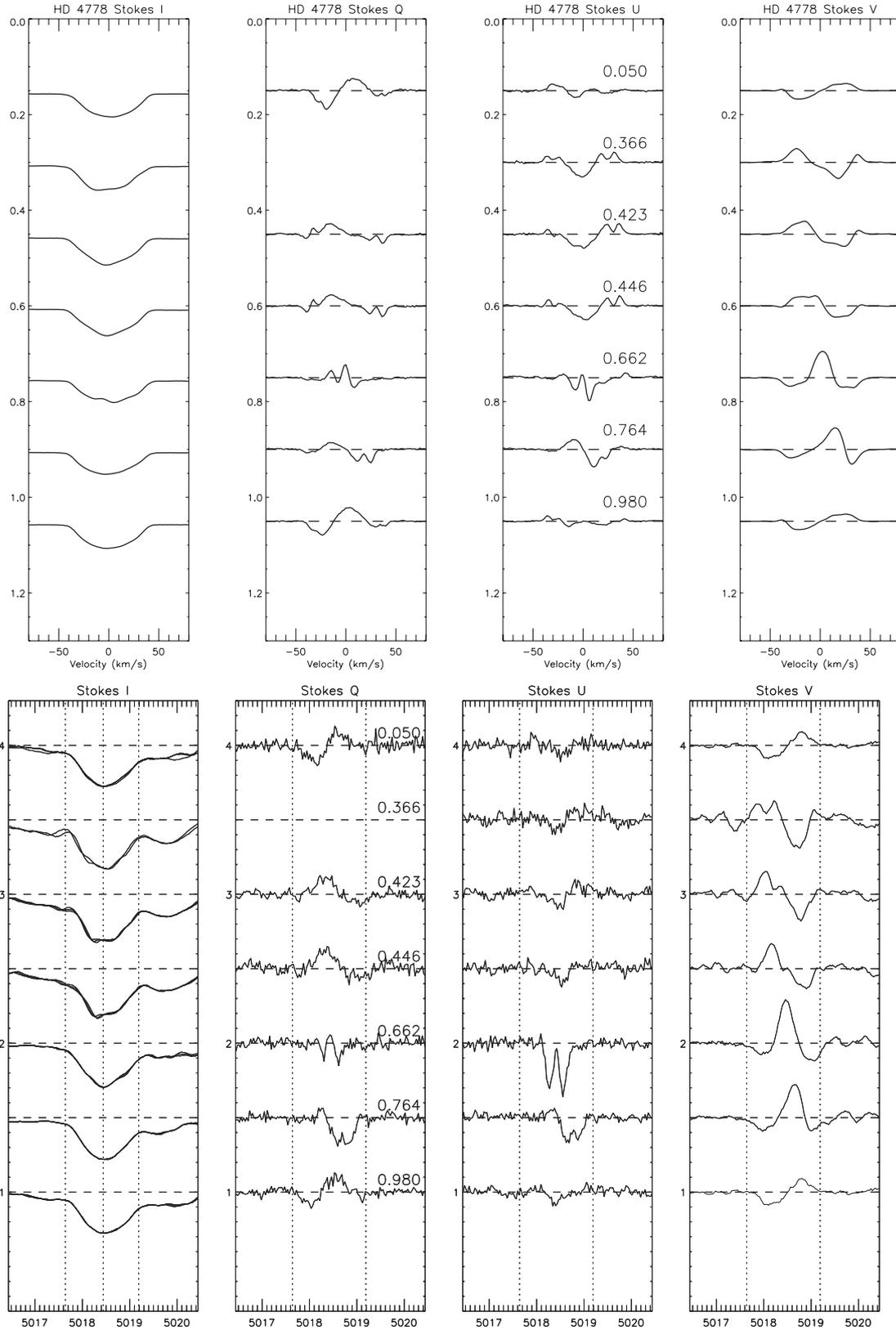
that when creating each line mask, a constant  $\log g$  of 4.0 was used for all stars. Input chemical abundances were determined by a rough abundance analysis, based typically on lines around the 4500 Å region. The abundance analysis was performed by using the SYNTH3 spectrum synthesis code (Kochukhov 2007) with ATLAS9 solar abundance model atmospheres used to create synthetic spectra that were compared directly to the observations. The input abundances were then adjusted in the synthetic spectrum until reasonable agreement was found with the observed spectrum. A comparison between two synthetic spectra, one computed using solar abundances and one using the final determined abundances, is shown in Fig. 6 for HD 32633. The final abundances were then adopted in the line mask creation. By doing so we ensure that any uncertainties caused by the mask when performing the LSD analysis are minimal.



**Figure 9.** Comparison between LSD profiles from ESPaDOnS/NARVAL observations and MuSiCoS observations taken at near identical phases for HD 32633 and HD 112413 (MuSiCoS data: dashed lines; ESPaDOnS/NARVAL: solid lines). Good agreement can be seen.



**Figure 10.** Two observations of HD 32633 obtained at close phases, one with NARVAL (phase 0.689, taken on 2006 December 17) and one with ESPaDOnS (phase 0.700, taken on 2008 January 26). The S/N is  $\approx 400$  and  $\approx 700$ , respectively. Excellent agreement between the signatures can be seen, illustrating the consistency between the two instruments.



**Figure 11.** Top: variation of Stokes  $I$ ,  $Q$ ,  $U$  and  $V$  LSD profiles for HD 4778. Rotational phase increases from top to bottom, represented on the vertical axis. A scaling factor of 25 was used for both the Stokes  $Q$  and  $U$  observations, and a factor of 5 was used for Stokes  $V$ . Bottom: variation of Stokes  $I$ ,  $Q$ ,  $U$  and  $V$  profiles for HD 4778 in the Fe  $\Pi$   $\lambda$ 5018 line (scaled by a factor of 20 for Stokes  $Q$  and  $U$ , and a factor of 4 for Stokes  $V$ ).

**Table 4.** Longitudinal magnetic field and net linear polarization measurements of magnetic A and B stars, measured from LSD Stokes  $V$  profiles, and Stokes  $Q$  and  $U$  profiles.

JD – 245 0000	Phase	$B_\ell \pm \sigma_B$ (G)	$Q/I \pm \sigma_Q$	$U/I \pm \sigma_U$
HD 4778				
4070.746	0.423	$1584 \pm 22$	$0.002\,535 \pm 0.000\,322$	$-0.002\,785 \pm 0.000\,348$
4074.734	0.980	$-1299 \pm 15$	$-0.000\,206 \pm 0.000\,341$	$-0.001\,108 \pm 0.000\,312$
4076.744	0.764	$-613 \pm 18$	$-0.004\,080 \pm 0.000\,390$	$-0.003\,208 \pm 0.000\,370$
4089.251	0.662	$191 \pm 21$	$0.001\,527 \pm 0.000\,436$	$0.007\,873 \pm 0.000\,454$
4090.244	0.050	$-1299 \pm 15$	$0.003\,215 \pm 0.000\,516$	$0.000\,027 \pm 0.000\,477$
4091.260	0.446	$1560 \pm 22$	$-0.003\,398 \pm 0.000\,387$	$0.003\,088 \pm 0.000\,355$
4475.394	0.366	$1415 \pm 24$	–	$-0.000\,771 \pm 0.000\,424$
$\theta$ Aur (HD 40132)				
4074.925	0.587	$284 \pm 11$	$-0.000\,802 \pm 0.000\,414$	$0.000\,139 \pm 0.000\,411$
4076.147	0.924	$-119 \pm 17$	$-0.000\,159 \pm 0.000\,810$	$0.001\,076 \pm 0.000\,817$
4473.417	0.710	$232 \pm 17$	$0.000\,107 \pm 0.000\,736$	$0.001\,391 \pm 0.000\,738$
4475.268	0.222	$43 \pm 14$	$-0.001\,876 \pm 0.000\,622$	$0.000\,512 \pm 0.000\,534$
4489.002	0.017	$-184 \pm 11$	$-0.001\,646 \pm 0.000\,456$	$0.000\,054 \pm 0.000\,389$
4490.012	0.296	$108 \pm 9$	$0.001\,905 \pm 0.000\,388$	$0.000\,731 \pm 0.000\,417$
4491.795	0.789	$156 \pm 10$	$-0.001\,068 \pm 0.000\,446$	$-0.000\,547 \pm 0.000\,460$
HD 32633				
4070.098	0.972	$-3641 \pm 34$	$0.000\,794 \pm 0.000\,830$	$-0.000\,148 \pm 0.001\,156$
4070.830	0.086	$-953 \pm 31$	–	–
4074.848	0.710	$-2281 \pm 16$	$0.000\,290 \pm 0.000\,383$	$0.000\,450 \pm 0.000\,384$
4075.094	0.749	$-2835 \pm 20$	$0.000\,576 \pm 0.000\,406$	$0.000\,601 \pm 0.000\,457$
4076.085	0.903	$-4387 \pm 28$	$0.000\,267 \pm 0.000\,384$	$0.000\,267 \pm 0.000\,401$
4082.577	0.916	$-4245 \pm 30$	$-0.000\,343 \pm 0.000\,688$	$-0.000\,122 \pm 0.000\,696$
4083.558	0.071	$-1174 \pm 33$	$0.000\,590 \pm 0.000\,790$	$-0.000\,113 \pm 0.000\,743$
4085.537	0.379	$423 \pm 11$	$-0.000\,766 \pm 0.000\,632$	$-0.001\,007 \pm 0.000\,632$
4087.528	0.689	$-2165 \pm 18$	$0.000\,202 \pm 0.000\,788$	$-0.000\,962 \pm 0.000\,907$
4089.430	0.984	$-3170 \pm 35$	$-0.000\,083 \pm 0.000\,536$	$-0.000\,059 \pm 0.000\,525$
4090.419	0.138	$352 \pm 32$	$0.000\,343 \pm 0.000\,699$	$0.000\,409 \pm 0.000\,673$
4091.434	0.296	$1388 \pm 18$	$-0.001\,486 \pm 0.000\,549$	$-0.000\,241 \pm 0.000\,539$
4475.482	0.017	$-2753 \pm 46$	$-0.000\,062 \pm 0.000\,569$	$-0.000\,249 \pm 0.000\,573$
4490.736	0.390	$404 \pm 12$	$0.001\,182 \pm 0.000\,388$	$0.000\,729 \pm 0.000\,395$
4491.741	0.546	$-1366 \pm 17$	$0.000\,510 \pm 0.000\,441$	$0.000\,231 \pm 0.000\,436$
4492.729	0.700	$-2185 \pm 20$	$0.000\,384 \pm 0.000\,449$	$0.000\,414 \pm 0.000\,449$
4493.730	0.855	$-4256 \pm 40$	$0.000\,263 \pm 0.000\,642$	$0.000\,013 \pm 0.000\,558$
4849.432	0.181	$898 \pm 18$	$0.000\,194 \pm 0.000\,515$	$-0.000\,166 \pm 0.000\,509$
5115.055	0.481	$-925 \pm 13$	$0.000\,590 \pm 0.000\,477$	$0.000\,413 \pm 0.000\,449$
5254.743	0.209	$1273 \pm 19$	$0.000\,285 \pm 0.000\,413$	$-0.000\,176 \pm 0.000\,422$
5263.774	0.613	$-1559 \pm 15$	$-0.000\,030 \pm 0.000\,536$	$0.000\,529 \pm 0.000\,547$
49 Cam (HD 62140)				
4080.681	0.381	$-1350 \pm 13$	$-0.001\,999 \pm 0.000\,437$	$-0.000\,067 \pm 0.000\,459$
4082.653	0.841	$1100 \pm 14$	$0.001\,359 \pm 0.000\,447$	$-0.004\,501 \pm 0.000\,452$
4083.684	0.081	$1086 \pm 14$	$0.002\,122 \pm 0.000\,461$	$-0.003\,331 \pm 0.000\,514$
4084.595	0.294	$-819 \pm 19$	$-0.000\,318 \pm 0.000\,594$	$0.001\,067 \pm 0.000\,563$
4087.715	0.017	$1253 \pm 23$	$-0.000\,198 \pm 0.001\,432$	$-0.000\,483 \pm 0.001\,584$
4089.557	0.451	$-1471 \pm 12$	$-0.002\,095 \pm 0.000\,476$	$0.000\,024 \pm 0.000\,511$
4090.542	0.681	$-307 \pm 21$	$0.005\,206 \pm 0.000\,519$	$0.003\,038 \pm 0.000\,486$
4091.557	0.918	$1298 \pm 13$	$-0.001\,324 \pm 0.000\,393$	$-0.000\,273 \pm 0.000\,438$
4475.706	0.520	$-1358 \pm 13$	$-0.002\,141 \pm 0.000\,512$	$0.001\,236 \pm 0.000\,513$
4491.003	0.089	$1112 \pm 13$	$0.003\,333 \pm 0.000\,323$	$-0.003\,265 \pm 0.000\,370$
4493.989	0.786	$598 \pm 17$	$0.003\,778 \pm 0.000\,473$	$-0.003\,123 \pm 0.000\,484$
4550.724	0.026	$1255 \pm 12$	$0.000\,867 \pm 0.000\,322$	$-0.001\,192 \pm 0.000\,352$
4551.723	0.260	$-434 \pm 21$	$0.001\,381 \pm 0.000\,389$	$0.001\,123 \pm 0.000\,438$
4581.349	0.174	$370 \pm 21$	$0.003\,960 \pm 0.000\,457$	$0.002\,064 \pm 0.000\,531$
4846.625	0.056	$1164 \pm 14$	$0.001\,132 \pm 0.000\,830$	$-0.002\,572 \pm 0.000\,928$
4847.620	0.288	$-749 \pm 20$	$0.000\,353 \pm 0.000\,463$	$0.001\,085 \pm 0.000\,497$
5085.125	0.677	$-478 \pm 20$	$0.005\,382 \pm 0.000\,443$	$0.002\,041 \pm 0.000\,419$
5250.873	0.348	$-1125 \pm 14$	$-0.001\,510 \pm 0.000\,417$	$0.000\,311 \pm 0.000\,436$
5251.972	0.604	$-1024 \pm 13$	$0.002\,344 \pm 0.000\,407$	$0.000\,870 \pm 0.000\,401$

Table 4 – continued

JD – 245 0000	Phase	$B_\ell \pm \sigma_B$ (G)	$Q/I \pm \sigma_Q$	$U/I \pm \sigma_U$
HD 71866				
4085.756	0.723	$-169 \pm 26$	$-0.001\,707 \pm 0.000\,558$	$-0.004\,811 \pm 0.000\,603$
4089.756	0.312	$-629 \pm 20$	$-0.007\,183 \pm 0.000\,630$	$-0.001\,153 \pm 0.000\,592$
4090.751	0.457	$-1785 \pm 19$	$-0.002\,275 \pm 0.000\,598$	$-0.003\,174 \pm 0.000\,559$
4091.748	0.603	$-1580 \pm 21$	$-0.003\,833 \pm 0.000\,532$	$0.001\,778 \pm 0.000\,484$
4167.957	0.819	$1316 \pm 24$	$-0.000\,151 \pm 0.000\,441$	$-0.003\,645 \pm 0.000\,506$
4168.975	0.969	$2223 \pm 21$	$-0.002\,215 \pm 0.000\,431$	$-0.000\,973 \pm 0.000\,672$
4227.424	0.554	$-1750 \pm 19$	$-0.002\,216 \pm 0.000\,554$	$0.000\,866 \pm 0.000\,535$
4473.530	0.746	$260 \pm 26$	$0.001\,818 \pm 0.000\,718$	$-0.005\,288 \pm 0.000\,736$
4475.566	0.056	$2290 \pm 22$	$0.000\,009 \pm 0.000\,456$	$-0.000\,153 \pm 0.000\,629$
4492.827	0.592	$-1688 \pm 18$	$-0.003\,414 \pm 0.000\,484$	$0.001\,299 \pm 0.000\,474$
4493.829	0.740	$43 \pm 25$	$0.002\,172 \pm 0.000\,653$	$-0.004\,723 \pm 0.000\,629$
4582.411	0.757	$461 \pm 26$	$0.001\,491 \pm 0.000\,686$	$-0.004\,704 \pm 0.000\,613$
4842.734	0.040	$2323 \pm 22$	$0.000\,023 \pm 0.000\,475$	$-0.000\,549 \pm 0.000\,659$
5224.799	0.228	$971 \pm 25$	$-0.004\,270 \pm 0.000\,440$	$0.000\,347 \pm 0.000\,471$
HD 112413				
4075.158	0.977	$-856 \pm 23$	$0.001\,780 \pm 0.000\,494$	$0.001\,723 \pm 0.000\,484$
4076.159	0.160	$-603 \pm 19$	$0.002\,334 \pm 0.000\,525$	$0.004\,213 \pm 0.000\,518$
4082.760	0.368	$624 \pm 21$	$0.000\,137 \pm 0.000\,617$	$-0.001\,208 \pm 0.000\,647$
4089.663	0.630	$231 \pm 23$	$0.000\,691 \pm 0.000\,657$	$0.001\,269 \pm 0.000\,687$
4090.647	0.810	$-785 \pm 19$	$0.003\,106 \pm 0.000\,790$	$0.002\,424 \pm 0.000\,788$
4163.087	0.053	$-795 \pm 21$	$-0.001\,239 \pm 0.000\,734$	$-0.000\,578 \pm 0.000\,718$
4214.514	0.457	$807 \pm 18$	$0.000\,226 \pm 0.001\,107$	$-0.001\,956 \pm 0.001\,125$
4214.635	0.479	$799 \pm 19$	$0.001\,272 \pm 0.000\,993$	$-0.001\,996 \pm 0.000\,981$
4216.523	0.824	$-830 \pm 20$	$0.003\,010 \pm 0.000\,838$	$0.002\,400 \pm 0.000\,780$
4227.516	0.834	$-849 \pm 19$	$0.003\,131 \pm 0.000\,848$	$0.001\,945 \pm 0.000\,738$
4228.467	0.007	$-818 \pm 24$	$-0.000\,271 \pm 0.001\,652$	$-0.001\,587 \pm 0.000\,992$
4489.024	0.646	$129 \pm 23$	$-0.000\,916 \pm 0.000\,367$	$0.001\,985 \pm 0.000\,378$
4490.031	0.830	$-851 \pm 22$	$0.002\,872 \pm 0.000\,452$	$0.002\,218 \pm 0.000\,433$
4549.044	0.621	$304 \pm 22$	$-0.000\,453 \pm 0.000\,489$	$-0.000\,584 \pm 0.000\,524$
4550.850	0.951	$-888 \pm 22$	$0.002\,450 \pm 0.000\,779$	$0.001\,239 \pm 0.000\,691$
4551.891	0.141	$-595 \pm 21$	$0.001\,459 \pm 0.000\,530$	$0.004\,084 \pm 0.000\,518$
4581.460	0.548	$657 \pm 20$	$-0.001\,276 \pm 0.000\,634$	$0.001\,097 \pm 0.000\,631$
4582.574	0.752	$-595 \pm 21$	$0.002\,091 \pm 0.000\,553$	$0.003\,702 \pm 0.000\,545$
4672.734	0.236	$-193 \pm 23$	$0.004\,147 \pm 0.000\,629$	$-0.001\,381 \pm 0.000\,543$
4703.724	0.901	$-854 \pm 31$	$0.002\,628 \pm 0.003\,134$	$-0.000\,760 \pm 0.001\,992$
4703.736	0.903	$-907 \pm 30$	$0.002\,653 \pm 0.002\,106$	$-0.000\,584 \pm 0.001\,505$
4843.728	0.500	$-785 \pm 18$	$0.001\,821 \pm 0.000\,779$	$-0.002\,499 \pm 0.000\,754$
4846.143	0.940	$-875 \pm 22$	$0.002\,773 \pm 0.000\,675$	$-0.001\,509 \pm 0.000\,701$
5223.163	0.873	$-884 \pm 19$	$0.002\,989 \pm 0.000\,488$	$0.000\,467 \pm 0.000\,503$
5351.852	0.402	$739 \pm 16$	$-0.000\,142 \pm 0.000\,361$	$0.002\,013 \pm 0.000\,375$
5367.761	0.311	$296 \pm 23$	$0.001\,792 \pm 0.000\,379$	$-0.001\,181 \pm 0.000\,359$
5369.856	0.694	$-225 \pm 21$	$0.000\,081 \pm 0.000\,584$	$0.003\,735 \pm 0.000\,598$
78 Vir (HD 118022)				
4214.570	0.623	$-905 \pm 12$	$-0.005\,410 \pm 0.000\,359$	$0.001\,493 \pm 0.000\,380$
4216.580	0.162	$-241 \pm 10$	$-0.003\,051 \pm 0.000\,700$	$-0.005\,213 \pm 0.000\,658$
4227.503	0.095	$-127 \pm 5$	$-0.001\,195 \pm 0.000\,454$	$-0.005\,993 \pm 0.000\,528$
4581.520	0.211	$-374 \pm 11$	$-0.005\,294 \pm 0.000\,757$	$-0.003\,943 \pm 0.000\,973$
4582.562	0.490	$-957 \pm 10$	$-0.004\,111 \pm 0.000\,320$	$0.002\,266 \pm 0.000\,365$

As discussed by Shorlin et al. (2002), the LSD S/N is only weakly sensitive to the line-depth cutoff employed to populate the mask. Following their results, we have chosen to employ a line-depth cutoff equal to 10 per cent of the continuum. Imposing such a cutoff has a related advantage: because weaker lines are less likely to have published experimental Landé factors (and to generally have more poorly determined atomic data), we pre-filter our line list to (statistically) exclude those lines with the poorest data. In addition, Balmer lines are removed from the mask and the mask is restricted

to the ESPaDOnS/NARVAL spectral range. Application of LSD to the data yields a set of mean profiles (Stokes  $I$ , Stokes  $V$  and  $N$ ) for each reduced spectrum.

The mean longitudinal magnetic field  $B_\ell$  was computed from each LSD profile set. This quantity was evaluated by computing the first-order moment of the Stokes  $V$  profile in velocity according to

$$B_\ell = -2.14 \times 10^{11} \frac{\int (v - v_0) V(v) dv}{\lambda g c \int [1 - I(v)] dv} \quad (1)$$

(Mathys 1989; Donati et al. 1997; Wade et al. 2000a), where  $v_0$  is the centre of gravity of the Stokes  $V$  profile,  $g$  is the integrated mean Landé factor and  $\lambda$  is the weighted mean wavelength of all the lines included in the mask. LSD profiles were locally renormalized to a continuum level of 1.0 before evaluation of equation (1). Uncertainties associated with  $B_l$  were computed by propagating the formal uncertainties of each LSD spectral pixel through equation (1). LSD profiles are extracted for each star uniformly weighted to the same Landé factor, line depth and wavelength.

To determine the net linear polarization (see, e.g., Wade et al. 2000a) the LSD Stokes  $Q$  or  $U$  profile was integrated to compute the normalized equivalent width of the line polarization using

$$\frac{Q}{I} = \frac{\int Q(v) dv}{\int [1 - I(v)] dv}. \quad (2)$$

It was found that both net linear polarization and longitudinal field measurements were sensitive to the integration range used to calculate equations (1) and (2). Integration ranges were carefully chosen to include the entire line profile while avoiding the inclusion of excess continuum outside of the profile (which contributes to only noise). This was accomplished initially by selecting limits based on the apparent extent of the wings of the Stokes  $I$  profile. We then evaluated visually if any significant polarized flux was located outside of the limits, adjusting the integration limits as necessary.

As was the case with Stokes  $V$ , LSD Stokes  $I$  profiles associated with Stokes  $Q$  and  $U$  profiles were locally renormalized to a continuum level of 1.0 before evaluation of equation (2).

Longitudinal field measurements are phased according to the ephemerides given in Table 1. We verified that all adopted periods were sufficiently precise that no significant relative phase uncertainties exist. A harmonic curve was fitted by least squares to the phased longitudinal field data. The degree of the harmonic function which yielded the lowest reduced  $\chi^2$ , while still providing a significant and systematic improvement to the fit, was chosen as the ‘best’ fit. The results of the fit are shown in Table 2. It should be noted that whilst the fits to the longitudinal field variations help compare different data sets, quantify the dispersion of the data and then verify the accuracy of the error bars, their parameters and degree do not necessarily have any easily interpretable physical meaning. The log of observations is given in Table 3.

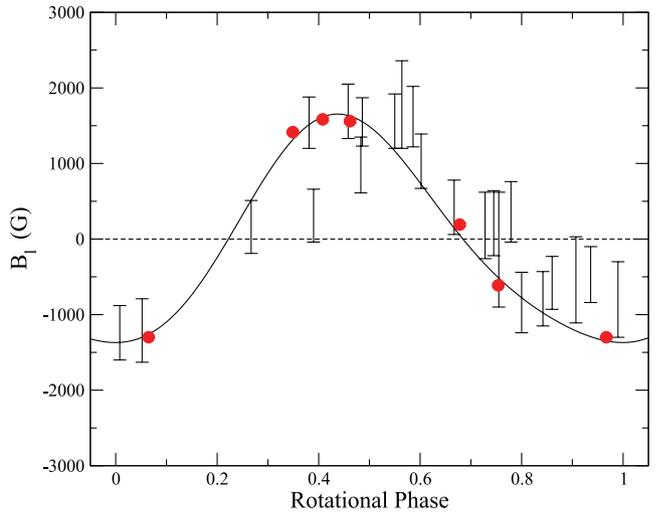
To allow a consistent comparison between the MuSiCoS and ESPaDOnS/NARVAL data, the normalization of the ESPaDOnS/NARVAL data had to be performed carefully. Care also had to be taken with the integration limits when determining both the longitudinal field and net linear polarization. A comparison between ESPaDOnS/NARVAL and MuSiCoS linear polarization for HD 32633 can be seen in Fig. 7. In addition, an example region of  $IVQU$  spectra is shown in Fig. 8 for HD 62140. Illustrating the wealth of individual lines.

When we compared the new longitudinal field measurements with those computed from MuSiCoS spectra, immediate agreement was found between the new measurements and those of Wade et al. (2000b), except in the case of HD 71866 and  $\alpha^2$  CVn which showed a slight discrepancy. By recomputing the Stokes  $V$  MuSiCoS LSD profiles for these stars with the new masks, the observations were brought into agreement. It is possible that if both HD 71866 and  $\alpha^2$  CVn have sufficiently peculiar abundances, a general Ap star line mask (with arbitrarily enhanced and depleted abundances of particular elements, as used in the original analysis of Wade et al. 2000b) may result in an underestimation of the longitudinal field. Because each mask used in this study is ‘tailor-made’ to the chemical abundances of the respective star, we are less susceptible to

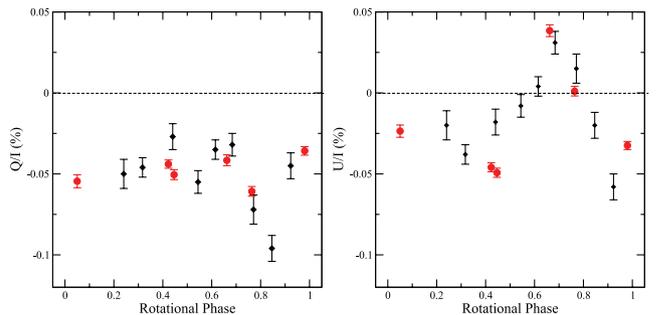
this problem and thus measure slightly increased longitudinal field values.

An example of the comparison between the new measurements and those corresponding to the two sets of MuSiCoS LSD profiles is described in Section 6.6 (Fig. 27) for  $\alpha^2$  CVn. This figure demonstrates that overall the agreement between the new measurements and the recomputed MuSiCoS measurements is acceptable. Using the new masks for the other stars makes little difference to the longitudinal field measurements, so the original measurements were kept for comparison. In addition, good agreement can be seen between LSD profiles obtained with MuSiCoS and ESPaDOnS/NARVAL observations. Fig. 9 shows that LSD profiles both from MuSiCoS and the current observations are consistent for an identical phase of observation in the case of HD 32633 and HD 112413, with the new data showing much reduced noise levels.

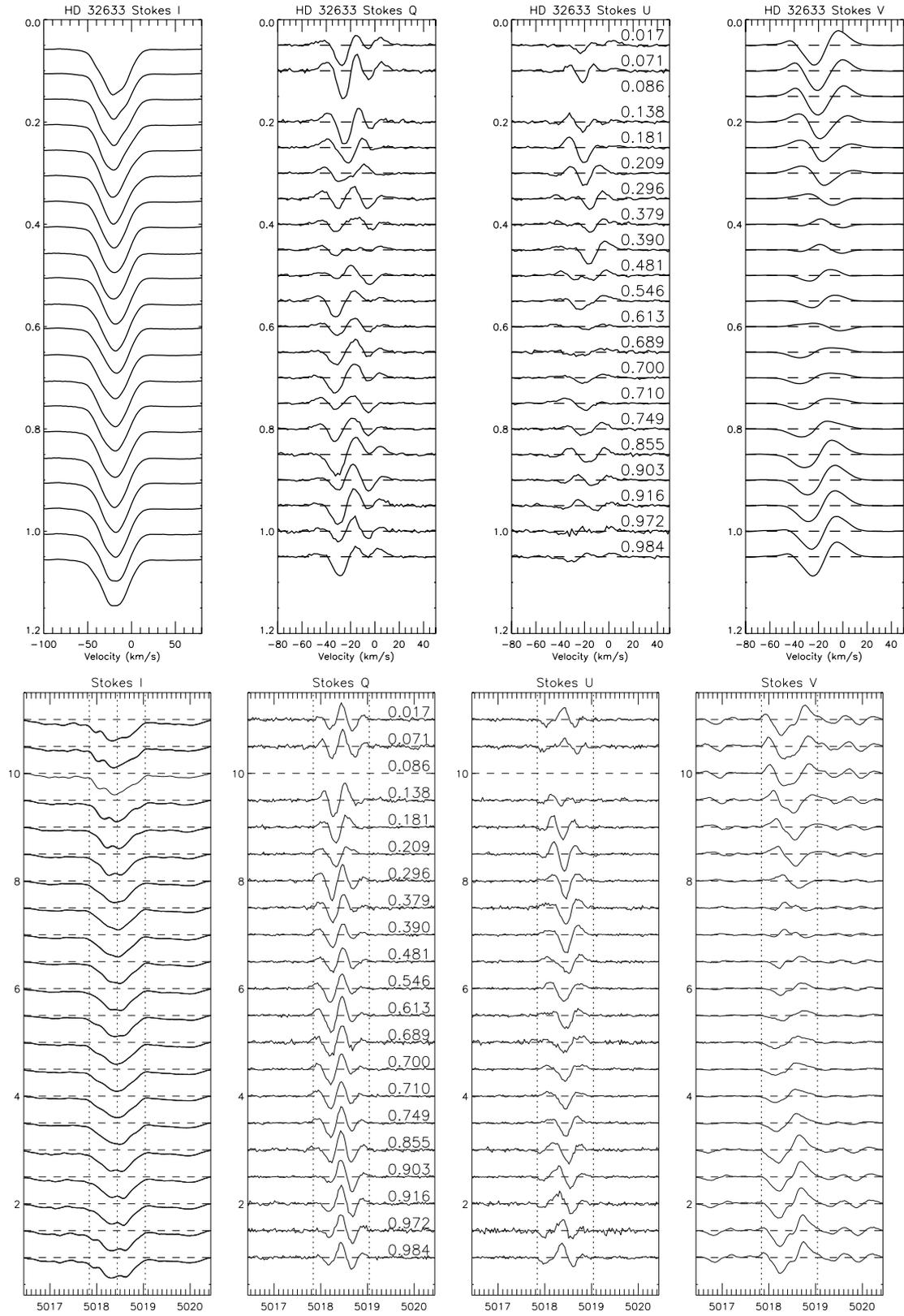
It was quickly seen that both the ESPaDOnS and NARVAL data were consistent with one another. An excellent example of this can be seen in Fig. 10, where two observations of HD 32633 are shown, one at a phase of 0.689 (taken with NARVAL) and one with a phase of 0.700 (taken with ESPaDOnS). The detailed agreement between the Stokes profiles indicates that both instruments appear to be performing extremely consistently with one another. Another test of



**Figure 12.** Longitudinal field measurements for HD 4778 obtained with ESPaDOnS/NARVAL (shown by filled circles) compared with those obtained by Bohlender (1989). A second-order Fourier fit to the ESPaDOnS/NARVAL data is given by the solid curve.



**Figure 13.** Net linear polarization (Stokes  $Q$  and  $U$ ) measurements for HD 4778 obtained with ESPaDOnS/NARVAL (filled circles, scaled 0.04) and those obtained by Leroy (1995, filled diamonds). The improvement in data quality is evidenced by the smaller error bars associated with the ESPaDOnS/NARVAL measurements.



**Figure 14.** Top: variation of Stokes  $I$ ,  $Q$ ,  $U$  and  $V$  LSD profiles for HD 32633. Rotational phase is represented on the vertical axis. A scaling factor of 20 was used for both the Stokes  $Q$  and  $U$  observations. Bottom: variation of Stokes  $I$ ,  $Q$ ,  $U$  and  $V$  profiles for HD 32633 in the 5018 Fe II line (scaled by a factor of 10 for Stokes  $Q$  and  $U$ , and 2 for Stokes  $V$ ).

consistency is described in Sections 6.2 and 6.6, where for certain phases, the ESPaDONs/NARVAL data for  $\alpha^2$  CVn and HD 32633 have been convolved to the same approximate resolution as MuSiCoS ( $R = 35\,000$ ) and then compared to MuSiCoS observations of a similar phase. Very good agreement can be seen in individual lines.

As proposed by Landi Degl’Innocenti et al. (1981), assumed by Landolfi et al. (1993) and Leroy et al. (1993, 1995a,b) and finally confirmed by Wade et al. (2000b), the phenomenon of broad-band or net linear polarization results from the differential saturation of  $\pi$  and  $\sigma$  Zeeman components of individual spectral lines. Broad-band polarization measurements, such as those reported by Leroy (1995), effectively average the signal from the net polarized spectral lines with regions of continuum (which may be unpolarized or polarized due to e.g. interstellar polarization). In contrast, net polarization measurements obtained using LSD measure only the net polarization in lines. Therefore, the measurements are not equivalent. In particular, Wade et al. (2000b) found that it was necessary to arbitrarily scale and shift the LSD measurements relative to the broad-band measurements in order to bring them into agreement. Here we adopt a similar procedure.

Table 2 also shows the reduced  $\chi^2$  of longitudinal field measurements taken from the null spectra. A value close to 1 indicates that the null spectra and their error bars are consistent with zero longitudinal field and that the instrument is performing as expected. We see that for HD 4778, HD 32633, 49 Cam and HD71866, this is not the case. These high values are a result of a handful of non-zero null measurements obtained during the 2006 December observations with NARVAL and very small number obtained with ESPaDONs, in which the null LSD profiles contained weak signatures (as discussed in Section 3).

In addition, for these stars, we find that the error bars on the Stokes  $V$  longitudinal field measurements are larger than the null longitudinal field error bars, typically by a factor of 3–4. This is significant, because it tells us that we begin to see that the Stokes  $V$  longitudinal field measurements are no longer limited by photon noise. Rather, the larger error bars for the longitudinal field in Stokes  $V$  tell us that the LSD model is failing to fit the  $V$  spectrum within the LIBRE-ESPRIT computed error bars, likely due to blends and the limitation of the weak-field approximation. This only happens when a significant number of individual lines show very strong Stokes  $V$  signatures, and (as demonstrated by our ability to fit the longitudinal field curves within the error bars) is fully compensated for in the error bar calculation for Stokes  $V$ . As discussed by Wade et al. (2000b) this compensation is achieved by scaling the photon noise statistical error bars of LSD profiles by a factor equal to the square root of the reduced  $\chi^2$ . Wade et al. (2000b) describes that this scaling is almost always required for Stokes  $V$  in Ap stars. If we recompute the reduced  $\chi^2$  calculation for the null spectra, but use the Stokes  $V$  error bars instead of the null error bars, the resulting reduced  $\chi^2$  is consistent with a zero longitudinal field (for example in the case of HD 32633, the reduced  $\chi^2$  goes from 5.42 to 0.36).

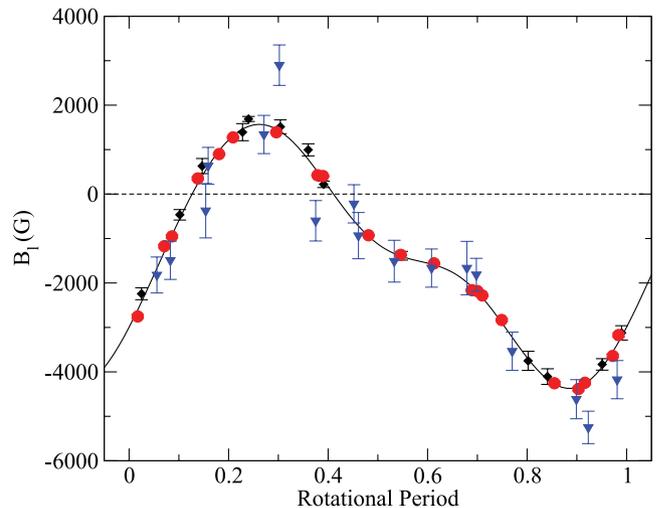
## 6 RESULTS FOR INDIVIDUAL STARS

In the following sections, the results for each star will be discussed, including figures showing the Stokes profile variation both in selected individual lines and in the LSD profiles. Also the variation in the longitudinal field and in some cases the variation in net linear polarization as a function of phase will be shown for each target. Results are compared with those of Wade et al. (2000a) and Leroy (1995) where possible. In the cases where the net linear polarization

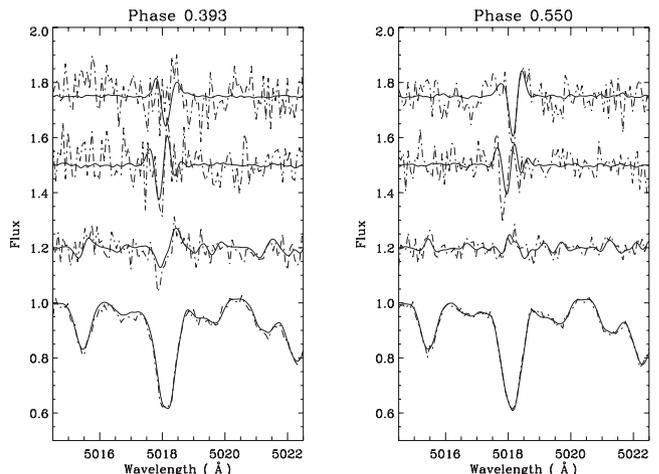
was close to or consistent with zero at all phases (such as HD 32633 and HD 40312), these plots have been omitted.

### 6.1 HD 4778

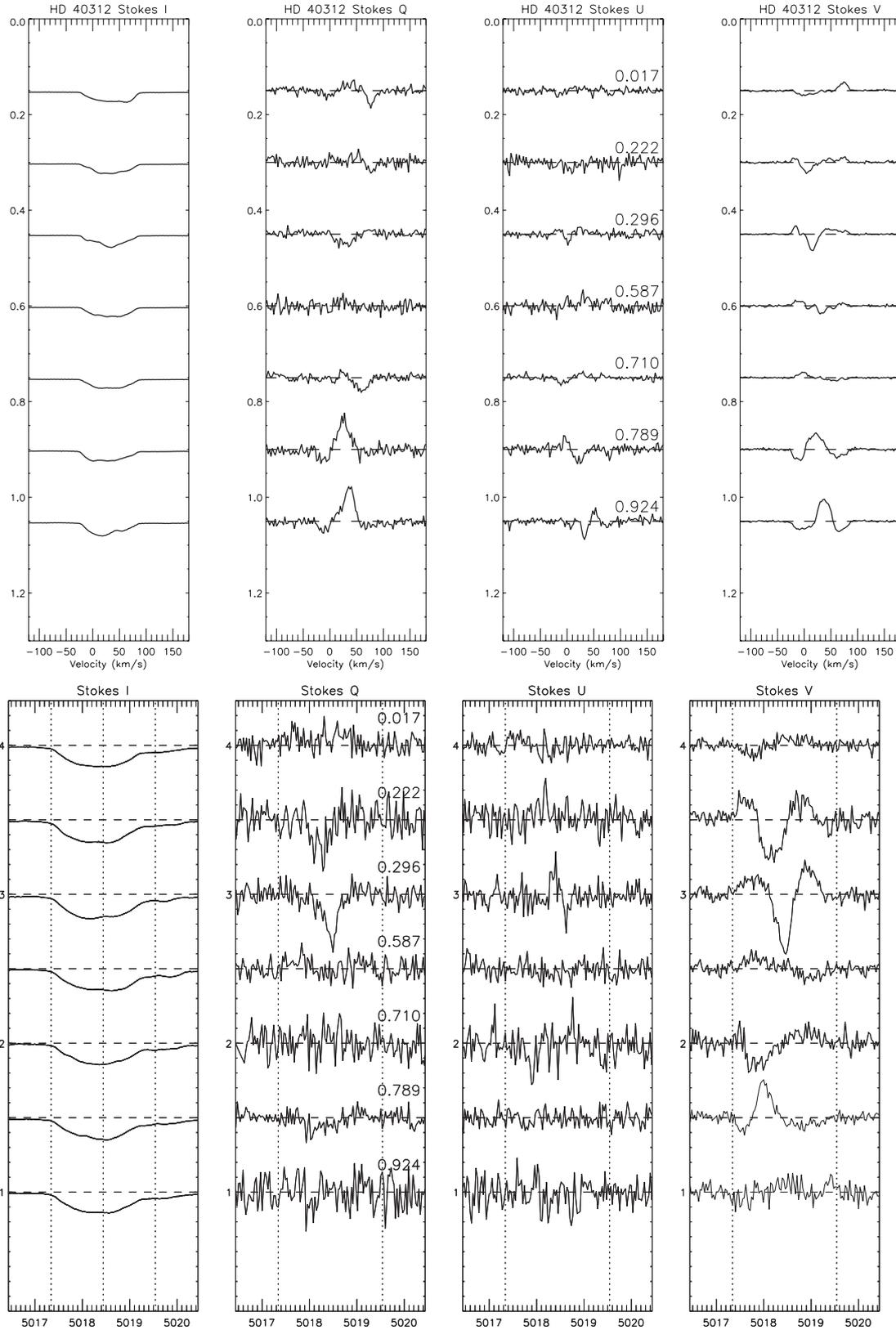
HD 4778 is the only star in our sample that was not studied by Wade et al. (2000a,b). It was classified by Renson & Manfroid (2009) as an A1CrSrEu star. We obtained Stokes  $V$ ,  $Q$  and  $U$  spectra for seven rotational phases (although no Stokes  $Q$  observation was obtained at phase 0.366). Using the method described in Section 5, the projected rotational velocity ( $v \sin i$ ) was determined to be  $36 \pm 2 \text{ km s}^{-1}$  which is little larger than the Abt & Morrell (1995) value of  $33 \text{ km s}^{-1}$ . The data have been phased according to the ephemeris



**Figure 15.** Longitudinal field measurements for HD 32633 obtained with ESPaDONs/NARVAL (shown by filled circles) compared with those obtained by Wade et al. (2000b) with MuSiCoS (shown by filled diamonds) and those obtained by Borra & Landstreet (1980, shown as triangles). Good agreement can be seen, confirming consistency between the instruments, and the improvement in data quality is evidenced by the smaller error bars associated with the ESPaDONs/NARVAL measurements. A third-order Fourier fit to the ESPaDONs/NARVAL data is given by the solid curve.



**Figure 16.** Comparison between two phases of the ESPaDONs/NARVAL observations convolved to a resolution of 35 000 and the MuSiCoS data at a similar phase for HD 32633 (MuSiCoS data: dashed lines; ESPaDONs/NARVAL: solid lines). Good overall agreement can be seen between the two epochs of data.



**Figure 17.** Top: variation of Stokes  $I$ ,  $Q$ ,  $U$  and  $V$  LSD profiles for HD 40312. Rotational phase is represented on the vertical axis. A scaling factor of 200 was used for both the Stokes  $Q$  and  $U$  observations, while a factor of 50 was used for Stokes  $V$ . Bottom: variation of Stokes  $I$ ,  $Q$ ,  $U$  and  $V$  profiles for HD 40312 in the Fe II 5018 line (scaled by a factor of 80 for Stokes  $Q$  and  $U$ , and 40 for Stokes  $V$ ).

reported by Leone et al. (2000):

$$\text{JD} = 244\,6674.006 + 2^d 561\,71 \cdot \text{E}. \quad (3)$$

The individual spectral lines show strong signatures in Stokes  $V$ , and although variation in Stokes  $Q$  and  $U$  can be seen in individual lines in this star (an example is shown for the Fe II  $\lambda 5018$  in Fig. 11), the signatures in Stokes  $Q$  and  $U$  are relatively weak. Signatures are more prominent in the LSD profiles shown in Fig. 11. Stokes  $I$  shows variability in the Fe II 4923, 5169 and 5018 Å lines (as shown in Fig. 11).

Both  $B_\ell$  and net linear polarization measurements have been obtained for each phase and are reported in Table 4. The average uncertainty on the longitudinal field measurement is 20 G. The LSD profiles are illustrated in Fig. 11, and the longitudinal field curve is shown in Fig. 12. The second-order Fourier fit to this curve, using only the new measurements, gives a reduced  $\chi^2$  of 1.41. Also included in Fig. 12 are measurements reported by Bohlender (1989) for comparison. The slight apparent phase shift could be due to a cumulative error in phase over an almost 30 yr span caused by the period uncertainty.

Leroy et al. (1995) obtained broad-band linear polarization measurements of HD 4778; their values are compared with the measurements obtained in this work in Fig. 13. Reasonable agreement in Stokes  $Q$  and  $U$  can be seen between the two epochs of data. The ESPaDONs/NARVAL data had their mean subtracted and then were scaled by 0.08 to bring them into agreement with those of Leroy et al. (1995).

## 6.2 HD 32633

HD 32633 is a fairly broad-lined B9p star with a strong, non-sinusoidal longitudinal magnetic field variation. We obtained Stokes  $V$ ,  $Q$  and  $U$  profiles for 20 rotational phases (Stokes  $Q$  and  $U$  observations are missing at phase 0.086). Using the method described in Section 5, the projected rotational velocity ( $v \sin i$ ) was determined to be  $19 \pm 2 \text{ km s}^{-1}$ , which agrees within uncertainty with the value adopted by Wade et al. (2000b).

The data have been phased according to the ephemeris of Adelman (1997b):

$$\text{JD} = 243\,7635.200 + 6^d 430\,00 \cdot \text{E}. \quad (4)$$

Very strong and variable Stokes  $Q$  and  $U$  polarization signatures can be seen in individual lines in the spectra of this star, an example of which is shown for the Fe II 5018 Å line in Fig. 14. Interestingly, the shapes of the Stokes  $Q$  and  $U$  profiles change relatively little as a function of phase (in individual lines, as well as LSD profiles of Fig. 14). This could be an indication that we are ‘seeing’ a limited part of the transverse geometry of the field due to the overall geometry of the star. It should be noted in Fig. 14 that the shapes of the Stokes  $Q$  and  $U$  profiles of individual Fe lines look quite different from the LSD profiles. This may be a result of differences between the Stokes profiles of weak versus strong lines, or possibly differences in the Stokes profiles of lines of different chemical elements. As with all the stars with strong linear polarization signatures in this study, HD 32633 has many more individual lines showing signatures in the new observations than in the MuSiCoS observations. Also Stokes  $I$  shows variability in lines such as in the Cr II 4588 Å line and Fe II 4824 and 5018 Å lines (as shown in Fig. 14).

Both  $B_\ell$  and net linear polarization measurements have been obtained for each phase and are reported in Table 4. The longitudinal field values obtained by Borra & Landstreet (1980) and Wade et al. (2000b) are compared to the new measurements in Fig. 15. Very

good agreement can be seen between the three epochs of data, with the clear reduction in uncertainties shown by the smaller error bars associated with the new data. The reduced  $\chi^2$  of the Fourier fit to the field curve is 1.95, with an order of fit of 3. The average uncertainty on the longitudinal field measurements is 26 G. The net linear polarization plot is not shown for HD 32633, due to the fact that the measurements show little or no variation and are consistently null at all phases. In addition to certain phases, the ESPaDONs/NARVAL data for HD 32633 have been convolved to the same approximate resolution as MuSiCoS ( $R = 35000$ ) and then compared to MuSiCoS observations of a similar phase. Within the limits imposed by noise, very good agreement can be seen in individual lines between the two data sets in Fig. 16.

## 6.3 HD 40312 – $\theta$ Aur

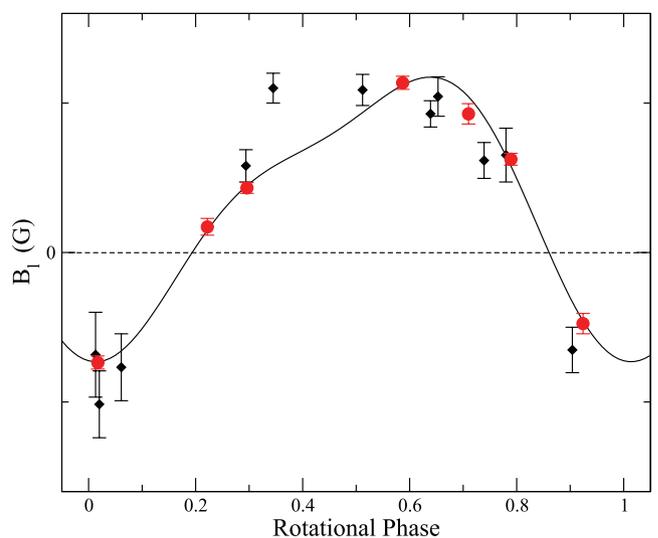
$\theta$  Aur is a broader lined A0p star with a weak magnetic field. We obtained Stokes  $V$ ,  $Q$  and  $U$  profiles for seven rotational phases. Using the method as described above, the projected rotational velocity ( $v \sin i$ ) was determined to be  $53 \pm 1 \text{ km s}^{-1}$ , which agrees within uncertainty with the value adopted by Wade et al. (2000b).

All magnetic measurements have been phased according to the ephemeris of Wade et al. (2000a):

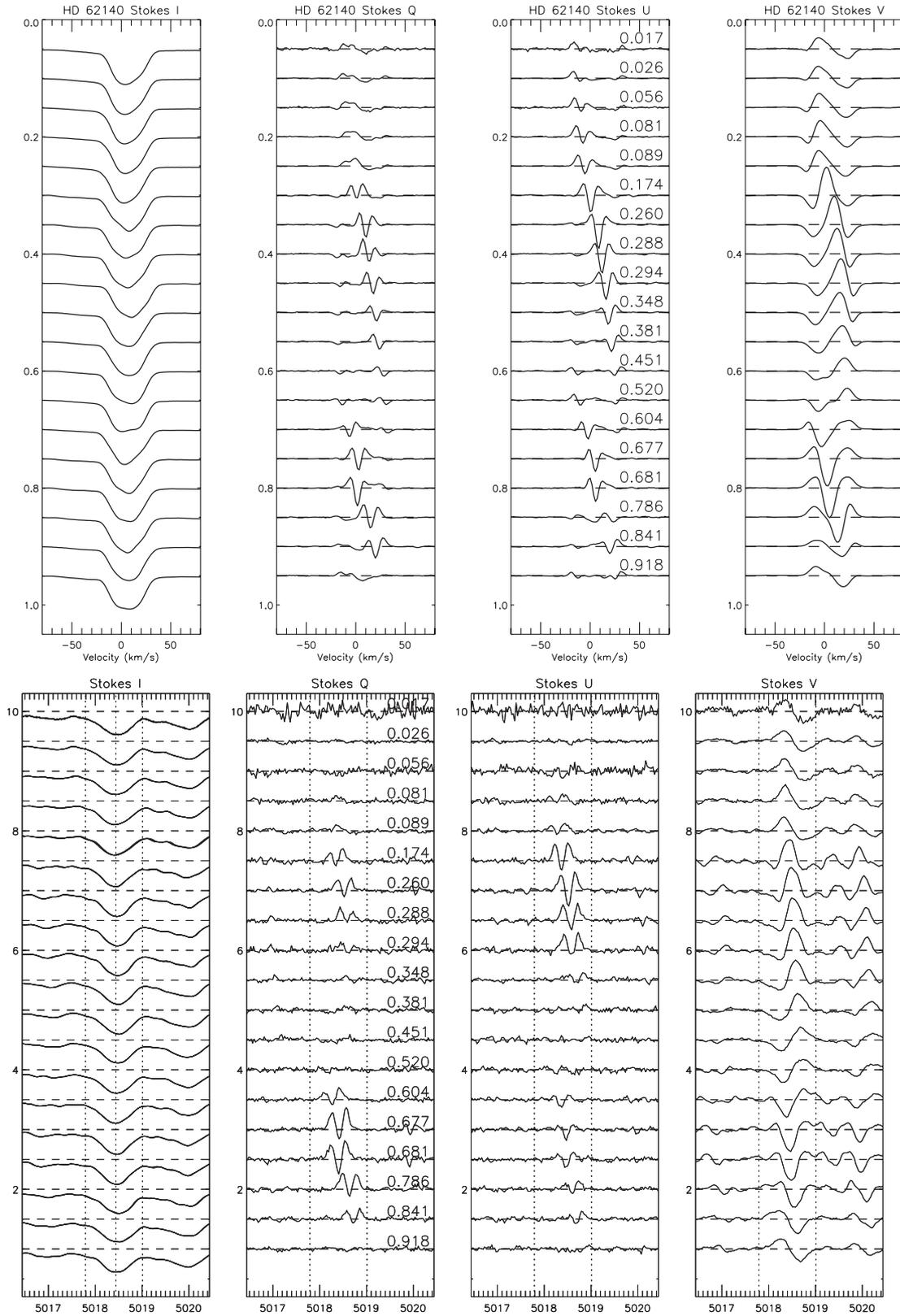
$$\text{JD} = 245\,0001.881 + 3^d 61860 \cdot \text{E}. \quad (5)$$

The individual line profiles show weak signatures in Stokes  $V$ ; marginal Stokes  $Q$  and  $U$  signatures are visible at phase 0.7–0.8 (an example for the Fe II 5018 Å line is shown in Fig. 17). Signatures are slightly more prominent in the LSD profiles. Stokes  $I$  exhibits small variations in the Fe II 5018 Å line and very subtle variation in the Fe II 4923 Å line. Because of the weak Stokes  $Q$  and  $U$  signatures in individual lines in the spectra of this star, it would be a challenging candidate for  $IQUV$  mapping, but would be very suitable for  $IV$  mapping.

Both  $B_\ell$  and net linear polarization measurements have been obtained for each phase and are reported in Table 4. The longitudinal field values obtained by Wade et al. (2000b) are compared to this



**Figure 18.** Longitudinal field measurements for HD 40312 obtained with ESPaDONs/NARVAL (shown by filled circles) compared with those obtained by Wade et al. (2000b) with MuSiCoS (shown by filled diamonds). A second-order Fourier fit to the ESPaDONs/NARVAL data is given by the solid curve.



**Figure 19.** Top: variation of Stokes  $I$ ,  $Q$ ,  $U$  and  $V$  LSD profiles for HD 62140. Rotational phase represented on the vertical axis. A scaling factor of 12 was used for both the Stokes  $Q$  and  $U$  observations, and a scaling factor of 2 was used for Stokes  $V$ . Bottom: variation of Stokes  $I$ ,  $Q$ ,  $U$  and  $V$  profiles for HD 62140 in the 5018 Fe II line (scaled by a factor of 12 for Stokes  $Q$  and  $U$ , and 4 for Stokes  $V$ ).

work in Fig. 18. Agreement can be seen between the two epochs of data. The average uncertainty on the new longitudinal field measurement is 14 G. A Fourier fit to the longitudinal field curve yields a reduced  $\chi^2$  of 1.53 with an order of fit of 2. Because of the limited phase coverage and the small net linear polarization measurements, a plot of net linear polarization is not shown for HD 40312.

#### 6.4 HD 62140 – 49 Cam

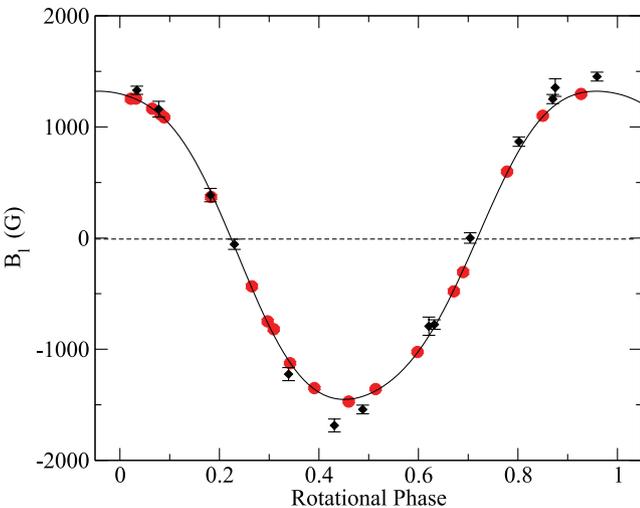
49 Cam is a fairly broad-lined F0p star with a moderately strong magnetic field. We obtained Stokes  $V$ ,  $Q$  and  $U$  profiles for 19 rotational phases. The projected rotational velocity ( $v \sin i$ ) was determined to be  $24 \pm 2 \text{ km s}^{-1}$ , which agrees within uncertainty with the value adopted by Wade et al. (2000b).

All measurements have been phased according to the ephemeris of Adelman (1997a):

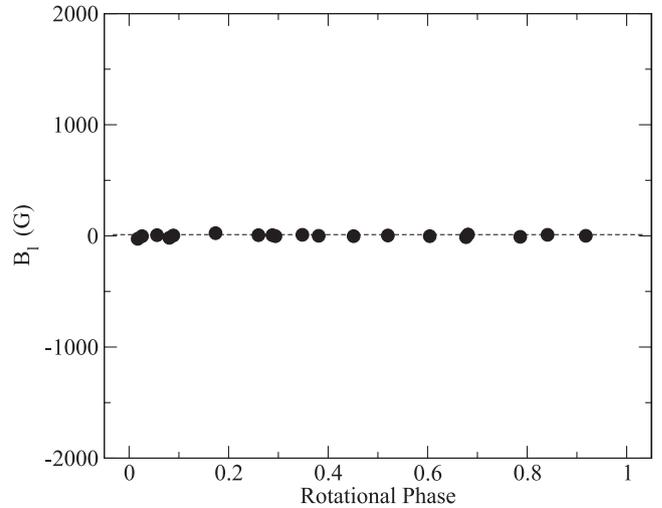
$$\text{JD} = 244\,1257.300 + 4^d 28679 \cdot E. \quad (6)$$

Variable signatures can be seen in Stokes  $Q$  and  $U$  in individual lines in this star, an example of which is shown for the Fe II 5018 Å line in Fig. 19. The signatures are even clearer in the LSD profiles (Fig. 19). The amplitudes of the Stokes  $Q$  and  $U$  profiles are small compared to HD 32633 (about 50 per cent smaller), in both the individual lines and LSD profiles. In Stokes  $I$ , 49 Cam shows subtle variability in most lines, but as with all the stars with strong linear polarization signatures in this study, many individual lines show signatures, as shown in Fig. 8.

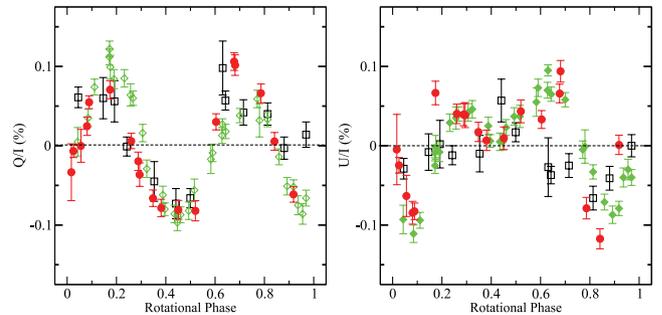
Both  $B_\ell$  and net linear polarization measurements have been obtained for each phase and are reported in Table 4. The longitudinal field values obtained by Wade et al. (2000b) are compared to those derived in this work in Fig. 20. Very good agreement can be seen between the two epochs of data. The average uncertainty on the new longitudinal field measurement is 14 G, and the Fourier fit to the new measurements gives a reduced  $\chi^2$  of 1.29 with an order of fit of 3. To further illustrate the quality of data, Fig. 21 shows the longitudinal measurements obtained from the null spectrum.



**Figure 20.** Longitudinal field measurements for HD 62140 obtained with ESPaDOnS/NARVAL (shown by filled circles) compared with those obtained by Wade et al. (2000b) with MuSiCoS (shown by filled diamonds). Good agreement can be seen, confirming consistency between the instruments, and the improvement in data quality is evidenced by the smaller error bars associated with the ESPaDOnS/NARVAL measurements. A third-order Fourier fit to the ESPaDOnS/NARVAL data is given by the solid curve.



**Figure 21.** Longitudinal field measurements for HD 62140 obtained with ESPaDOnS/NARVAL for the null spectrum.



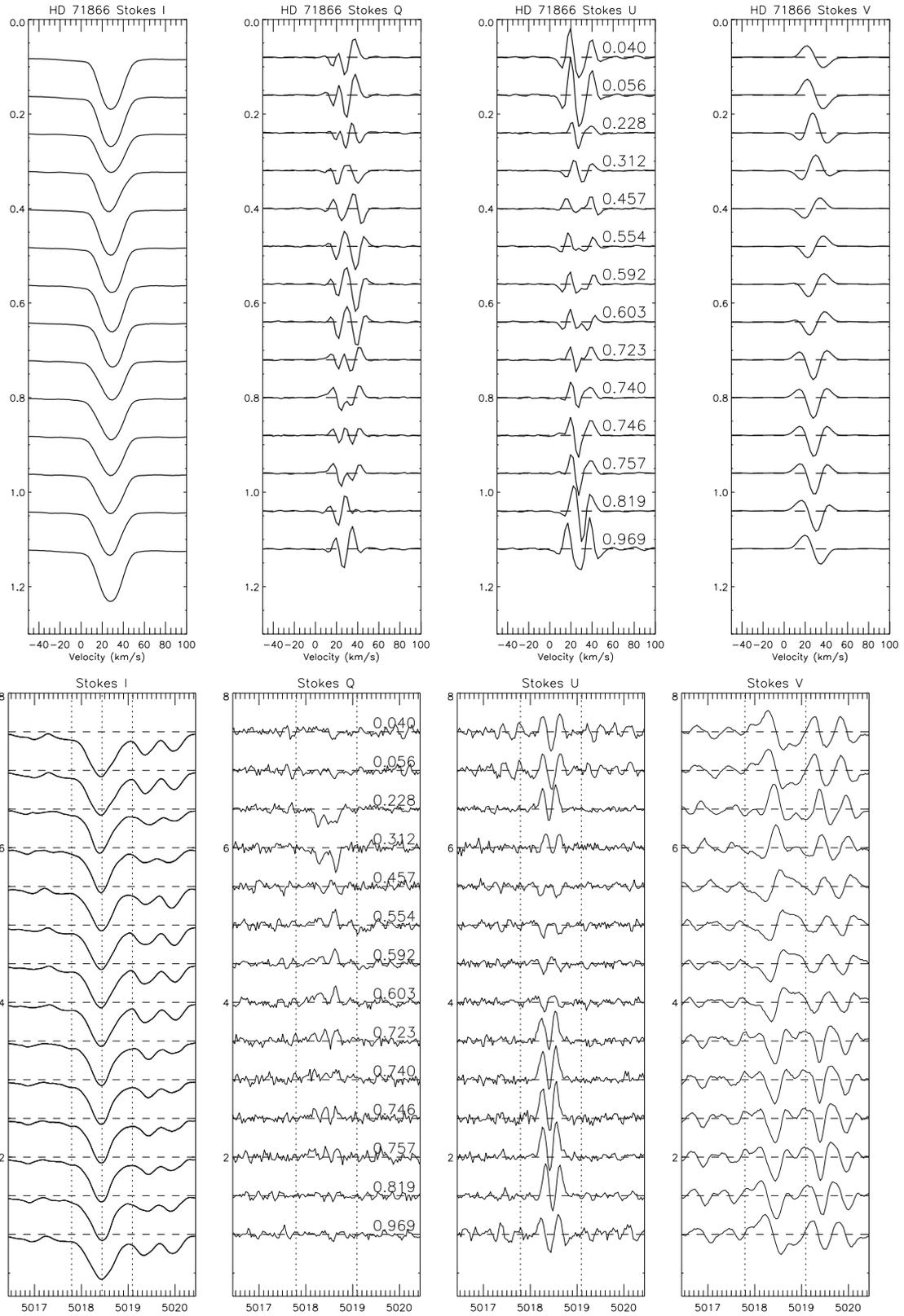
**Figure 22.** Net linear polarization measurements (Stokes  $Q$  and  $U$ ) for 49 Cam obtained with ESPaDOnS/NARVAL (shown by filled circles, scaled by a factor of 0.28) compared with those obtained by Wade et al. (2000b) with MuSiCoS (shown by filled diamonds) and those obtained by Leroy (1995) (shown by open squares). Good agreement can be seen, confirming consistency between the instruments, and the improvement in data quality is evidenced by the smaller error bars associated with the ESPaDOnS/NARVAL measurements.

In the case of data free from any spurious signals (e.g. caused by instrumental polarization effects), the longitudinal field in the null spectrum should be consistent with zero, and this is clearly the case with HD 62140 and is representative of most observations in this sample.

The net linear polarization as a function of phase is shown in Fig. 22. For both Stokes  $Q$  and  $U$  measurements, there is reasonable agreement at most phases with the measurements reported by Wade et al. (2000b) and in most cases very good agreement with values reported by Leroy et al. (1995). The net linear polarization was scaled by the same factor (0.28), double the scaling factor employed by Wade et al. (2000b). In addition, the mean was subtracted from all measurements, as prescribed by Wade et al. (2000b).

#### 6.5 HD 71866

HD 71866 is a sharp-lined A2p star with a moderately strong magnetic field. We have obtained Stokes  $V$ ,  $Q$  and  $U$  profiles for 14 rotational phases. The projected rotational velocity ( $v \sin i$ ) we



**Figure 23.** Top: variation of Stokes  $I$ ,  $Q$ ,  $U$  and  $V$  LSD profiles for HD 71866. Rotational phase represented on the vertical axis. A scaling factor of 15 was used for both the Stokes  $Q$  and  $U$  observations, and a scaling factor of 1 was used for Stokes  $V$ . Bottom: variation of Stokes  $I$ ,  $Q$ ,  $U$  and  $V$  profiles for HD 71866 in the 5018 Fe II line (scaled by a factor of 16 for Stokes  $Q$ , 16 for Stokes  $U$  and 4 for Stokes  $V$ ).

determine is  $15 \pm 2 \text{ km s}^{-1}$ , which agrees within uncertainty with the value adopted by Wade et al. (2000b).

All measurements have been phased according to the ephemeris of Bagnulo et al. (1995):

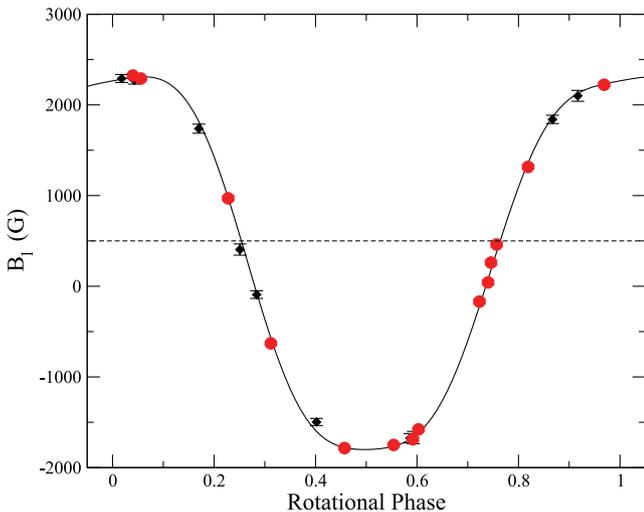
$$\text{JD} = 243\,8297.5 + 6^{\text{d}}800\,22 \cdot \text{E}. \quad (7)$$

Clear signatures and strong variability can be seen in Stokes  $Q$  and  $U$  in the individual lines of this star, an example of which is shown for the Fe II 5018 Å line in Fig. 23. The amplitudes of the Stokes  $Q$  and  $U$  profiles are large compared to many of the other targets in this study (in both the individual lines and LSD profiles, see Fig. 23). Stokes  $I$  appears to vary very slightly with phase, with only small changes in line shape in the Fe II lines at 4923, 5018 and 5169 Å.

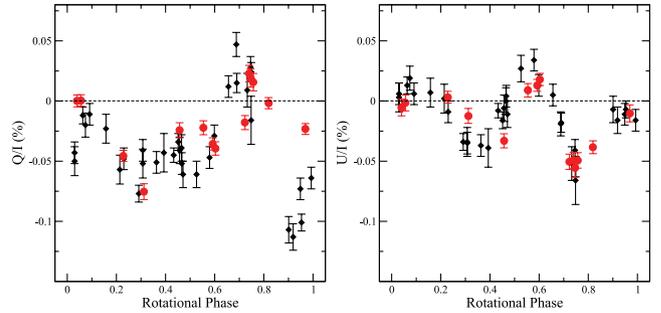
Both  $B_{\ell}$  and net linear polarization measurements have been obtained for each phase and are reported in Table 4. To get good agreement between the two epochs of data required a re-analysis of the Wade et al. (2000b) data, using the updated and abundance-specific line mask (as used for the ESPaDOnS/NARVAL data). It is likely that this star is in a temperature regime where measurements are very sensitive to the line mask chosen or because of peculiar abundances. The longitudinal field variation is illustrated in Fig. 24. Good agreement can be seen between the two epochs of data. The average uncertainty on the longitudinal field measured from the ESPaDOnS and NARVAL spectra was 27 G. A Fourier fit to the longitudinal field curve yields a  $\chi^2$  of 1.02, with order of fit of 4.

The net linear polarization as a function of phase is shown in Fig. 25. The Stokes  $Q$  and  $U$  measurements obtained in this study have been scaled by a factor of 0.07 to bring them into agreement with the measurements of Leroy et al. (1995), and this scaling is consistent with that of Wade et al. (2000b). There is approximate agreement at most phases.

HD 71866 is an ideal candidate for MDI. However, more observations will be required to complete phase coverage. As with all the



**Figure 24.** Longitudinal field measurements for HD 71866 obtained with ESPaDOnS/NARVAL (shown by filled circles) compared with those obtained by Wade et al. (2000b) with MuSiCoS (shown by filled diamonds). Good agreement can be seen (with a re-analysis of the MuSiCoS observations), confirming consistency between the instruments, and the improvement in data quality is evidenced by the smaller error bars associated with the ESPaDOnS/NARVAL measurements. A fourth-order Fourier fit to the ESPaDOnS/NARVAL data is given by the solid curve.



**Figure 25.** Net linear polarization (Stokes  $Q$  and  $U$ ) measurements for HD 71866 obtained with ESPaDOnS/NARVAL (shown by filled circles and scaled by a factor of 0.10) compared with those of Leroy et al. (1995, shown by filled diamonds).

stars with strong linear polarization in this study, many individual lines show clear signatures.

### 6.6 HD 112413 – $\alpha^2$ CVn

$\alpha^2$  CVn is a fairly sharp lined A0p star with a moderately strong magnetic field. We obtained Stokes  $V$ ,  $Q$  and  $U$  profiles at 24 rotational phases. The projected rotational velocity ( $v \sin i$ ) was determined to be  $17 \pm 1 \text{ km s}^{-1}$ , which agrees within uncertainty with the values adopted by Wade et al. (2000b) and Kochukhov & Wade (2010).

All measurements have been phased according to the ephemeris of Farnsworth (1932):

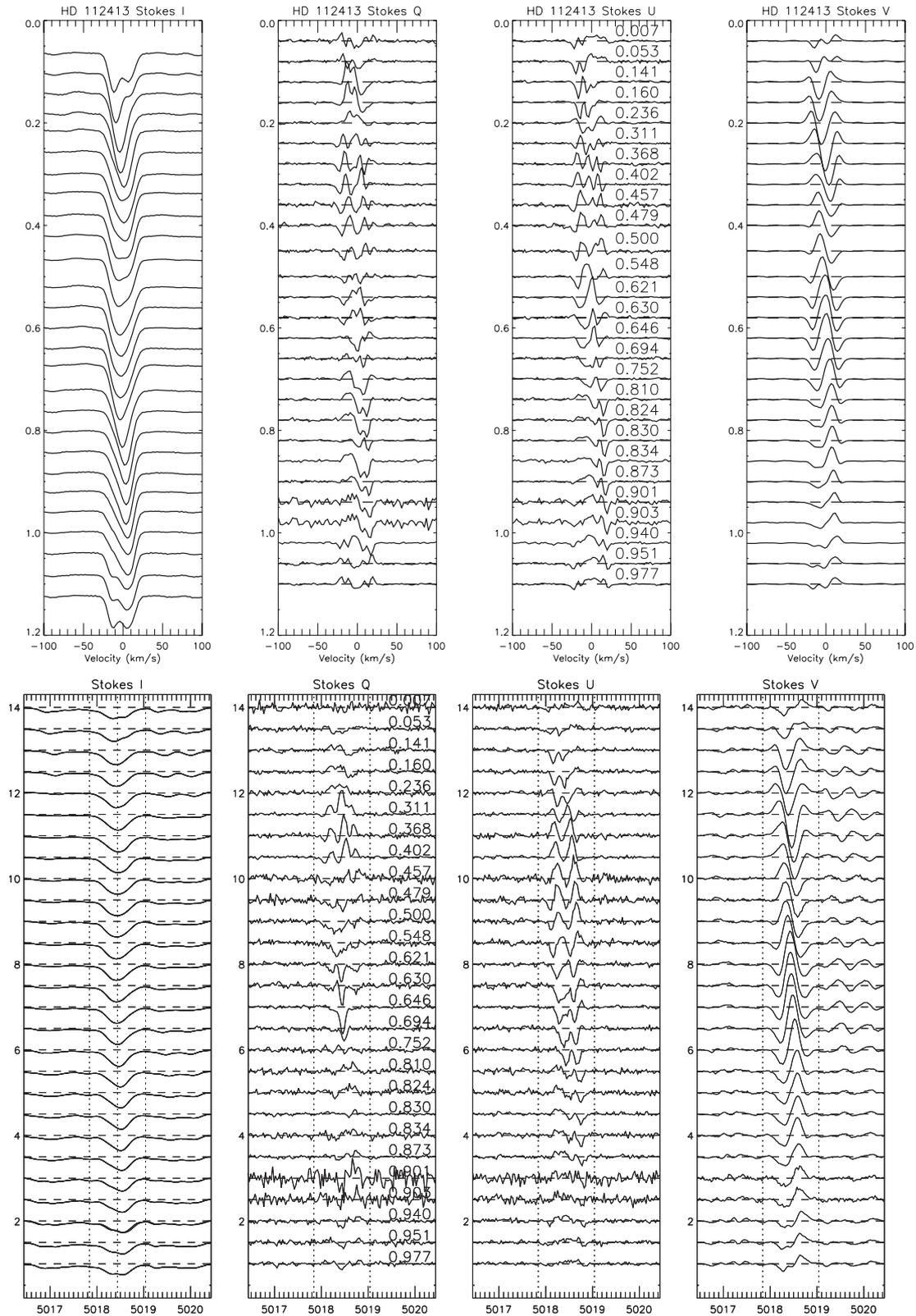
$$\text{JD} = 241\,9869.720 + 5^{\text{d}}469\,39 \cdot \text{E}. \quad (8)$$

Clear signatures and strong variability can be seen in Stokes  $Q$  and  $U$  in individual spectral lines of this star. Profile variations of the Fe II  $\lambda$  5018 line and LSD profiles are shown in Fig. 26. Very strong variation in Stokes  $I$  can be seen in the Fe II lines at 4923, 5018 and 5169 Å. Both  $B_{\ell}$  and net linear polarization measurements have been obtained for each phase and are reported in Table 4. The longitudinal field values are compared to the results of Wade et al. (2000b) in Fig. 27. To get good agreement between the two epochs of data required a re-analysis of the Wade et al. (2000b) data, using the updated and abundance-specific line mask (as used for the ESPaDOnS/NARVAL data). It is likely that this star is in a temperature regime where measurements are more sensitive to the line mask chosen or that the abundances are sufficiently peculiar to cause a difference. In addition, the new measurements are much more precise. The average uncertainty on the longitudinal field measurement was 27 G, with a Fourier fit to the longitudinal field curve yielding a reduced  $\chi^2$  of 0.83, with an order of fit of 3.

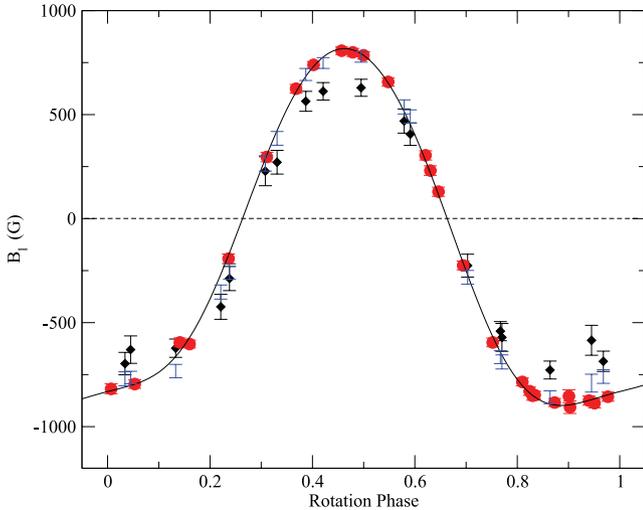
The net linear polarization as a function of phase is shown in Fig. 28. Both Stokes  $Q$  and  $U$  measurements were compared with the measurements of Wade et al. (2000b), showing very good agreement at most phases. For certain phases, the ESPaDOnS/NARVAL data for  $\alpha^2$  CVn have been convolved to the same approximate resolution as MuSiCoS ( $R = 35\,000$ ) and then compared to MuSiCoS, as described for HD 32633. Again good agreement can be seen in individual lines between the two data sets in Fig. 29.

### 6.7 HD 118022 – 78 Vir

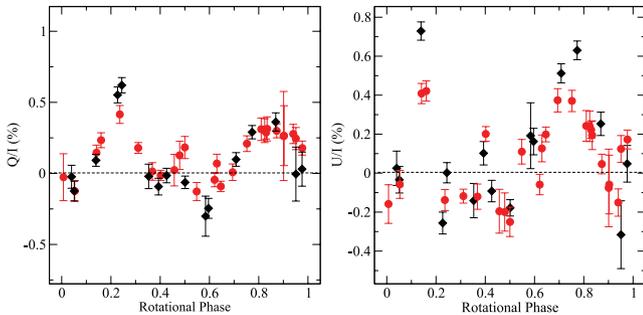
78 Vir is a sharp-lined A1p star with a moderately strong magnetic field. We obtained Stokes  $V$ ,  $Q$  and  $U$  profiles for five rotational



**Figure 26.** Top: variation of Stokes  $I$ ,  $Q$ ,  $U$  and  $V$  LSD profiles for HD 112413. Rotational phase is represented on the vertical axis. A scaling factor of 25 was used for both the Stokes  $Q$  and  $U$  observations, and a scaling factor of 2 was used for Stokes  $V$ . Bottom: variation of Stokes  $I$ ,  $Q$ ,  $U$  and  $V$  profiles of HD 112413 in the Fe II 5018 line (scaled by a factor of 30 for Stokes  $Q$  and  $U$ , and 7 for Stokes  $V$ ).



**Figure 27.** Longitudinal field measurements for HD 112413 obtained with ESPaDOnS/NARVAL (shown by filled circles) compared with those obtained by Wade et al. (2000b) with MuSiCoS (shown by filled diamonds). The measurements shown without symbols are the result of a re-analysis of the original Wade et al. (2000a) observations with the new line mask. The improvement in data quality is evidenced by the smaller error bars associated with the ESPaDOnS/NARVAL measurements. A fourth-order Fourier fit to the ESPaDOnS/NARVAL data is given by the solid curve.



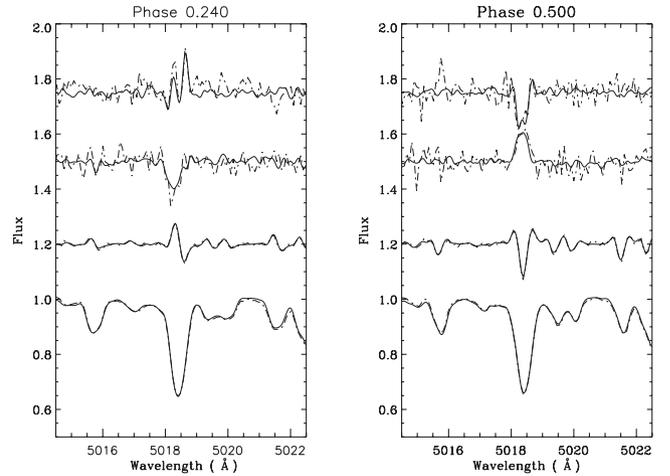
**Figure 28.** Net linear polarization (Stokes  $Q$  and  $U$ ) measurements for HD 112413 obtained with ESPaDOnS/NARVAL (shown by filled circles) compared with those obtained by Wade et al. (2000b) with MuSiCoS (shown by filled diamonds). Good agreement can be seen, confirming consistency between the instruments, and the improvement in data quality is evidenced by the smaller error bars associated with the ESPaDOnS/NARVAL measurements.

phases. The projected rotational velocity ( $v \sin i$ ) was determined to be  $13 \pm 1 \text{ km s}^{-1}$ , which agrees within uncertainty with the value adopted by Wade et al. (2000b). All measurements have been phased according to the ephemeris of Preston (1969):

$$\text{JD} = 243\,4816.90 + 3^d 7220 \cdot E. \quad (9)$$

Clear signatures and strong variability can be seen in Stokes  $Q$  and  $U$  in the individual lines of this star, as shown for the Fe II  $\lambda$  5018 line and LSD profiles in Fig. 30. Stokes  $I$  appears to vary only slightly between phases, with only small changes in the Fe II lines at 4923, 5018 and 5169 Å.

HD 118022 was studied by Khalack & Wade (2006) who constrained the global magnetic field of the star and determined the abundance distributions of titanium and chromium. This was performed using the magnetic charge distribution method (Gerth,



**Figure 29.** Comparison between two phases of the ESPaDOnS/NARVAL observations convolved to a resolution of  $R = 35\,000$  and the MuSiCoS data at a similar phase for HD 112413 (MuSiCoS data: dashed lines; ESPaDOnS/NARVAL: solid lines).

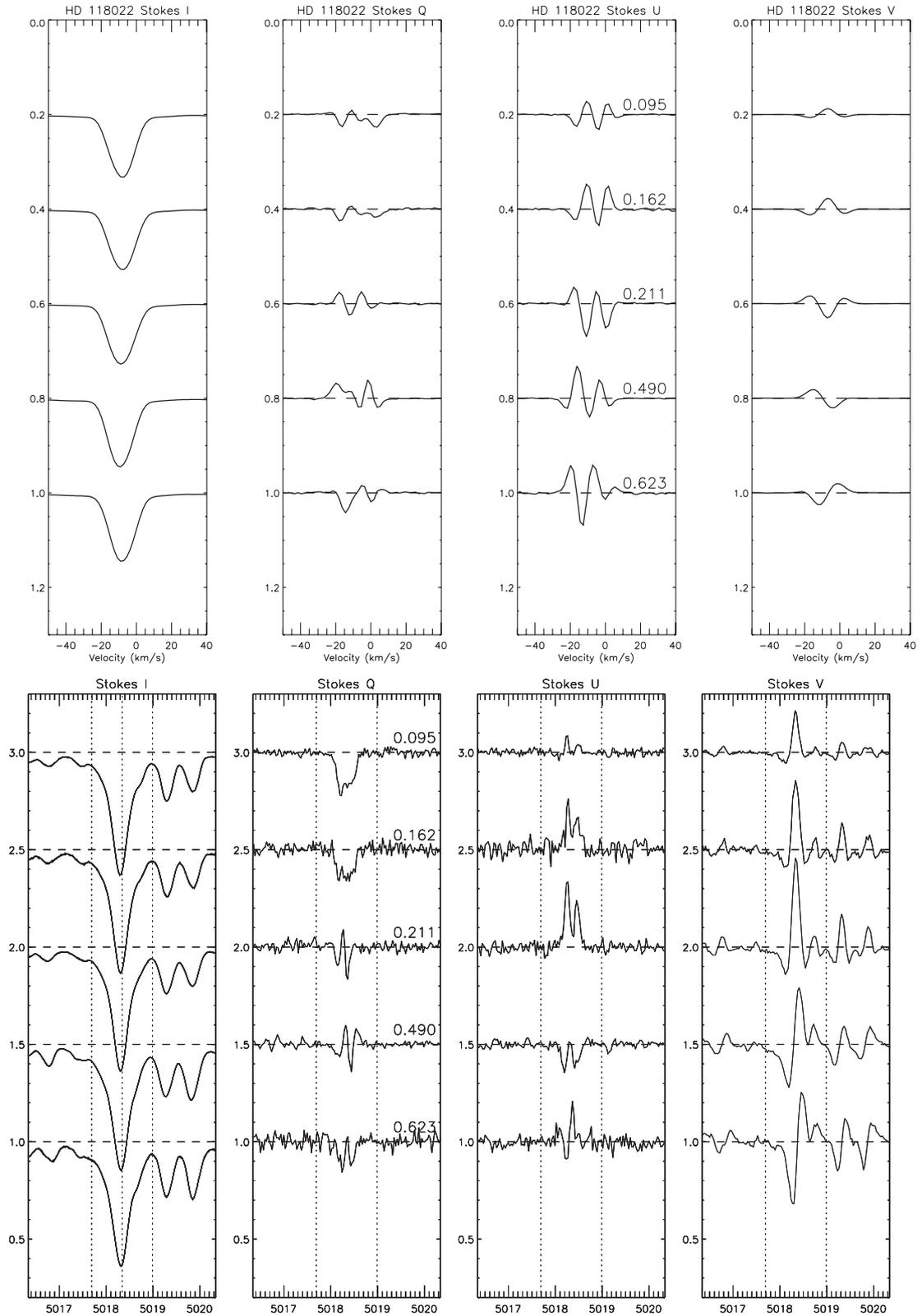
Glagolevskij & Scholz 1997; Khalack et al. 2001) and lower resolution MuSiCoS spectra, which limit the number of lines that could be modelled. This star is an excellent target for MDI, and further observations are warranted to supplement the currently rather sparse phase coverage.

Both  $B_l$  and net linear polarization measurements have been obtained for each phase and are reported in Table 4. The longitudinal field values obtained by Wade et al. (2000b) are compared to this work in Fig. 31. The average uncertainty on the longitudinal field measurements was 12 G. The net linear polarization as a function of phase is shown in Fig. 32 and compared with the measurements of Wade et al. (2000b) and Leroy (1995), with good agreement achieved at most phases. The net linear polarization measurements were scaled by the same factor (0.10) as reported by Wade et al. (2000b).

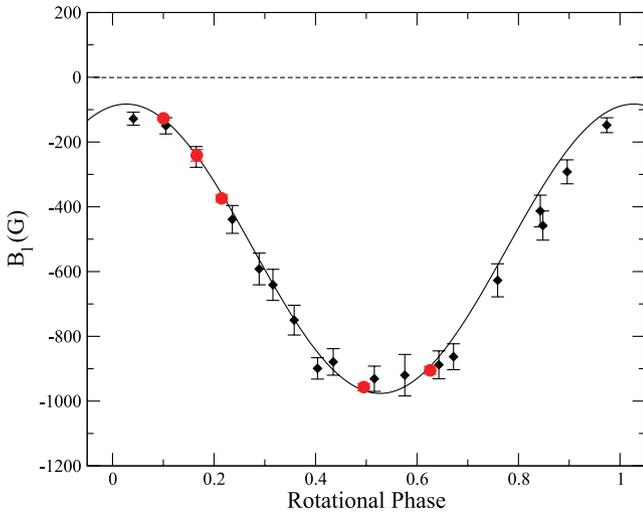
## 7 DISCUSSION AND CONCLUSIONS

The goal of this project was to obtain a new data set in all four Stokes parameters for a selection of well-studied Ap stars, with the ultimate aim to map these stars using MDI. The target list contained stars which span a large part of the parameter space of interest, with sufficient S/N to not only greatly improve on the previous observations, but to also be suitable for MDI mapping. The final selection was based primarily on stars already identified by Wade et al. (2000a) as promising candidates for such study.

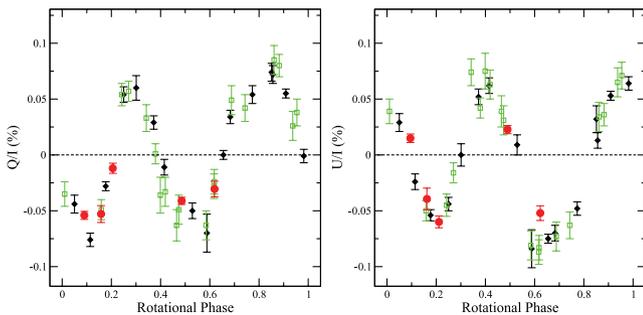
Early on in the project it was clear that both ESPaDOnS and NARVAL have greatly improved the level of detail at which Ap stars can be studied. The resulting data set obtained for this study is far superior to that obtained previously with MuSiCoS and represents some of the highest resolution phase-resolved observations of Ap stars acquired to date. This data set has more individual lines showing variation, much improved S/N and smaller error bars associated with measurements of the longitudinal field and net linear polarization. The new data have been shown to be consistent with the previous observations of Wade et al. (2000a) and also with those of Leroy et al. (1995), with most targets agreeing well between the different epochs.



**Figure 30.** Top: variation of Stokes  $I$ ,  $Q$ ,  $U$  and  $V$  LSD profiles for HD 118022. Rotational phase represented on the vertical axis. A scaling factor of 25 was used for both the Stokes  $Q$  and  $U$  observations, and no scaling on Stokes  $V$ . Bottom: variation of Stokes  $I$ ,  $Q$ ,  $U$  and  $V$  profiles for HD 118022 in the 5018 Fe II line (scaled by a factor of 10 for Stokes  $Q$  and  $U$ , and 3 for Stokes  $V$ ).



**Figure 31.** Longitudinal field measurements for HD 118022 obtained with ESPADONs/NARVAL (shown by filled circles) compared with those obtained by Wade et al. (2000b) with MuSiCoS (shown by filled diamonds). A first-order Fourier fit to the ESPADONs/NARVAL data is given by the solid curve.



**Figure 32.** Net linear polarization (Stokes  $Q$  and  $U$ ) measurements for HD 118022 obtained with ESPADONs/NARVAL (shown by filled circles; scaled by 0.10 and consistent with Wade et al. 2000b) compared with those obtained by Wade et al. (2000b) with MuSiCoS (shown by filled diamonds) and the broad-band linear polarization measurements by Leroy et al. (1995, shown by open squares). Good agreement can be seen, confirming consistency between the instruments. The improvement in data quality is evidenced by the smaller error bars associated with the ESPADONs/NARVAL measurements.

Surprisingly, we found that even when the data are of such high S/N and when the magnetic fields are strong, the LSD analysis is sensitive to the normalization and the measured magnetic field is rather sensitive to the integration ranges chosen, with variations of sometimes of the order of 100 G with very small changes in the integration range. A key conclusion of this work is that even with such high-quality data, extreme care must still be taken with all stages of analysis to ensure consistent results at this level of precision.

Although cross-talk was originally a concern, through a series of experiments we have shown that it is at a level where it should not have a significant impact on the results. By using observations of  $\gamma$  Equ, we have seen that the highest level of cross-talk in Stokes  $Q$  is still within the noise and slightly above the noise in Stokes  $U$  (around the 5 per cent level). It should be noted that even at these levels, this effect will be less significant in the broader lined stars studied here.

Considering this, we believe that the other uncertainties associated with the analysis techniques have a greater effect: normal-

ization, blending, line masks used for LSD and the choice of integration ranges used for longitudinal field measurements. But as regards the final impact of the cross-talk on MDI mapping, this will be discussed and addressed in a future paper (Silvester et al., in preparation). With these high-quality observations, we suspect that the limitations for mapping will in fact not come from the data (with strong Stokes  $Q$  and  $U$  signatures seen in many individual lines) but the ability to deal with line blends within the MDI code.

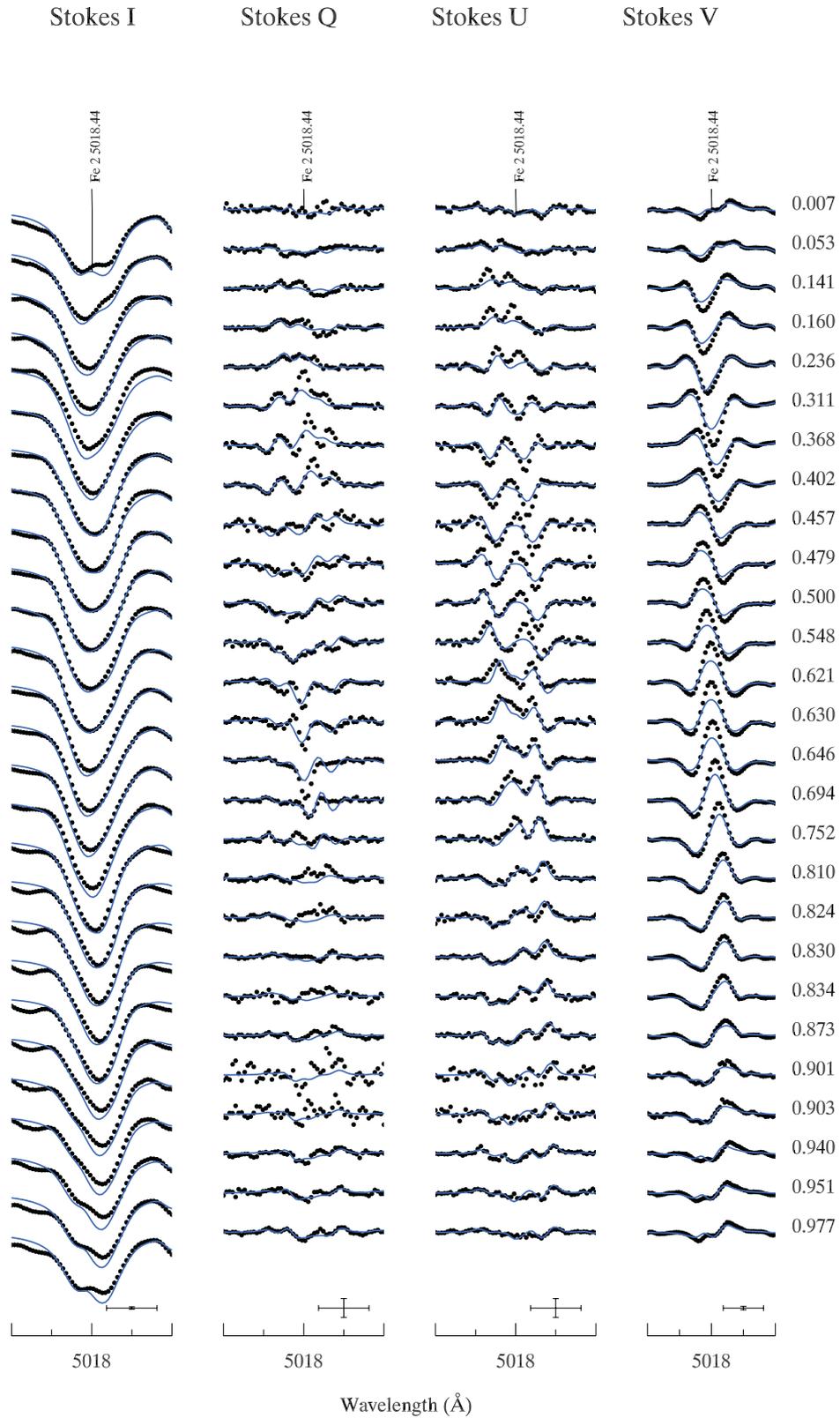
An important result of this study is the confirmed stability of the global properties of the magnetic fields of these Ap stars. Over multiple epochs of observations the fields have remained constant, with little variation in both longitudinal field and linear polarization measurements. In some cases, measurements separated by over a decade still agree with each other within the uncertainties. By comparing MDI maps produced from the new observations of  $\alpha^2$  CVn with those of Kochukhov & Wade (2010), we can potentially test for evolution of the field geometry which may occur on small spatial scales.

Considering the longitudinal magnetic field measurements, agreement was found between the new measurements and those of Wade et al. (2000b), with the exception of HD 71866 and  $\alpha^2$  CVn which showed a slight discrepancy. By re-reducing the MuSiCoS data with the new masks, the observations were brought into agreement. Whilst the general shapes of the net linear polarization variations were in agreement, one needed to invoke free parameters such as scaling which is consistent with what was described in Wade et al. (2000b).

For any study which requires observations over multiple semesters, it is imperative that the instrument is stable and consistent throughout the campaign. Both NARVAL and ESPADONs proved to be very stable instruments, with resolution and S/N being constant over the four years of data. Indeed, we have also shown that the two instruments are consistent with one another with close to identical result from similar phases. These facts demonstrate that ESPADONs and NARVAL are both very capable instruments, well suited to high-resolution four Stokes measurements of magnetic stars over multiple year time-scales.

One of the targets ( $\alpha^2$  CVn) has already been mapped with MDI using MuSiCoS data by Kochukhov & Wade (2010). HD 112413 is an ideal star for determining how much improvement the new polarimetric data could give to MDI mapping. To quantify this improvement and to further confirm consistency, the new observations of  $\alpha^2$  CVn were compared with the profiles predicted by the model of Kochukhov & Wade (2010). As shown in Fig. 33, good general agreement between the new observations and the MuSiCoS-derived model is observed. However, the new profiles show more complexity than was present in the MuSiCoS data, which is not fully reproduced in the current model and would likely require a more complex magnetic field distribution. In addition, the Stokes  $V$  profile amplitude is significantly underestimated by the model at a number of phases.

The complete phase coverage of 49 Cam (Silvester et al., in preparation),  $\alpha^2$  CVn and HD 32633 will allow the completion of four Stokes parameter MDI maps for these stars, doubling the number of Ap stars studied using this technique. Out of the remaining targets HD 4778, HD 71866 and HD 118022 would also be worthwhile candidates for MDI; conversely, HD 40312 has small linear polarization signatures, relative to the noise in their spectra, making them less suitable for MDI analysis. The mapping of 49 Cam is well underway and the results will be presented in Paper II (Silvester et al., in preparation).



**Figure 33.** Comparison between the new ESPaDOnS/NARVAL data for the Fe II 5018 line (shown by black points) with the final model profiles adopted in the mapping of HD 112413 (by solid blue lines) by Kochukhov & Wade (2010)

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